



Limits on Spin-Dependent WIMP-Nucleon Cross Sections from 225 Live Days of XENON100 Data

E. Aprile,¹ M. Alfonsi,² K. Arisaka,³ F. Arneodo,⁴ C. Balan,⁵ L. Baudis,^{6,*} B. Bauermeister,⁷ A. Behrens,⁶ P. Beltrame,^{8,3} K. Bokeloh,⁹ A. Brown,¹⁰ E. Brown,⁹ G. Bruno,⁴ R. Budnik,¹ J. M. R. Cardoso,⁵ W.-T. Chen,¹¹ B. Choi,¹ A. P. Colijn,² H. Contreras,¹ J. P. Cussonneau,¹¹ M. P. Decowski,² E. Duchovni,⁸ S. Fattori,⁷ A. D. Ferella,^{4,6} W. Fulgione,¹² F. Gao,¹³ M. Garbini,¹⁴ C. Ghag,³ K.-L. Giboni,¹ L. W. Goetzke,¹ C. Grignon,⁷ E. Gross,⁸ W. Hampel,¹⁵ F. Kaether,¹⁵ A. Kish,⁶ J. Lamblin,¹¹ H. Landsman,^{8,†} R. F. Lang,¹⁰ M. Le Calloch,¹¹ D. Lellouch,⁸ C. Levy,⁹ K. E. Lim,¹ Q. Lin,¹³ S. Lindemann,¹⁵ M. Lindner,¹⁵ J. A. M. Lopes,⁵ K. Lung,³ T. Marrodán Undagoitia,^{15,6} F. V. Massoli,¹⁴ A. J. Melgarejo Fernandez,¹ Y. Meng,³ M. Messina,¹ A. Molinaro,¹² K. Ni,¹³ U. Oberlack,⁷ S. E. A. Orrigo,⁵ E. Pantic,³ R. Persiani,¹⁴ G. Plante,¹ N. Priel,⁸ A. Rizzo,¹ S. Rosendahl,⁹ J. M. F. dos Santos,⁵ G. Sartorelli,¹⁴ J. Schreiner,¹⁵ M. Schumann,^{16,6} L. Scotto Lavina,¹¹ P. R. Scovell,³ M. Selvi,¹⁴ P. Shagin,¹⁷ H. Simgen,¹⁵ A. Teymourian,³ D. Thers,¹¹ O. Vitells,⁸ H. Wang,³ M. Weber,¹⁵ and C. Weinheimer⁹

(XENON100 Collaboration)

¹Physics Department, Columbia University, New York, New York 10027, USA

²Nikhef and the University of Amsterdam, Science Park, Amsterdam, Netherlands

³Physics and Astronomy Department, University of California, Los Angeles, California 90095, USA

⁴INFN, Laboratori Nazionali del Gran Sasso, Assergi, 67100, Italy

⁵Department of Physics, University of Coimbra, R. Larga, 3004-516, Coimbra, Portugal

⁶Physics Institute, University of Zürich, Winterthurerstrasse, 190 8057 Zurich, Switzerland

⁷Institut für Physik, Johannes Gutenberg-Universität Mainz, 55099 Mainz, Germany

⁸Department of Particle Physics and Astrophysics, Weizmann Institute of Science, 76100 Rehovot, Israel

⁹Institut für Kernphysik, Wilhelms-Universität Münster, 48149 Münster, Germany

¹⁰Department of Physics, Purdue University, West Lafayette, Indiana 47907, USA

¹¹SUBATECH, Ecole des Mines de Nantes, CNRS/In2p3, Université de Nantes, 44307 Nantes, France

¹²INFN-Torino and Osservatorio Astrofisico di Torino, 10100 Torino, Italy

¹³Department of Physics, Shanghai Jiao Tong University, Shanghai 200240, China

¹⁴University of Bologna and INFN-Bologna, Bologna 40126, Italy

¹⁵Max-Planck-Institut für Kernphysik, Saupfercheckweg 1, 69117 Heidelberg, Germany

