New Features of Electron Phase Space Holes Observed by the THEMIS Mission

L. Andersson,[1] R. E. Ergun,[1,2] J. Tao,[1,2] A. Roux,[3] O. LeContel,[3] V. Angelopoulos,[4] J. Bonnell,[5] J. P. McFadden,[5] D. E. Larson,[5] S. Eriksson,[1] T. Johansson,[1] C. M. Cully,[6] D. L. Newman,[7] M. V. Goldman,[7] K.-H. Glassmeier,[8] and W. Baumjohann[9]

[1]*Laboratory for Atmospheric and Space Physics, University of Colorado, Boulder, Colorado 80309, USA*
[2]*Department of Astrophysical and Planetary Sciences, University of Colorado, Boulder, Colorado 80309, USA*
[3]*Centre d'étude des Environnements Terrestre et Planétaires, Velizy, France*
[4]*Institute of Geophysics and Planetary Physics, University of California, Los Angeles, California 90055, USA*
[5]*Space Sciences Laboratory, University of California, Berkeley, California, 94720, USA*
[6]*Swedish Institute of Space Physics, Uppsala, Sweden*
[7]*Center for Integrated Plasma Studies, University of Colorado, Boulder, Colorado 80309, USA*
[8]*TUBS, Braunschweig, D-38106, Germany*
[9]*Space Research Institute, Austrian Academy of Sciences, A-8042 Graz, Austria*

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Observations of electron phase-space holes (EHs) in Earth's plasma sheet by the THEMIS satellites include the first detection of a magnetic perturbation ($\delta B_\parallel$) parallel to the ambient magnetic field ($\boldsymbol{B}_0$). EHs with a detectable $\delta B_\parallel$ have several distinguishing features including large electric field amplitudes, a magnetic perturbation perpendicular to $\boldsymbol{B}_0$, high speeds ($\sim 0.3c$) along $\boldsymbol{B}_0$, and sizes along $\boldsymbol{B}_0$ of tens of Debye lengths. These EHs have a significant center potential ($\Phi \sim k_B T_e/e$), suggesting strongly nonlinear behavior nearby such as double layers or magnetic reconnection.

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Electron phase-space holes (EHs) [1–3] are ubiquitous in space plasmas. Observations have been made in the plasma sheet [4], auroral zone [5–7], magnetosheath [8], magnetopause [9], bow shock transition region [10], and solar wind [11]. EHs can be described as BGK [12] structures [3,13] or modeled by small potential expansion [1,14]. Theoretical treatments on generation concentrate on the electron two-stream instability [15,16] and Bunemann instability [17]. Most importantly, EHs are associated with processes such as double layers [18–20] and magnetic reconnection [17,21,22] making EHs a reliable indicator of strongly nonlinear behavior in plasmas.

The observational characteristics of EHs have been reported in a number of articles [7,8,23]. They are detected as bipolar electric field signals ($\delta E_\parallel$) parallel to $\boldsymbol{B}_0$ [4–6]. The parallel scale sizes ($L_\parallel$, defined here as the distance between peaks in $\delta E_\parallel$) are most often several electron Debye lengths (λ_D) and the speeds (v_{EH}) are near to, but often less than, the electron thermal speed (v_e). The perpendicular scale sizes ($L_\perp$) are comparable to $L_\parallel$ in the low-altitude auroral region [23], whereas it has been reported that $L_\perp \gg L_\parallel$ in most other space environments [7,8]. EHs are most often weak ($e\Phi_0/k_B T_e \ll 1$, where Φ_0 is the center potential, e is the electron charge, and T_e is the electron temperature). THEMIS observations largely support these earlier results.

Space-based measurements record a profile in time, so the derivation of Φ and $L_\parallel$ depends on the speed of the EH. The statistical characteristics described above relate to "slow-moving" EHs. By "slow-moving", we mean that the speeds of the EHs are derived from the time delay between the signals of two spatially separated electric field probes [7,23]. Most instruments are limited to measuring $v_{\mathrm{EH}} < \sim 1000$ km/s with this technique.

