

Shape of cosmic string loops

Craig J. Copi*

CERCA, Department of Physics, Case Western Reserve University, Cleveland, Ohio 44106-7079, USA

Tanmay Vachaspati

Physics Department, Arizona State University, Tempe, Arizona 85287, USA

(Received 1 November 2010; published 27 January 2011)

Complicated cosmic string loops will fragment until they reach simple, nonintersecting (“stable”) configurations. Through extensive numerical study we characterize these attractor loop shapes including their length, velocity, kink, and cusp distributions. We find that an initial loop containing M harmonic modes will, on average, split into $3M$ stable loops. These stable loops are approximately described by the degenerate kinky loop, which is planar and rectangular, independently of the number of modes on the initial loop. This is confirmed by an analytic construction of a stable family of perturbed degenerate kinky loops. The average stable loop is also found to have a 40% chance of containing a cusp. We examine the properties of stable loops of different lengths and find only slight variation. Finally we develop a new analytic scheme to explicitly solve the string constraint equations.

DOI: 10.1103/PhysRevD.83.023529

PACS numbers: 98.80.Cq, 11.27.+d

I. INTRODUCTION

Cosmic strings may be observed through their gravitational wave emission, gravitational lensing of background galaxies, imprints on the cosmic microwave background, particle emission, and various other signatures. To work out the signatures of cosmic strings it is essential to know the properties of the string network, the number density of loops, and their length and shape distributions. Several observational signatures, including all burst events, whether gravitational or other, depend on the sudden whip-like motion of the string, a feature called a “cusp.” Yet other signatures are affected by sharp corners on a string, called “kinks.” (See Ref. [1] for a review.) Hence, to make reliable observational predictions from cosmic strings, it is important to characterize the generic properties of cosmic string loops and to quantify the frequency with which kinks and cusps occur.

Scherrer and Press (SP) [2] undertook a study of the dynamics of large loops starting with randomly chosen loops containing 10 harmonic modes. As each loop oscillates, it self-intersects and reconnects, breaking into smaller loops. In SP the loop is approximated by 128 discrete segments, thus limiting the resolution of their study. The study in SP was followed up by Scherrer, Quashnock, Spergel, and Press [3] with a focus on evaluating the gravitational power emitted from string loops. A similar study was undertaken by Casper and Allen (CA) [4] and is the most extensive study done to date. CA considered loops with 10 harmonics and discretized the loop using 600 segments.

In the present work we have built upon the SP scheme, improvements coming from more computational power

and choice of algorithms. To put our current study in perspective, we have considered loops with 3–50 harmonics and discretized using several times 10 000 segments. Our very fine resolution ensures that discretization effects are unimportant as will be clear from our results. A comparison of previous loop fragmentation studies with this one is given in Table I.

Our conclusions can be summarized as follows. Large loops with many oscillation harmonics split into smaller loops until the small loops have a *minimal* number of harmonics. In other words, the number of fragments is directly proportional to the number of harmonics on the

TABLE I. Parameters from studies of loop fragmentation. Modes refers to M in Eq. (4). The numbers provided here are for the “type A” loops defined by SP and used by CA; see the discussion after Eq. (4) for more details. For SP see [2] and for CA see [4].

Source	Segments	Modes (M)	Initial loops (N)	Stable loops
SP	128 ^a	10	20	561
CA	600 ^b	10	200	5723
Present work	10 000	3	1000	8308
	10 000	10	3000	94 628
	10 000	20	1000	63 490
	10 000	30	1000	96 207
	10 000	40	1000	128 764
	10 000	50	1000	157 968
	50 000	10	1000	32 158
	50 000	50	1000	162 157

^aLoops were also run with 256 segments without producing significantly more small loops.

^bThe loops were rerun with 800 segments resulting in almost no new daughter loops being produced.

*<http://www.phys.cwru.edu/projects/strings/>

initial loop; on average $3M$ loops are created from an initial loop with M modes. The final non-self-intersecting (“stable”) loops are approximately planar and their left- and right-moving modes tend to be orthogonal. The average stable loop contains four kinks and has a 40% chance of containing a cusp. Hence, realistic cosmic string loops should be visualized as oscillating rectangular shapes, like the “degenerate kinky loops” discussed in Ref. [5]. This picture is strikingly clear in the animations available in Ref. [6].

The conclusion that stable loops are close to degenerate kinky loops also means that gravitational and other radiation from stable loops can be estimated by the radiation from degenerate kinky loops, which is simple to calculate analytically [5]. Knowing the shape of loops can also help with other observational signatures, e.g. lensing.

Our detailed results can be found in Sec. III. At a finer level, the shape of a cosmic string loop may also depend on its length. We investigate this question in Sec. IV and find only weak dependence. It is a fair approximation to think of long loops as also being approximately planar and rectangular. In Sec. V we give analytical arguments that help us understand our numerical results. Notable is our demonstration of a set of perturbed degenerate kinky loops that are stable. We conjecture that such loops are attractors for the evolution of cosmic string loops.

In the appendixes we describe technical details of the numerical algorithm, formulas to boost to the rest frame of a loop, and finally, a new scheme to explicitly solve the Nambu-Goto equations and the string constraint equations. This new scheme is very suitable for constructing cosmic string loops but in our numerical work we use the SP scheme to aid comparison to previous work.

II. PROCEDURE

A string in flat spacetime is described by $X^\mu = (t, \mathbf{x}(\sigma, t))$, where σ is the parameter along the string and t is Minkowski time. The solution of the Nambu-Goto string equations yields independent left- and right-moving modes

$$\mathbf{x}(\sigma, t) = \frac{1}{2}[\mathbf{a}(\sigma_-) + \mathbf{b}(\sigma_+)], \quad (1)$$

where $\sigma_\pm = \sigma \pm t$. We choose equal intervals of σ to label equal amounts of energy which gives the additional constraint

$$|\mathbf{p}| = 1 = |\mathbf{q}|, \quad (2)$$

where $\mathbf{p} \equiv \mathbf{a}'$ and $\mathbf{q} \equiv \mathbf{b}'$. Closure of a loop of length L implies that

$$\int_0^L \mathbf{p} d\sigma = - \int_0^L \mathbf{q} d\sigma \quad (3)$$

while, in the center of momentum of the loop, each of the two integrals vanishes. A helpful geometrical picture is that the vector functions \mathbf{p} and \mathbf{q} trace out a path on the

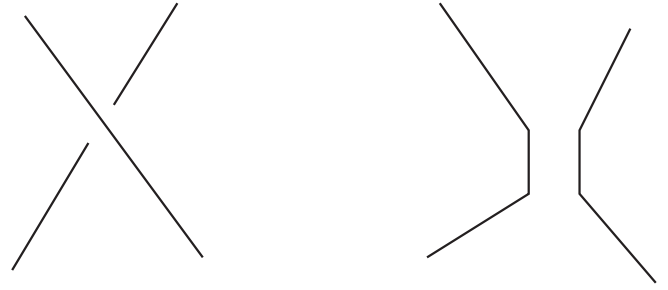


FIG. 1. When two strings intersect, they reconnect in an intercommutation event with the production of four kinks (discontinuities in the tangent to the string), two on each string.

surface of a two-dimensional sphere, also called the “Kibble-Turok (KT) sphere.”