¹⁶Albert Einstein Center for Fundamental Physics, University of Bern, Sidlerstrasse 5, 3012 Bern, Switzerland

¹⁷Department of Physics and Astronomy, Rice University, Houston, Texas 77005 - 1892, USA

(Received 13 February 2013; revised manuscript received 5 May 2013; published 9 July 2013)

We present new experimental constraints on the elastic, spin-dependent WIMP-nucleon cross section using recent data from the XENON100 experiment, operated in the Laboratori Nazionali del Gran Sasso in Italy. An analysis of 224.6 live days \times 34 kg of exposure acquired during 2011 and 2012 revealed no excess signal due to axial-vector WIMP interactions with ^{129}Xe and ^{131}Xe nuclei. This leads to the most stringent upper limits on WIMP-neutron cross sections for WIMP masses above 6 GeV/ c^2 , with a minimum cross section of 3.5×10^{-40} cm² at a WIMP mass of 45 GeV/ c^2 , at 90% confidence level.

DOI: [10.1103/PhysRevLett.111.021301](https://doi.org/10.1103/PhysRevLett.111.021301)

PACS numbers: 95.35.+d, 14.80.Ly, 29.40.-n

XENON100 was built to search for hypothetical, weakly interacting massive particles (WIMPs), which could explain the nonbaryonic, cold dark matter in our Universe [1]. Independently of astrophysical and cosmological observations, WIMPs are a consequence of many extensions of the standard model of particle physics, as new, stable, or long-lived neutral particles. The WIMP dark matter hypothesis is testable by experiment, the most compelling avenue is to directly observe WIMPs scattering off atomic nuclei in ultralow background terrestrial detectors [2,3]. XENON100 is a double-phase xenon time projection

chamber operated at the Laboratori Nazionali del Gran Sasso (LNGS) in Italy. A total of 178 low-radioactivity, UV-sensitive photomultiplier tubes detect the prompt ($S1$) and proportional ($S2$) light signals induced by particles interacting in the sensitive volume, containing 62 kg of ultrapure liquid xenon. The background level in the energy region of interest for dark matter searches (< 50 keV_{nr}) is 5.3×10^{-3} events kg⁻¹d⁻¹keV⁻¹, before discrimination of electronic and nuclear recoils based on their $S2/S1$ -ratio [4,5]. The instrument is described in [6], the analysis procedure is detailed in [7].

WIMPs in the halo of our Galaxy are expected to be highly nonrelativistic and their interactions with nuclei can be characterized in terms of scalar (or spin-independent, SI) and axial-vector (or spin-dependent, SD) couplings [1,2]. In the case of SI interactions, the leading contribution of the scattering is coherent across the nucleus, and roughly scales with A^2 , where A is the number of nucleons. Our SI result was presented in [4] and excludes a WIMP-nucleon cross section above $2 \times 10^{-45} \text{ cm}^2$ at a WIMP mass of 55 GeV/ c^2 at 90% confidence level. Here we use the same data set, with an exposure of 224.6 live days, a fiducial mass of 34 kg, identical event selection cuts, acceptances, relative scintillation efficiency, and background model to derive limits on spin-dependent interactions.

If the WIMP is a spin-1/2 or a spin-1 field, the contributions to the WIMP-nucleus scattering cross section arise from couplings of the WIMP field to the quark axial current. In the case of the lightest neutralino in supersymmetric models for instance, scattering occurs through the exchange of Z bosons or squarks [1]. To predict actual rates, these fundamental interactions are first translated into interactions with nucleons by evaluating the matrix element of the quark axial-vector current in a nucleon. Finally, the spin components of the nucleons must be added coherently using nuclear wave functions to yield the matrix element for the SD WIMP-nucleus cross section as a function of momentum transfer. The SD differential WIMP-nucleus cross section as a function of momentum transfer q can be written as [8]:

$$\frac{d\sigma_{\text{SD}}(q)}{dq^2} = \frac{8G_F^2}{(2J+1)v^2} S_A(q), \quad (1)$$