In this article, we present the first 3D observations of magnetic field perturbations caused by EHs including the detection of a $\delta B_\parallel$ signal. We show that the perpendicular magnetic perturbation ($\delta B_\perp$) is primarily caused by the motion of a quasielectrostatic EH. In other words, $\delta B_\perp$ is consistent with the Lorentz transformation of $\delta E_\perp$ [5,23]. If EHs are quasielectrostatic in their rest frame (see later discussion on the "rest" frame), $\delta B_\perp$ and $\delta E_\perp$ can be used to accurately determine their speed, particularly if they are "fast-moving" (>1000 km/s) and, subsequently, accurately derive Φ and $L_\parallel$. We also show that EHs with a detectable $\delta B_\parallel$ have quite different characteristics than reported by earlier observations. They have large electric field amplitudes, $\delta \boldsymbol{E} \sim O(100\ \mathrm{mV/m})$, high speeds ($v_{\mathrm{EH}} > v_e$), large parallel sizes ($L_\parallel > 10\lambda_D$), moderate to strong center potentials ($e\Phi/k_B T_e \sim 0.5$), and elongated shapes ($L_\parallel > L_\perp$). We suggest that $\delta B_\parallel$ arises from the $\delta \boldsymbol{E} \times \boldsymbol{B}_0$ electron motion in the EH and that $\delta B_\parallel \propto \Phi$. These observations have a number of similarities to laboratory observations of elongated, high-speed holes associated with magnetic reconnection [22].

The observations are from the THEMIS mission [24], which has five identical satellites in highly eccentric orbits at low inclination with apogees that range from $10R_E$ to

$30R_E$. The satellites carry electron and ion analyzers [25], a three-axis electric field instrument (dc—8 kHz) [26], a dc magnetometer [27], and a search coil magnetometer [28].

Figure 1 presents five minutes of observations from THEMIS Probe A at $\sim 10R_E$ from Earth's center. The top panel [Fig. 1(a)] displays a spectrogram of the electron differential energy flux as a function of energy (vertical axis) and time (horizontal axis). The energy flux is calculated from a set of two-dimensional, energy-angle measurements averaged over a satellite spin period ($\sim$3 s) [25]. Data from two detectors are combined. The lower-energy electrons ($\sim$10 eV to $\sim$30 keV) are measured by an electrostatic analyzer whereas the higher-energy electrons are detected by a solid state telescope. The lowest-energy electron fluxes ($<$25 eV) are spacecraft photoelectrons. The black trace overlying the spectrogram is T_e in eV based on the electrostatic analyzer data.