The velocity of a point on the string is $(\mathbf{q} - \mathbf{p})/2$. Hence, at points where the curves corresponding to \mathbf{p} and $-\mathbf{q}$ intersect, the velocity of the string reaches the speed of light. Such momentarily lightlike points on the string are known as cusps. Finally note that \mathbf{p} and \mathbf{q} need not be continuous. If there is a break in either of the curves, it implies a discontinuity in the tangent vector to the string, which will appear as a sharp corner or kink.

The Nambu-Goto description of a cosmic string breaks down when two string segments intersect. The Nambu-Goto dynamics has to be supplemented by the condition that the strings intercommute, i.e. reconnect, at the point of intersection. This is sketched in Fig. 1, which also shows the four kinks that are created during intercommutation, two on each string.

To numerically model strings we follow SP and choose initial loops from an ensemble constructed by expanding the left and right movers in Fourier series

$$\begin{aligned} a_i(s) &= \sum_{m=1}^M a_{m,i} \cos(ms + \phi_{a,m,i}), \\ b_i(s) &= \sum_{m=1}^M b_{m,i} \cos(ms + \phi_{b,m,i}), \end{aligned} \quad (4)$$

where i stands for the three components of the vectors. Here we have taken $s \in [0, 2\pi]$. The mode coefficient vectors \mathbf{a}_m and \mathbf{b}_m are chosen so that each component of these vectors is uniformly distributed in the interval $[0, 1]$. The phases $\phi_{a,m,i}$ and $\phi_{b,m,i}$ are chosen to be uniformly distributed in the interval $[0, 2\pi]$. This corresponds to the “type A” scheme in SP and CA. Had we chosen the length of the mode coefficient vectors in the interval $[0, 1/m^2]$, we would have a direct correspondence with the “type B” scheme of SP and CA. As already discussed in CA, the type A loops have more power on small scales and lead to more fragmentations. For this reason we have only studied type A loops in this work. However, the shape of the final stable loop population is expected to be the same in both schemes. We have checked this by considering several

different values of M since this controls the power on small scales. A comparison of our choice of parameters with those of earlier work is given in Table I.

The challenge now is to meet the constraints in Eq. (2). This is done by recognizing that we are free to choose the parameter s as a function of the parameter σ . By differentiating Eq. (4) we obtain

$$1 = |\mathbf{p}| = \left| \frac{ds}{d\sigma} \right| \left| \sum_{m=1}^M \mathbf{a}_m m \sin(ms + \phi_{a,m}) \right|. \quad (5)$$

This provides a differential equation for σ in terms of s which we solve numerically with the additional requirement that $ds/d\sigma \geq 0$. To ensure that $s(\sigma) \in [0, 2\pi]$ for $\sigma \in [0, 1]$ we first rescale \mathbf{a}_m and \mathbf{b}_m before solving (5). A numerical inversion then gives s as a function of σ which finally allows us to obtain $\mathbf{p}(\sigma)$ and $\mathbf{q}(\sigma)$ at the initial time $t = 0$. Note that the value of M corresponds to the “number of harmonics” on the loop in terms of the parameter s but not in terms of the invariant length σ . However we will still refer to M as the number of harmonics.

Once we know the function $\mathbf{p}(\sigma)$ and $\mathbf{q}(\sigma)$, the Nambu-Goto evolution is straightforward since only the arguments of the functions change as given in Eq. (1). The numerical implementation of intercommutation events is more involved and described in Appendix A.

An intercommutation causes the original loop to fragment into two smaller loops, with each of the two loops having one left-moving kink and one right-moving kink. In this way, the loops keep fragmenting with the production of more and more kinks. The fragmentation ceases when a loop reaches a non-self-intersecting configuration. It is these stable loops that we wish to study and characterize.

In principle, two fragments of the initial loop can intersect and reconnect to form a bigger loop. In practice, since the fragments are flying apart, such collisions do not happen frequently. In any case, we ignore such mergers.

III. NUMERICAL RESULTS

Most of our data is collected for loops of 10000 segments and 10 harmonic modes; see Table I. In Fig. 2 we show a sample initial loop. We evolve the loop as described in Sec. II and keep track of the fragmentations. We define loop “generations” in the following way. The initial loop is the first generation. When it splits into two loops, those are in the second generation, and so on for successive generations. Once a loop is in a stable configuration, it is counted in each subsequent generation. For example, a stable loop that is generated in the third generation will be counted in the third, fourth, etc., generations. At every generation, we can plot the length distribution of loops as shown in Fig. 3. The plot shows that the second generation has roughly a uniform distribution of lengths but that the stable loop distribution is sharply peaked at small lengths. The stable loop length distributions are plotted in Fig. 4.

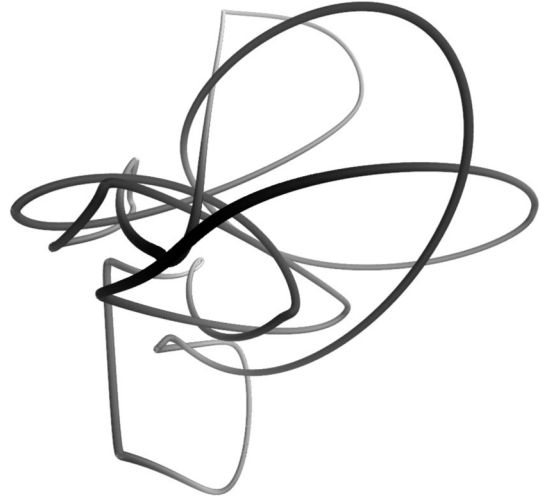


FIG. 2. A sample initial loop with 10 harmonic modes. The darker, thicker segments of the loop are nearer the viewer.