where G_F is the Fermi constant, v is the WIMP speed relative to the target, J is the total angular momentum of the nucleus and S_A is the axial-vector structure function. In the limit of zero momentum transfer (at finite momentum transfer, or when WIMP couplings to two nucleons are included [9], the neutron-only coupling case implies also coupling to protons and vice versa) the structure function reduces to the form [10]:

$$S_A(0) = \frac{(2J+1)(J+1)}{\pi J} [a_p \langle S_p \rangle + a_n \langle S_n \rangle]^2, \quad (2)$$

where $\langle S_{p,n} \rangle = \langle J | \hat{S}_{p,n} | J \rangle$ are the expectation values of the total proton and neutron spin operators in the nucleus, and the effective WIMP couplings to protons and neutrons are defined in terms of the isoscalar $a_0 = a_p + a_n$ and isovector $a_1 = a_p - a_n$ couplings.

WIMPs will thus couple to the total angular momentum of a nucleus and only nuclei with an odd number of protons or/and neutrons will yield a significant sensitivity to this channel. Natural xenon contains two nonzero spin isotopes, ^{129}Xe (spin-1/2) and ^{131}Xe (spin-3/2), with an abundance of 26.4% and 21.2%, respectively. In XENON100, the isotopic abundances of ^{129}Xe and ^{131}Xe

are changed to 26.2% and 21.8%, respectively, due to the addition of isotopically modified xenon to the available natural xenon.

To compare results from different target materials, a common practice is to report the cross section for the interaction with a single nucleon (σ_p , σ_n) [11–13]. Assuming that WIMPs couple predominantly to protons ($a_n = 0$) or neutrons ($a_p = 0$), the WIMP-nucleon cross section becomes

$$\sigma_{p,n}(q) = \frac{3}{4} \frac{\mu_{p,n}^2}{\mu_A^2} \frac{2J+1}{\pi} \frac{\sigma_{\text{SD}}(q)}{S_A^{a_0=\pm a_1}(q)}, \quad (3)$$

where σ_{SD} is the total WIMP-nucleus cross section, μ_A and $\mu_{p,n}$ are the WIMP-nucleus and WIMP-nucleon reduced masses, respectively.

Calculations of the structure functions $S_A(q)$ are traditionally based on the nuclear shell model, but differ in the effective nucleon-nucleon interactions and in the valence space and truncation used for the computation. For xenon as a WIMP target material, we consider three large-scale shell-model calculations: by Ressel and Dean [14] with the Bonn-A [15] two-nucleon potential, by Toivanen *et al.* [16], using the CD-Bonn potential [15], and the recent results by Menendez *et al.* [9], using state-of-the-art valence shell interactions [17,18] and less severe truncations of the valence space. Menendez *et al.* [9] also use for the first time chiral effective field theory (EFT) currents [19] to determine the couplings of WIMPs to nucleons. The currents for spin-dependent scattering are derived at the one-body level and the leading long-range two-nucleon currents are included, resulting in a reduction of the isovector part of the one-body axial-vector WIMP currents [9]. The resulting chiral EFT currents are then used to calculate the structure functions for the WIMP-xenon scattering. Theoretical errors due to nuclear uncertainties can be provided when chiral two-body currents are included [9]; we show their effect on our limits in this Letter. The shell-model calculations are based on the largest many-body spaces accessible with nuclear interactions, also used to calculate double-beta decay matrix elements for nuclei up to ^{136}Xe and to study nuclear structures [17,18,20].