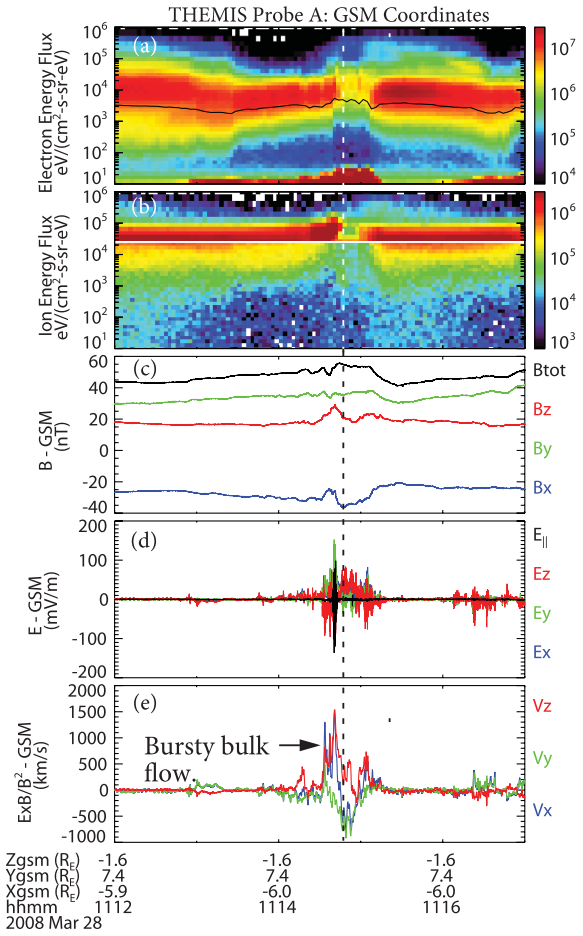

FIG. 1 (color). (a) Electron differential energy flux as a function of energy (vertical axis) and time (horizontal axis). The black trace is T_e. (b) Ion differential energy flux. (c) Magnetic field in GSM coordinates at 128 samples/s. The black trace is $|\boldsymbol{B}_0|$. (d) $\boldsymbol{E}_0$ in GSM coordinates at 128 samples/s. The black trace is $E_{0\|}$. (e) $\delta\boldsymbol{E} \times \boldsymbol{B}_0/|\boldsymbol{B}_0|^2$ low-pass filtered to 1 Hz in GSM coordinates. The vertical dashed line marks the period of the EHs in Fig. 2.

Figure 1(b) displays the differential energy flux of ions in the same format. A small gap in energy coverage is seen in the plot as white space. Figure 1(c) plots the dc-50 Hz magnetic field ($\boldsymbol{B}_0$) at 128 samples/s in geocentric solar magnetospheric (GSM) coordinates. The absolute accuracy is better than 1 nT [27]. The color represents direction: blue is towards the Sun, red is near Earth's magnetic north, and green completes the set. The black trace in Fig. 1(c) is $|\boldsymbol{B}_0|$.

The dc-50 Hz electric field [$\boldsymbol{E}_0$, Fig. 1(d)] is measured by three orthogonal, dipole antennas [26]. The black trace in Fig. 1(d) represents the parallel electric field, $E_{0\|}$. The antennas in the spin plane of the spacecraft, mostly covering the GSM x and y directions, have $\sim$40 m and $\sim$50 m physical lengths and are accurate to approximately ± 2 mV/m, depending on plasma conditions. The spin-axis dipole, predominantly the GSM z direction, is $\sim$7 m and is accurate to ± 20 mV/m.

Figure 1(e) plots the quantity $\boldsymbol{E}_0 \times \boldsymbol{B}_0/|\boldsymbol{B}_0|^2$ low-pass filtered to 1 Hz representing the flow perpendicular to $\boldsymbol{B}_0$. The x-component of the flow (towards Earth; blue trace) rises to over 1000 km/s at $\sim$11:14.5 UT indicating a bursty bulk flow event [29]. Such events are associated with magnetic reconnection occurring anti-Earthward of the spacecraft's position. During the bursty bulk flow event, the electron and ion energies increase [Figs. 1(a) and 1(b)] and $\boldsymbol{E}_0$ and $\boldsymbol{B}_0$ display strong variations [Figs. 1(c) and 1(d)].

Figure 2 presents 0.2 seconds of high-time resolution $\delta\boldsymbol{E}$ and $\delta\boldsymbol{B}$ signals (filtered from $\sim$5 Hz to $\sim$3.3 kHz; 8192 samples/s) during the time marked with a vertical dashed line in Fig. 1. The signals are in a magnetic coordinate system such that $\delta E_\|$ is parallel to $\boldsymbol{B}_0$, δE_X (accurate to ± 2 mV/m) is the perpendicular component measured only by the spin-plane booms, and δE_Y (accurate to ± 20 mV/m) completes the vector. The ac magnetic field signals are in the same coordinate system.

The $\delta E_\|$ signal [Fig. 2(a)] shows a series of bipolar structures, a defining signature of EHs [4–6]. All of the EHs have a positive then negative polarity indicating that they are traveling in the same direction and, consequently, are likely to come from the same source. The perpendicular electric field signals [Figs. 2(c) and 2(e)] have a corresponding unipolar perturbation, again, typical of EHs. Some of the EHs are such that δE_X or δE_Y are greater than $\delta E_\|$, a sufficient but not necessary condition for $L_\| \geq L_\perp$ (since the spacecraft may pass through the center of the EH rather than the edge, a small perpendicular signal does not, by itself, reveal the relation between $L_\|$ and $L_\perp$). Almost all of the EHs have a corresponding positive unipolar perturbation in $\delta B_\|$ [Fig. 2(b)].