Though very small loops are formed through fragmentation, the peaks in this figure at $L \sim 2/N$ show the resolution limit of the simulations. Clearly the bulk of the stable loops are above our resolution limit. For comparison the SP and CA results are shown and are in good agreement with our results.

From the length distribution we calculate the average length per generation for several values of M . In Fig. 5 we show the average length scaled by M for the stable loops. All the plots for $M = 10, \dots, 50$ asymptote to the same value, 0.33, which clearly demonstrates that the average length of stable loops is proportional to $1/M$. The plot for $M = 3$ asymptotes to 0.36 and suggests that loops with only very low harmonics could behave differently from

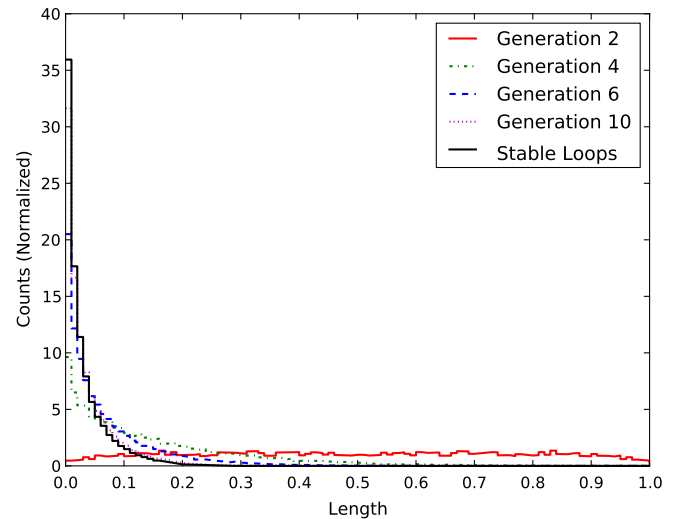


FIG. 3 (color online). Length of loops versus generation. By generation 10 the distribution is nearly identical to the stable loop length distribution. As fragmentation occurs the length of loops decreases. The distribution of the predominantly small length stable loops can be seen in Fig. 4.

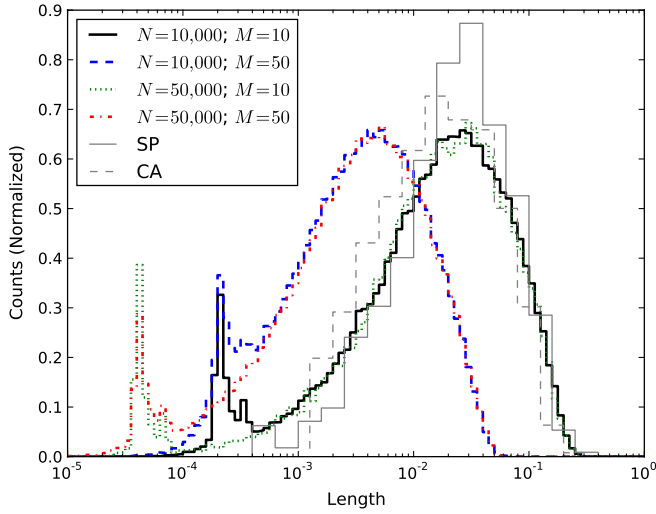


FIG. 4 (color online). Length distribution of stable loops shown on a logarithmic scale to better see the small length behavior. The resolution limit of the simulations is seen by the small peaks at $L \sim 2/N$. The length distribution is otherwise independent of resolution, N . Comparison to the SP (solid, gray line) and CA (dashed, gray line) shows good general agreement.

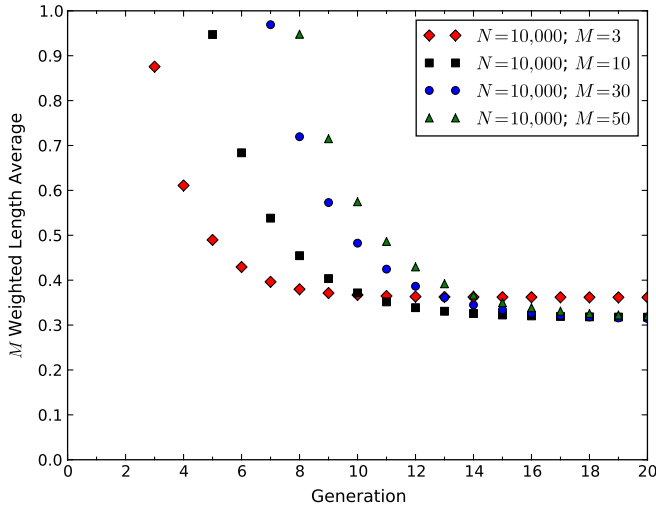


FIG. 5 (color online). Average loop length multiplied by the number of harmonics in the parent loop versus generation. This weighted length asymptotes to ~ 0.33 for all initial loops.

loops with high harmonics. However, even the $M = 3$ value of 0.36 is very close to 0.33, and we will disregard this difference in the discussion below. Then the asymptote at $0.33 = 1/3$ implies that an initial loop with M harmonics fragments into $3M$ stable loops, as can also be read off from Table I. Given that the initial loop is straight on the length scale $1/M$, the stable loop is also composed of roughly straight segments of length $\sim 1/3M$. This result solidifies and quantifies the finding in Ref. [7] that fragmentation does not continue indefinitely to smaller and smaller loops.

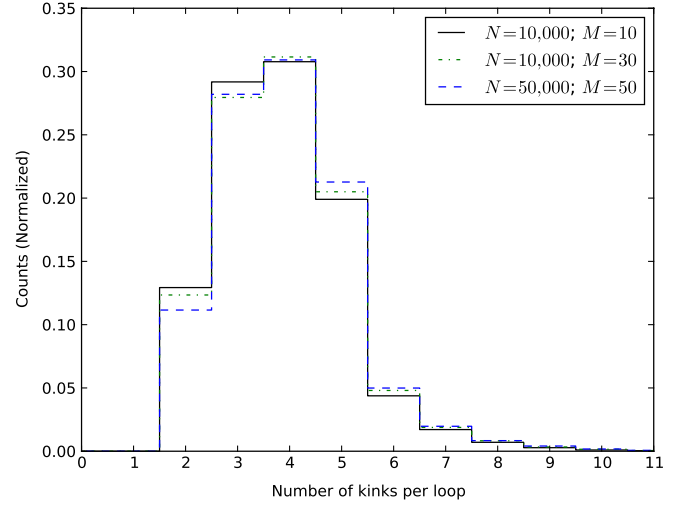


FIG. 6 (color online). The number of kinks on stable loops. Each loop must have at least 2 kinks. Most of the loops have between 2 and 5 kinks independent of resolution and the number of harmonic modes.