The new calculations by Menendez *et al.* [9] yield a far superior agreement between calculated and measured spectra of the ^{129}Xe and ^{131}Xe nuclei, both in energy and in the ordering of the nuclear levels, compared to older [16] results. The values for $\langle S_{p,n} \rangle$ are close to those of Ressel and Dean [14], but quite different from the results of Toivanen *et al.* [16], as summarized in Table I. We thus use the Menendez *et al.* [9] structure functions for our benchmark upper limits on WIMP-neutron and WIMP-proton cross sections. We also provide a comparison to the limits obtained when using the calculations by Ressel and Dean [14] and Toivanen *et al.* [16]. In all cases, $|\langle S_n \rangle| \gg |\langle S_p \rangle|$, as expected for the two xenon nuclei

TABLE I. Parameters of the xenon isotopes used in this analysis: nuclear total angular momentum and parity of the ground state, J^P , predicted expectation values of the total proton and neutron spin operators in the nucleus ($\langle S_{n,p} \rangle$) by the Ressel and Dean (Bonn A potential) [14], Toivanen *et al.* (Bonn CD potential) [16], and Menendez *et al.* (state-of-the art valence shell interactions) [9] calculations.

Nucleus	J^P	Ressel and Dean [14]		Toivanen <i>et al.</i> [16]		Menendez <i>et al.</i> [9]	
		$\langle S_n \rangle$	$\langle S_p \rangle$	$\langle S_n \rangle$	$\langle S_p \rangle$	$\langle S_n \rangle$	$\langle S_p \rangle$
^{129}Xe	$(1/2)_{g.s.}^+$	0.359	0.028	0.273	-0.0019	0.329	0.010
^{131}Xe	$(3/2)_{g.s.}^+$	-0.227	-0.009	-0.125	-0.00069	-0.272	-0.009

with an odd number of neutrons and an even number of protons.

Figure 1 shows the structure functions $S_A(q)$ obtained from the three calculations for pure neutron and pure proton couplings as a function of nuclear recoil energy. For the neutron coupling case, for which xenon has the best sensitivity, the functions are rather similar. For the proton coupling case, the structure function by Toivanen *et al.* [16] differs significantly from the other two results. We note that, for xenon, a significant effect in the proton channel had already been pointed out in [21], in a comparison between the results of Ressel and Dean with the Bonn-A potential, and Toivanen *et al.* using the Bonn-CD nucleon-nucleon potential.

Table I summarizes the expectation values of the total proton and neutron spin operators in the nucleus for ^{129}Xe and ^{131}Xe in the zero momentum transfer limit.

Constraints on the spin-dependent WIMP-nucleon cross sections are calculated using the Profile Likelihood

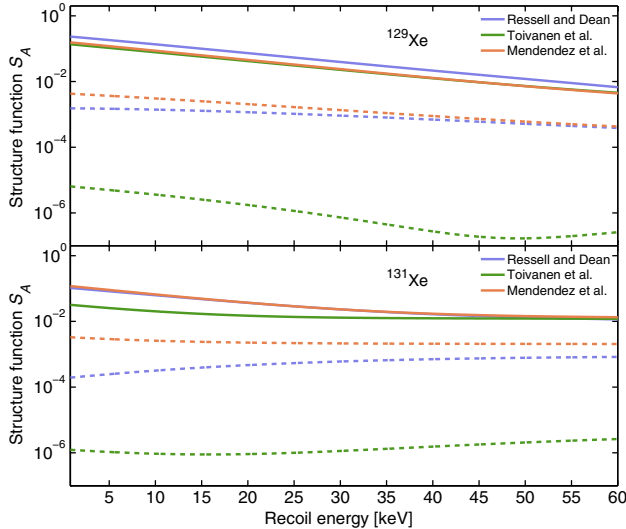


FIG. 1 (color online). Structure functions for ^{129}Xe (top) and ^{131}Xe (bottom) for the case of neutron (plain) and proton (dashed) couplings, as a function of recoil energy using the calculations of Ressel and Dean [14], Toivanen *et al.* [16], and Menendez *et al.* [9]. The difference is most significant in the case of the proton coupling for the Toivanen *et al.* results.