The perpendicular $\delta\boldsymbol{E}$ and $\delta\boldsymbol{B}$ signals [Figs. 2(c)–2(f)] are arranged in orthogonal pairs [δE_X, Fig. 2(c), is orthogonal to δB_Y, Fig. 2(d), etc.]. δE_X and δB_Y are well correlated and δE_Y and δB_X have a negative correlation,

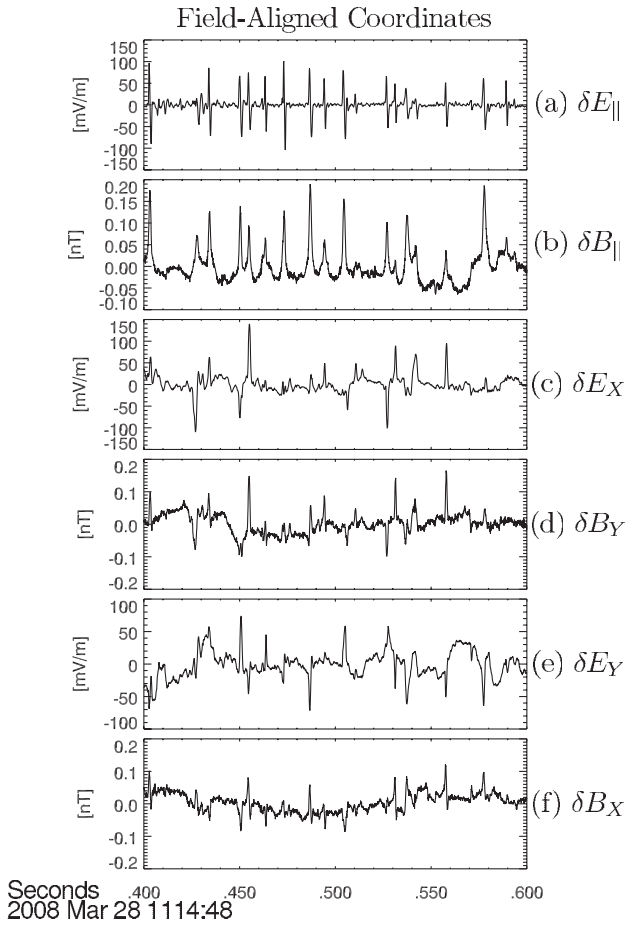

FIG. 2. (a) $\delta E_{\|}$ (5 Hz–3.3 kHz) at 8192 samples/s during the period marked on Fig. 1. (b) $\delta B_{\|}$ (5 Hz–3.3 kHz) at 8192 samples/s. (c) δE_X is from the long wire antennas and accurate to ± 2 mV/m. (d) δB_Y is orthogonal to δE_X. One can see that δE_X and δB_Y signals of EHs are well correlated. (e) δE_Y (± 20 mV/m) is derived from a combination of all electric field dipole antennas including the short (7 m) dipole along the spacecraft spin axis [26]. (f) δB_X.

albeit somewhat weaker. These $\delta \boldsymbol{E}$ and $\delta \boldsymbol{B}$ signals are consistent with a Lorentz transformation of a moving quasielectrostatic structure ([30], changed to SI units):

$$\boldsymbol{B}' = \gamma(\boldsymbol{B} - \boldsymbol{v} \times \boldsymbol{E}/c^2) - \frac{\gamma^2}{1+\gamma}\frac{\boldsymbol{v}(\boldsymbol{v}\cdot\boldsymbol{B})}{c^2}. \quad (1)$$

In their rest frame, the perpendicular $\delta \boldsymbol{B}'$ signals nearly vanish. With v_{EH} parallel to $\boldsymbol{B}_0$, the perpendicular components in Eq. (1) reduce to

$$\delta B_Y = \frac{v_{\mathrm{EH}}}{c^2}\delta E_X, \qquad \delta B_X = -\frac{v_{\mathrm{EH}}}{c^2}\delta E_Y. \quad (2)$$

Most importantly, the data indicate that a quasielectrostatic frame exists. In other words, there is a frame in which the perpendicular $\delta \boldsymbol{B}'$ signals nearly vanish (are minimum). The velocity of this frame, and presumably that of the EH, can be derived from δE_X and δB_Y, the more accurate of the orthogonal pairs.