In terms of the \mathbf{p} and \mathbf{q} curves on the KT sphere, since a loop with only the fundamental harmonic wraps around a great circle once, we expect that stable loops will wrap the sphere only a third of the way. Also, the discussion below shows that the \mathbf{p} and \mathbf{q} curves for the stable loops occur in two disconnected segments on the KT sphere. Hence each segment only covers a sixth of the KT sphere and can be thought of as an arc on the KT sphere that extends 60° on average.

Our result that the stable loops are made from roughly straight segments can only work if the loops have kinks on them. It can be argued that the average number of kinks should be four [4]. Assume that there is an initial loop with N_k preexisting kinks. (In our numerical work, the initial loops were smooth and so $N_k = 0$.) Every intercommutation event adds four kinks and one extra loop to the system. Therefore after n intercommutations there will be $4n + N_k$ kinks and $n + 1$ loops. So the average number of kinks per loop is $(4n + N_k)/(n + 1)$, which goes to 4 in the large n limit. The distribution of kinks, however, needs to be calculated numerically and is shown in Fig. 6. From this argument and the numerical simulations we have the picture that stable cosmic string loops have roughly straight sides and four corners (kinks) on average.

In general, stable loops have relativistic center of mass velocities as shown in Fig. 7. To analyze the shape of a stable loop, we must first transform to its center of mass frame. Care must be taken to transform the coordinates so that the gauge conditions in Eqs. (2) are satisfied as described in Appendix B. Let us denote the resulting curves on the KT sphere in the rest frame by \mathbf{P} and \mathbf{Q} .

To analyze the planarity of a stable loop, we calculate the “moment of inertia” tensors for \mathbf{P} and \mathbf{Q} via

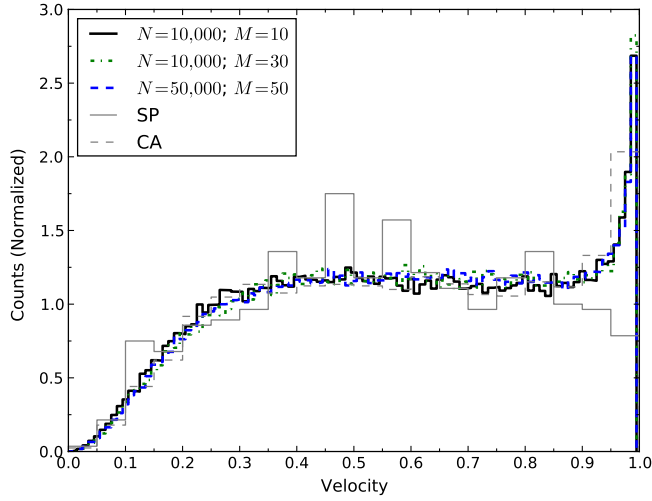


FIG. 7 (color online). The velocity distribution of stable loops for various resolutions and numbers of modes. The distribution is universal and sharply peaked at $v \approx 1$. Also shown are the SP (solid, gray line) and CA (dashed, gray line) for comparison. The resolution of SP did not allow them to produce high velocity loops. Our results refine the peak at high velocity found by CA.

$$A_{ij} = \int P_i P_j d\sigma_-, \quad B_{ij} = \int Q_i Q_j d\sigma_+. \quad (6)$$

If the \mathbf{P} (and similarly for the \mathbf{Q}) curve is uniformly distributed on the KT sphere, the three eigenvalues of the A tensor would be equal. If one of the eigenvalues of A vanishes, then \mathbf{P} is distributed in a plane; and if two eigenvalues vanish, the distribution is lineal. In Fig. 8 we show the average eigenvalues as the loop continues to fragment. It is clear that one of the eigenvalues vanishes, and the largest eigenvalue is much larger than the middle eigenvalue. So the \mathbf{P} curve is mostly in one direction on the KT sphere (say, around the z axis), with a little spread in some other direction. Since the centroid of the \mathbf{P} curve has to vanish in the center of mass frame [see the discussion below Eq. (3)], the \mathbf{P} curve must contain two short segments, say, one near the north pole and the other near the south pole. In an idealized (“degenerate”) case, the \mathbf{P} curve would consist of just 2 points on the KT sphere. Similarly the \mathbf{Q} curve would consist of just 2 other points. Since the loop is a sum of the left and right movers, this implies that the loop is planar and defined by the plane of the \mathbf{P} and \mathbf{Q} curves.

As another check of this picture, we have plotted the kink sharpness for the \mathbf{p} curves, defined as [8]

$$\psi_a = \frac{1}{2}(1 - \mathbf{P}_- \cdot \mathbf{P}_+), \quad (7)$$

where \mathbf{P}_\pm are the values of \mathbf{P} on either side of the kink. The plot in Fig. 9 demonstrates a peak near sharpness of 1, which implies $\mathbf{P}_- \approx -\mathbf{P}_+$. This shows that there are 180° jumps on the KT sphere, in accordance with our picture

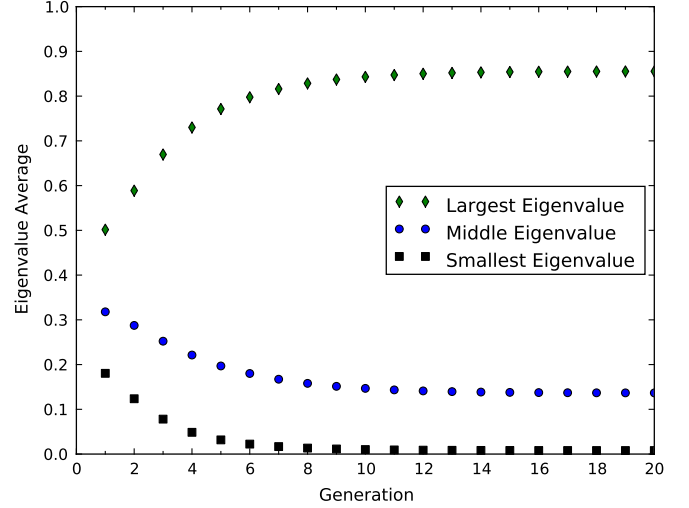


FIG. 8 (color online). The average of the eigenvalues of the \mathbf{P} and \mathbf{Q} “moment of inertia” tensors (6) versus loop generation. Since both \mathbf{P} and \mathbf{Q} have the same properties, both are included in the average to increase the statistics. We notice that the smaller eigenvalue approaches zero and the largest eigenvalue is much larger than the middle eigenvalue showing that the \mathbf{P} and \mathbf{Q} curves are localized on the KT sphere. This plot is for $N = 10000$, $M = 10$. The asymptotic behavior is independent of resolution and the number of modes.