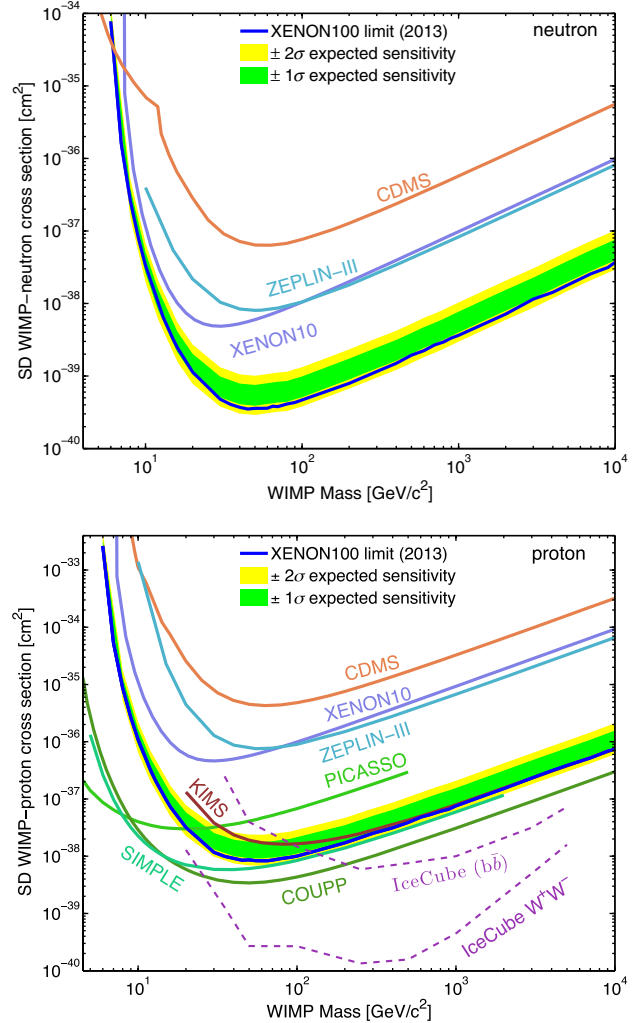


FIG. 2 (color online). XENON100 90% C.L. upper limits on the WIMP SD cross section on neutrons (top) and protons (bottom) using Menendez *et al.* [9]. The 1σ (2σ) uncertainty band on the expected sensitivity of this run is shown as a green (yellow) band. Also shown are results from XENON10 [24] (using Ressel and Dean [14]), CDMS [25,26], ZEPLIN-III [21,27] (using Toivanen *et al.* [16] and Ressel and Dean [14] for the neutron and proton case, respectively), PICASSO [28], COUPP [29], SIMPLE [30], KIMS [31], IceCube [32] in the hard ($W^+ W^-$, $\tau^+ \tau^-$ for WIMP masses $< 80.4 \text{ GeV}/c^2$), and soft ($b\bar{b}$) annihilation channels.

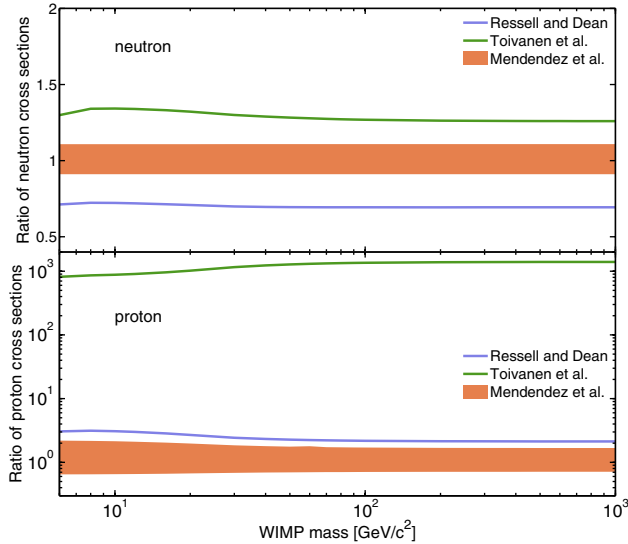


FIG. 3 (color online). The ratio of upper limits calculated with the Ressel and Dean [14] and the Toivanen *et al.* [16] results for the structure functions to the ones obtained using Menendez *et al.* [9] for the case of neutron (top) and proton (bottom) couplings, along with the theoretical uncertainty band due to chiral two-body currents [9].