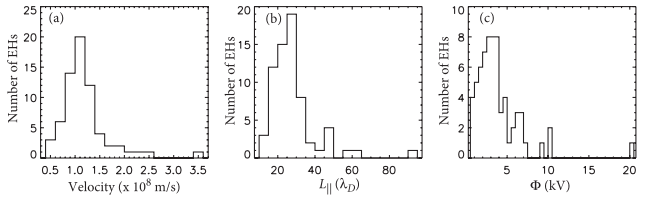

FIG. 3. (a) A histogram of the velocity of 67 EHs observed during a ~16 s wave burst period (11:14:41 UT to 11:14:57 UT on March 28, 2008). The velocity was derived as $c^2\delta B_Y/\delta E_X$. We note that one event has a derived speed greater than c which is either due to the uncertainty in δB_Y and δE_X or a strong electromagnetic (δB_r) contribution. (b) A histogram of the parallel size of the EHs. $L_{\|}$ is the distance between the negative and positive peaks in $\delta E_{\|}$ assuming the EH is traveling at the derived velocity. (c) A histogram of the potential of the EHs derived from the $\delta E_{\|}$ signal. The spacecraft may not have passed through the center of the EH, so Φ represents a lower bound.

Figure 3(a) displays the derived velocity ($c^2\delta B_Y/\delta E_X$) of 67 EHs detected in a ~16 s "wave burst" period (11:14:41 UT to 11:14:57 UT on March 28, 2008) of high-time resolution (8192 samples/s) waveform that includes the data in Fig. 2 (see Refs. [24,26] for discussion on wave burst data collection). The mean velocity of the EHs is $\sim 1 \times 10^8$ m/s. This high speed implies that these EHs are traveling faster than the thermal velocity ($v_e \sim 4 \times 10^7$ m/s). Using the derived velocity, the size of the EHs along $\boldsymbol{B}_0$ is displayed in Fig. 3(b). $L_{\|}$ is roughly 30 λ_D, where $\lambda_D \sim 3.0$ km (derived from a 3 s average electron distribution). The mean value of Φ [Fig. 3(c)] is ~3 keV. Within uncertainties, $T_e \sim 8$ keV (parallel to $\boldsymbol{B}_0$). We cannot determine the radial offset of the measurements (distance perpendicular to $\boldsymbol{B}_0$ from the center of the EH), so Φ represents a lower bound. These EH observations have moderate potentials ($e\Phi/k_BT_e \sim 0.5$) and are unusual in that $v_{\mathrm{EH}} > v_e$, and $L_{\|}$ is tens of λ_D. Similar results were reported from laboratory experiments on magnetic reconnection [22].

The presence of the $\delta B_{\|}$ signal supports the above conclusions. This signal can be explained from the electron $\delta \boldsymbol{E} \times \boldsymbol{B}_0$ currents generated by the perpendicular electric field signal. In the spacecraft frame, the duration of the EHs (~1.2 ms) is about 2 times the electron gyro-period (~0.66 ms), so an electron drift can be established whereas the ion motion is negligible. The resulting perpendicular current loop is around the center of the EH with $J_\phi \cong -en_e\delta E_r/|\boldsymbol{B}_0|$. Here, δE_r represents the radial perpendicular electric field perturbation, and n_e is the ambient electron density. This current will generate a magnetic field in the same direction as $\boldsymbol{B}_0$ in the center of the EH; hence, $\delta B_{\|}$ is always positive. The amplitude of $\delta B_{\|}$ depends on δE_r and the shape of the EH.