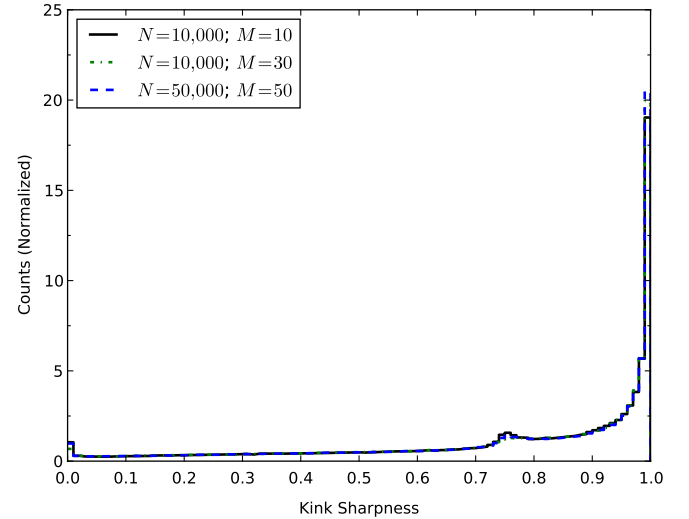


FIG. 9 (color online). The kink sharpness distribution [see Eq. (7)]. The distribution is independent of resolution and the number of modes on the initial string. We see that the distribution is sharply peaked at 1, that is, where the \mathbf{P} before and after the kink are antiparallel.

that the \mathbf{P} curve corresponds to two small antipodal regions on the KT sphere.

We can go a bit further and ask for correlations between the left and right movers. For this we plot the distribution of $\langle (\mathbf{P} \cdot \mathbf{Q})^2 \rangle$ in Fig. 10. (Angular brackets denote average over a loop.) This plot shows that \mathbf{P} and \mathbf{Q} tend to get more orthogonal with fragmentation.

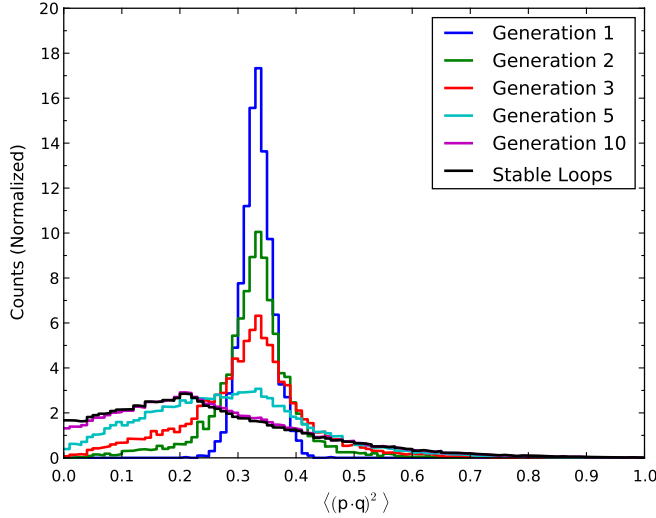


FIG. 10 (color online). The distribution of $\langle (\mathbf{P} \cdot \mathbf{Q})^2 \rangle$ showing that the left and right movers become more orthogonal to each other with more fragmentation.

This picture of stable loops implies that the \mathbf{P} and $-\mathbf{Q}$ curves rarely cross each other and hence that cusps should be suppressed. For an initial loop with M modes we expect it to contain approximately M^2 cusps. (Each mode roughly corresponds to a great circle on the KT sphere so the number of intersections is proportional to M^2 .)

We have seen that on average $3M$ stable loops are produced. If f is the “cusp survival fraction”—the fraction of cusps on the initial loop that survive on the stable loops—then on average we expect $M^2 f / 3M$ cusps on each loop. Figure 11 shows that, on average,

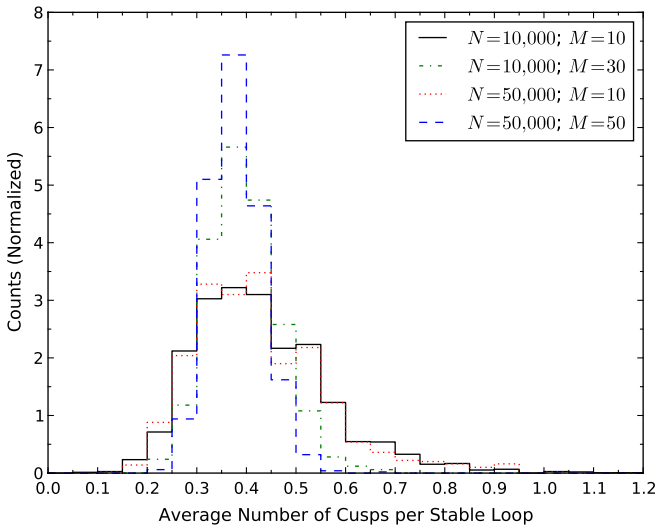


FIG. 11 (color online). The average number of cusps on each stable loop. On average we see that there is approximately a 40% chance of a stable loop to contain a cusp. This result is independent of resolution and the number of modes on the initial loop.

loop has a 40% of containing a cusp independent of resolution and the number of modes on the initial loop so the cusp survival fraction is found to be $f \approx 6/5M$. (We have verified this is true in our simulations but do not show the results here.)