approach described in [22]. Systematic uncertainties in the energy scale and in the background expectation are taken into account when constructing the Profile Likelihood model and are reflected in the actual limit. It is given at 90% C.L. after taking into account statistical downward fluctuations in the background. We assume that the dark matter is distributed in an isothermal halo with a truncated Maxwellian velocity distribution with a local circular speed of $v_c = 220$ km/s, galactic escape velocity $v_{\text{esc}} = 544$ km/s and a local density of $\rho = 0.3$ GeV cm $^{-3}$ [23].

The resulting upper limits from XENON100, along with results from other experiments, are shown in Fig. 2 for neutron couplings (top panel) and proton couplings (lower panel). The 1σ (2σ) uncertainty on the sensitivity of this run, namely the expected limit in absence of a signal above the background, is shown as a green (yellow) band in Fig. 2. The impact on these limits when using the Toivanen *et al.* and the Ressel and Dean calculations are shown in Fig. 3. XENON100 provides the most stringent limits for pure neutron couplings for WIMP masses above 6 GeV/ c^2 , excluding previously unexplored regions in the allowed parameter space. The minimum WIMP-neutron cross section is 3.5×10^{-40} cm 2 at a WIMP mass of 45 GeV/ c^2 , using Menendez *et al.* [9]. It changes to 2.5×10^{-40} cm 2 and 4.5×10^{-40} cm 2 when using Ressel and Dean and Toivanen *et al.*, respectively. The sensitivity to proton couplings (Fig. 2, bottom panel) is much weaker because, as detailed above, both ^{129}Xe and ^{131}Xe have an unpaired neutron but an even number of protons, thus $|\langle S_p \rangle| \ll |\langle S_n \rangle|$ (see Table I). Upper limits from other direct and indirect detection experiments are shown for comparison.

In conclusion, we have analyzed data from 224.6 live days \times 34 kg exposure acquired by XENON100 during 13 months of operation in 2011/2012 for SD WIMP interactions. We saw no evidence for a dark matter signal and have obtained new experimental upper limits on the spin-dependent WIMP-nucleon cross section. For our limits, we use the new calculations by Menendez *et al.* [9], where the WIMP couplings to nucleons are derived using chiral EFT currents and which yield a good agreement between the calculated and measured energy spectra of the ^{129}Xe and ^{131}Xe nuclei. We note that the interpretation of the results in terms of SD pure-proton cross section strongly depends on the used nuclear model. However, regardless of the nuclear model, we obtain the most stringent limits to date on spin-dependent WIMP-neutron couplings for WIMP masses above 6 GeV/ c^2 at 90% C.L.

We gratefully acknowledge support from NSF, DOE, SNF, UZH, FCT, INFN, Région des Pays de la Loire, STCSM, NSFC, DFG, Stichting voor Fundamenteel Onderzoek der Materie (FOM), the Max Planck Society, the Weizmann Institute of Science, and the EMG research center. We thank Achim Schwenk and Javier Menendez for many helpful discussions and for providing their structure functions in numerical form. We thank Jouni Suhonen for providing us the numerical data for Fig. 3 and Michael Pitt (WIS) for his contribution. We are grateful to LNGS for hosting and supporting XENON100.