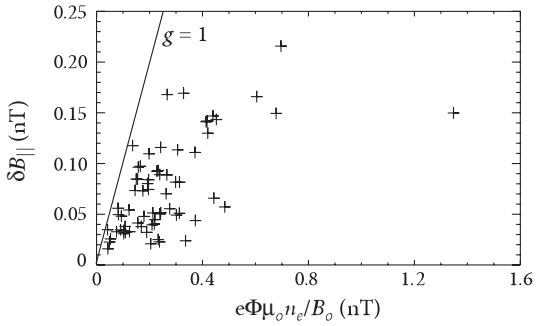

FIG. 4. The maximum value of $\delta B_\parallel$ versus $e\Phi\mu_0 n_e/B_0$ for the 67 EH measured in a 16 s burst period. There is no correction for radial offset, so both $\delta B_\parallel$ versus Φ represent lower bounds. $\delta B_\parallel$ and Φ do not have the same behavior as a function of radial offset, so the data exhibit significant scatter.

By modeling the hole as cylindrically symmetric with a Gaussian shape:

$$\Phi(r, z) = \Phi_0 e^{-r^2/2L_\perp^2} e^{-z^2/2L_\parallel^2}, \tag{3}$$

$\delta B_\parallel$ at the center of the EH can be derived by integrating the Biot-Savart equation:

$$\delta B_\parallel(r = 0, z = 0) = \frac{e\Phi_0\mu_0 n_e}{\boldsymbol{B}_0} g(L_\parallel, L_\perp), \tag{4}$$

where $g(L_\parallel, L_\perp) < 1$, is a dimensionless geometric factor.

Figure 4 presents $\delta B_\parallel$ versus $e\Phi\mu_0 n_e/B_0$ for the 67 EHs measured in the 16 s wave burst period. We do not correct for radial offset ($r \neq 0$), so both $\delta B_\parallel$ and Φ represent lower bounds. $\delta B_\parallel$ and Φ do not have the same behavior as a function of radial offset, so the data exhibit significant scatter. There are, however, two important properties. The values of $\delta B_\parallel$ are nearly equal to but always less than that of $e\Phi\mu_0 n_e/B_0$ consistent with $g(L_\parallel, L_\perp) < 1$ and, furthermore, $\delta B_\parallel$ increases with increasing Φ. These data, along with the observation that $\delta B_\parallel > 0$, support our supposition that $\delta B_\parallel$ results from electron $\delta\boldsymbol{E} \times \boldsymbol{B}_0$ currents. Thus the $\delta B_\parallel$ signal is in consort with the large amplitudes, high speeds, moderate to strong potentials, and elongated shape of the EHs.

Since the $\delta B_\parallel$ exists in all frames, the EHs cannot be entirely electrostatic. If an EH is cylindrically symmetric, then a radial magnetic field must be present, even in the rest frame ($\delta B_r \neq 0$ since $\nabla \cdot \boldsymbol{B} = 0$). The rest frame is best defined as the frame in which the azimuthal magnetic field vanishes. δB_r is due to J_ϕ, so it should be detected by a spacecraft as bipolar signal (the radial magnetic field of a current ring has opposite signs for $z > 0$ and $z < 0$). Careful examination of the measured magnetic field signals in Fig. 2 show that they are predominantly unipolar, so $\delta B_r \ll \delta B_\parallel$. δB_r is expected to be small if $L_\parallel \geq L_\perp$. A small δB_r is consistent with the elongated shape.

In conclusion, we have presented observations of the perturbation magnetic field and the first report of $\delta B_\parallel$ associated electron phase-space holes. These EHs differ from earlier observations in that they have high speeds ($v_{\mathrm{EH}} > v_e$), large parallel sizes ($L_\parallel > 10\lambda_D$), significant center potentials ($e\Phi/k_B T_e \sim 1$), and elongated shapes ($L_\parallel > L_\perp$). In particular, these EHs have many characteristics that are similar to those generated by magnetic reconnection in a laboratory experiment [22]. EHs also are known to be generated by double layers [18–20], so observations of EHs are an indicator of nonlinear, kinetic behavior in the active plasma sheet.

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