IV. LENGTH-DEPENDENT RESULTS

The focus of this work is on the average properties of all stable loops. The statistics we have presented give equal weights to all loops. For some observational signatures, however, long loops may be more relevant than short loops. Hence, in this section, we provide some statistics as a function of loop length. The results shown here are for the $N = 10\,000$ and $M = 10$ runs. Logarithmic length bins are chosen based on the distribution of loop lengths (see Fig. 4). Ten bins are chosen between lengths of 10^{-3} and 0.1. The initial bin begins at a length of 5×10^{-4} to ensure that all loops are above our resolution limit; approximately 94% of the loops are included. The final bin includes all loops with lengths longer than 0.1.

The results are shown in Figs. 12–15. These figures show the velocity, number of kinks, number of cusps, and eigenvalues, respectively, for each length bin. In the plots the x error bar shows the width of the bin and the y error bar provides the 95 percentile range for the values. The results are as expected. The smallest loops are predominantly high velocity, contain approximately four kinks per loop, have a small chance of containing a cusp, and are planar. The largest loops are predominantly lower velocity, contain

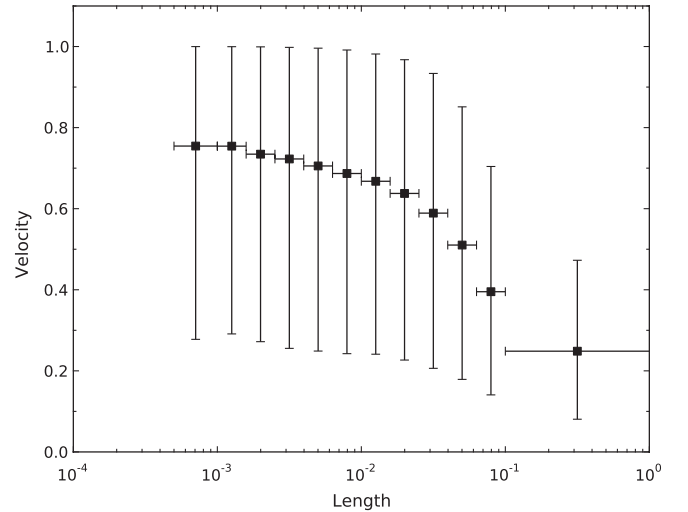


FIG. 12. The velocity of stable loops versus the length of the loop. The x error bars show the width of the length bins employed, and the y error bars show the 95 percentile ranges in each length bin. The square symbol shows the average in the bin. Short loops have center of mass velocity $\sim 1/\sqrt{2}$, which is the root-mean-squared velocity of the string in the initial loop. Longer loops tend to have a lower velocity consistent with the fact that the initial loop is at rest.

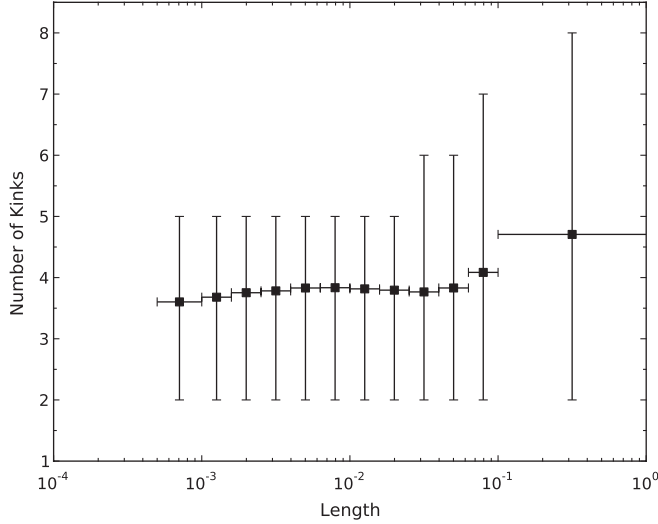


FIG. 13. Similar to Fig. 12 now for the number of kinks on stable loops versus the length of the loop. Longer loops tend to have a larger number of kinks though the variation is slight. The loop with the maximum number of kinks has 8 kinks instead of the canonical 4.

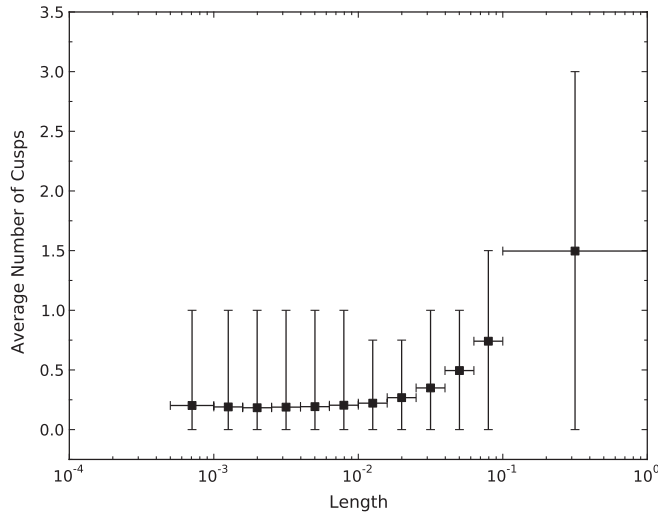


FIG. 14. Similar to Fig. 12 now for the average number of cusps on stable loops versus the length of the loop. Longer loops typically contain more cusps on average. The loop with the maximum number of cusps has only 3 cusps.

more than four kinks, have a higher chance of containing a cusp, and are less planar.

V. DISCUSSION

Our numerical results point to the picture that stable loops are deformations of degenerate kinky loops. In this section we provide further evidence for this picture by demonstrating that there exists a class of perturbed degenerate kinky loops that are stable.

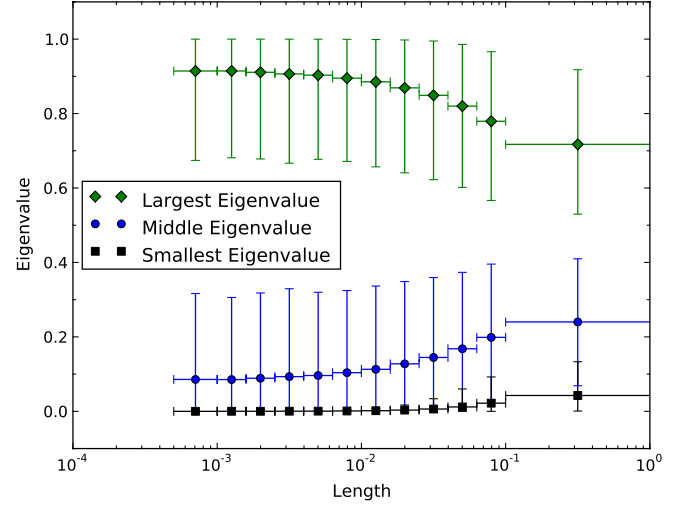


FIG. 15 (color online). Similar to Fig. 12 now for the eigenvalues of the moment of inertia tensors of the stable loops versus the length. Longer loops are also planar and only slightly less so than smaller loops.