*laura.baudis@physik.uzh.ch

†hagar.landsman@weizmann.ac.il

- [1] G. Jungman, M. Kamionkowski, and K. Griest, *Phys. Rep.* **267**, 195 (1996).
- [2] M. W. Goodman and E. Witten, *Phys. Rev. D* **31**, 3059 (1985).
- [3] L. Baudis, *Phys. Dark Univ.* **1**, 94 (2012).
- [4] E. Aprile *et al.* (XENON100 Collaboration), *Phys. Rev. Lett.* **109**, 181301 (2012).
- [5] E. Aprile *et al.* (XENON100 Collaboration), *Phys. Rev. D* **83**, 082001 (2011).
- [6] E. Aprile *et al.* (XENON100 Collaboration), *Astropart. Phys.* **35**, 573 (2012).
- [7] E. Aprile *et al.* (XENON100 Collaboration), [arXiv:1207.3458](https://arxiv.org/abs/1207.3458).
- [8] J. Engel, S. Pittel, and P. Vogel, *Int. J. Mod. Phys. E* **01**, 1 (1992).
- [9] J. Menendez, D. Gazit, and A. Schwenk, *Phys. Rev. D* **86**, 103511 (2012).
- [10] M. T. Ressel, M. Aufderheide, S. Bloom, K. Griest, G. Mathews, and D. Resler, *Phys. Rev. D* **48**, 5519 (1993).
- [11] D. Tovey, R. J. Gaitskell, P. Gondolo, Y. Ramachers, and L. Roszkowski, *Phys. Lett. B* **488**, 17 (2000).
- [12] F. Giuliani, *Phys. Rev. Lett.* **93**, 161301 (2004).
- [13] C. Savage, P. Gondolo, and K. Freese, *Phys. Rev. D* **70**, 123513 (2004).
- [14] M. T. Ressel and D. J. Dean, *Phys. Rev. C* **56**, 535 (1997).

- [15] M. Hjorth-Jensen, T. Kuo, and E. Osnes, *Phys. Rep.* **261**, 125 (1995).
- [16] P. Toivanen, M. Kortelainen, J. Suhonen, and J. Toivanen, *Phys. Rev. C* **79**, 044302 (2009).
- [17] E. Caurier, J. Menendez, F. Nowacki, and A. Poves, *Phys. Rev. Lett.* **100**, 052503 (2008).
- [18] J. Menendez, A. Poves, E. Caurier, and F. Nowacki, *Nucl. Phys.* **A818**, 139 (2009).
- [19] T. Park, L. Marcucci, R. Schiavilla, M. Viviani, A. Kievsky, S. Rosati, K. Kubodera, D.-P. Min, and M. Rho, *Phys. Rev. C* **67**, 055206 (2003).
- [20] J. Menendez, D. Gazit, and A. Schwenk, *Phys. Rev. Lett.* **107**, 062501 (2011).
- [21] V. Lebedenko *et al.* (ZEPLIN-III Collaboration), *Phys. Rev. Lett.* **103**, 151302 (2009).
- [22] E. Aprile *et al.* (XENON100 Collaboration), *Phys. Rev. D* **84**, 052003 (2011).
- [23] A.M. Green, *Mod. Phys. Lett. A* **27**, 1230004 (2012).
- [24] J. Angle *et al.* (XENON10 Collaboration), *Phys. Rev. Lett.* **101**, 091301 (2008).
- [25] Z. Ahmed *et al.* (CDMS Collaboration), *Phys. Rev. Lett.* **102**, 011301 (2009).
- [26] Z. Ahmed *et al.* (CDMS Collaboration), *Phys. Rev. Lett.* **106**, 131302 (2011).
- [27] D. Akimov *et al.* (ZEPLIN-III Collaboration), *Phys. Lett. B* **709**, 14 (2012).
- [28] S. Archambault *et al.* (PICASSO Collaboration), *Phys. Lett. B* **711**, 153 (2012).
- [29] E. Behnke *et al.* (COUPP Collaboration), *Phys. Rev. D* **86**, 052001 (2012).
- [30] M. Felizardo *et al.* (SIMPLE Collaboration), *Phys. Rev. Lett.* **108**, 201302 (2012).
- [31] S. Kim *et al.* (KIMS Collaboration), *Phys. Rev. Lett.* **108**, 181301 (2012).
- [32] M.G. Aartsen *et al.* (IceCube Collaboration), *Phys. Rev. Lett.* **110**, 131302 (2013).