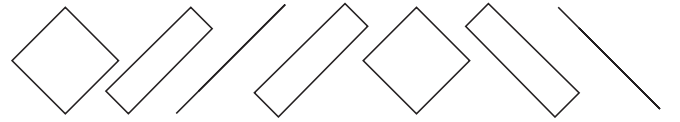


FIG. 16. Snapshots during the evolution of the degenerate kinky loop. Note that the loop periodically collapses to a double line.

We begin by describing the degenerate kinky loop [5]

$$\mathbf{p}_{\text{dk}} = \cos(\alpha_0)\hat{z}, \quad \mathbf{q}_{\text{dk}} = \cos(\beta_0)\hat{x}, \quad (8)$$

where

$$\alpha_0 = \pi[2\sigma_-], \quad \beta_0 = \pi[2\sigma_+], \quad (9)$$

$[x]$ denotes the greatest integer less than or equal to x , and $\sigma_{\pm} \in [0, 1]$. In Fig. 16 we show some snapshots of a degenerate kinky loop.

Note that the degenerate kinky loop collapses to a double line twice during an oscillation period, yet the stable loops do not self-intersect. This must be due to the fact that stable loops are not exactly degenerate. To show that this is the case, consider

$$\mathbf{p} = \cos(\alpha)\hat{z}, \quad (10)$$

$$\mathbf{q} = \cos(2\pi\epsilon\sigma_+ + \beta)\hat{x} + \sin(2\pi\epsilon\sigma_+ + \beta)\hat{y},$$

where $0 < \epsilon < 1$ and now

$$\alpha = \pi[2\sigma_-], \quad \beta = (1 - \epsilon)\pi[2\sigma_+]. \quad (11)$$

As illustrated in Fig. 17, these perturbations stretch out the point corresponding to \mathbf{q} on the KT sphere for the degenerate kinky loop to an arc of length $\epsilon\pi$ [5].

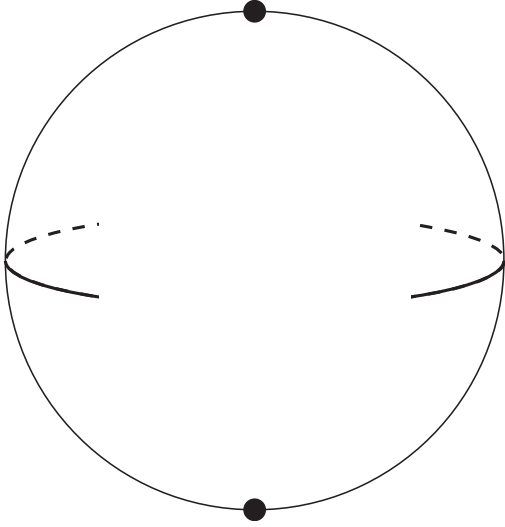


FIG. 17. The \mathbf{p} and \mathbf{q} curves for a perturbed degenerate kinky loop (10) plotted on the KT sphere. The \mathbf{p} curve is represented by the two points at the poles. The \mathbf{q} curve is represented by the arcs along the equator.

We can check explicitly that the self-intersection condition

$$\mathbf{x}(\sigma, t) = \mathbf{x}(\sigma', t) \quad (12)$$

is satisfied for such a loop if and only if $\sigma = \sigma' \pmod{1}$, for any t ; that is, the loop does not intersect at any time. We first write the intersection condition in terms of \mathbf{a} and \mathbf{b} :

$$\mathbf{a}(\sigma - t) + \mathbf{b}(\sigma + t) = \mathbf{a}(\sigma' - t) + \mathbf{b}(\sigma' + t). \quad (13)$$

Since for our loop (10) we have $\mathbf{p} \cdot \mathbf{q} = 0$ for all σ and t , we also have $\mathbf{a} \cdot \mathbf{b} = 0$, and self-intersection requires

$$\mathbf{a}(\sigma - t) = \mathbf{a}(\sigma' - t), \quad \mathbf{b}(\sigma + t) = \mathbf{b}(\sigma' + t). \quad (14)$$

The first condition is satisfied for $\sigma_- = \sigma'_-$ or for $\sigma'_- = 1 - \sigma_-$. Only the latter possibility is of interest since we are looking for an intersection with $\sigma \neq \sigma'$. The second condition for the x and y components leads to

$$\begin{aligned} \sin(2\epsilon\pi\sigma_+) &= \sin(\epsilon\pi) - \sin(2\epsilon\pi\sigma'_+ - \epsilon\pi), \\ \cos(2\epsilon\pi\sigma_+) &= 1 + \cos(\epsilon\pi) - \cos(2\epsilon\pi\sigma'_+ - \epsilon\pi). \end{aligned} \quad (15)$$

Note that *both* equations must be satisfied for a self-intersection to occur.

With the help of some trigonometric identities, we can rewrite (15) as

$$\begin{aligned} &\sin\left[\epsilon\pi\left(\sigma_+ + \sigma'_+ - \frac{1}{2}\right)\right] \cos\left[\epsilon\pi\left(\sigma_+ - \sigma'_+ + \frac{1}{2}\right)\right] \\ &= \sin\left(\frac{\epsilon\pi}{2}\right) \cos\left(\frac{\epsilon\pi}{2}\right), \\ &\cos\left[\epsilon\pi\left(\sigma_+ + \sigma'_+ - \frac{1}{2}\right)\right] \cos\left[\epsilon\pi\left(\sigma_+ - \sigma'_+ + \frac{1}{2}\right)\right] \\ &= \cos^2\left(\frac{\epsilon\pi}{2}\right). \end{aligned} \quad (16)$$

Taking the ratio of these equations leads to

$$\tan\left[\epsilon\pi\left(\sigma_+ + \sigma'_+ - \frac{1}{2}\right)\right] = \tan\left(\frac{\epsilon\pi}{2}\right), \quad (17)$$

so that $\sigma'_+ + \sigma_+ = 1$. By squaring and adding the equations, we obtain

$$\cos\left[\epsilon\pi\left(\sigma_+ - \sigma'_+ + \frac{1}{2}\right)\right] = \pm \cos\left(\frac{\epsilon\pi}{2}\right), \quad (18)$$

so that the second relation in Eq. (16) leads to

$$\cos[\epsilon\pi(\sigma_+ - \sigma'_+ + \frac{1}{2})] = \cos[\epsilon\pi(\sigma_+ + \sigma'_+ - \frac{1}{2})] \quad (19)$$

or

$$\sin(\epsilon\pi\sigma_+) \sin[\epsilon\pi(\sigma'_+ - \frac{1}{2})] = 0, \quad (20)$$

which implies $\sigma_+ = 0$ or else $\sigma'_+ = 1/2$. (Solutions such as $\sigma_+ = 1/\epsilon$ can be ignored because we restrict σ_+ to lie in the interval $[0, 2\pi]$ and $|\epsilon| < 1$.) Therefore the only solution is of the type $\sigma_+ = 0$, $\sigma'_+ = 1$ or else $\sigma'_+ = 1/2$, $\sigma_+ = 1/2$, which are both trivial. This shows that there are no self-intersections. The conclusion is intuitive since \mathbf{b} is obtained by integrating \mathbf{q} and hence is a vector that lies in the xy plane (see Fig. 17) and rotates in this plane. Since $0 < \epsilon < 1$, \mathbf{b} never attains the same vector value twice for $\sigma_+ \in [0, 1]$.

The perturbed degenerate kinky loop discussed above does not contain any cusps while our numerical results show that roughly half of the stable loops have a cusp. To understand the \mathbf{P} and \mathbf{Q} curves of a stable loop with cusps we show an example of these curves for a sample loop from our numerical runs in Fig. 18, which contains three kinks and one cusp. In this example the \mathbf{P} curve is like our perturbed degenerate kinky loop but the \mathbf{Q} curve is much more elongated. The two curves are in planes nearly perpendicular to each other.

To construct a stable loop with cusps consider

$$\begin{aligned} \mathbf{p} &= \sin(\beta_-)\hat{x} + \cos(\beta_-)\hat{z}, \\ \mathbf{q} &= \sin(\beta_+)[\sin(\phi_+)\hat{x} + \cos(\phi_+)\hat{y}] - \cos(\beta_+)\hat{z}, \end{aligned} \quad (21)$$

where

$$\beta_{\pm} = 2\pi\epsilon_{\pm}\sigma_{\pm} + (1 - \epsilon_{\pm})\pi[2\sigma_{\pm}] \quad (22)$$

and ϕ_+ is an arbitrary phase. There are two cusps due to the intersections at $\pm\hat{z}$. These occur at $\sigma = t = 0$ and

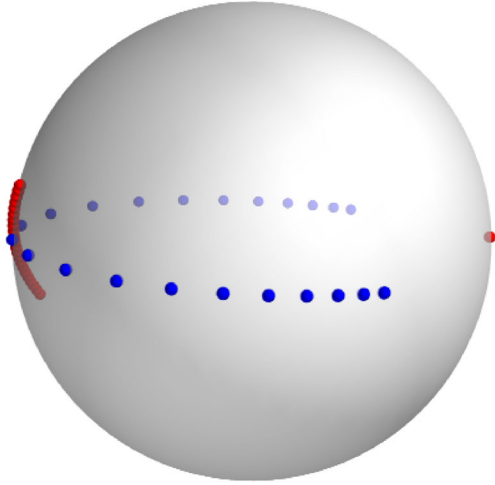


FIG. 18 (color online). The P and $-Q$ curves plotted on the KT sphere for a stable loop from our simulations. This loop contains three kinks and one cusp. The two curves are represented by the dots. The Q curve is the dots along the equator. The P curve is the dots perpendicular to the equator (running vertically on the left of the sphere) and includes the lone antipodal point (on the right of the sphere). Note that the points on the sphere are not weighted equally and the centroids of both P and Q curves are at the center of the sphere.

$\sigma = t = 1/2$. We have numerically confirmed that this loop is stable for the parameter choice

$$\epsilon_- = \frac{1}{\sqrt{58}}, \quad \epsilon_+ = \frac{1}{\sqrt{95}}, \quad \phi_+ = \frac{\pi}{\sqrt{67}}. \quad (23)$$

In Fig. 18 we show the p and q curves for a stable loop with cusps similar to (21) extracted from our simulations.

VI. CONCLUSIONS

We have shown that the fragmentation of a large complicated loop yields stable quasipolar, quasirectangular loops that are similar to degenerate kinky loops. If cosmic string loops are the result of a large number of intercommutings—even as few as five or six intercommutings may be sufficient (e.g. Fig. 8)—we expect that they too will be quasirectangular. Earlier work has examined the stability of relatively simple loops (with a small number of harmonics) and found stability over a region of parameter space [9–12]. It is unclear if cosmological evolution can directly produce these stable loops with few harmonics. Cosmological loops will certainly contain kinks, and this feature is missing in the loops investigated in these earlier studies.

In a cosmological setting, when a loop is produced from the infinite string network, it will fragment down to stable loops within two oscillation periods, and therefore only stable loops are relevant for cosmological signatures. Indeed the aim of the CA and Scherrer, Quashnock,

Spergel, and Press studies was to obtain the gravitational power emitted from realistic cosmic string loops. We can also estimate the gravitational power based on the analytical results of [5] for emission from degenerate, kinky loops and obtain an estimate that is within 13% of the CA result. The gravitational power radiated from degenerate kinky loops is analytically calculated to be $64 \ln(2) G \mu^2 \approx 44 G \mu^2$, whereas CA numerically estimate $39 G \mu^2$ from their simulations, where μ is the string tension.

For cosmic strings, Hubble expansion, frictional damping, and radiation damping also come into play, though the effects are generally on very long time scales and depend on the environment in which the strings are placed. Hence we expect that when a large loop is produced from the network, it will fragment into stable loops within two oscillation periods and then, as these stable loops oscillate for many oscillation periods, damping effects start to play a role and to change the Nambu-Goto dynamics. So to characterize the effect of damping on the cosmic string network loops, we can limit our study to the damping effects on the perturbed degenerate kinky loops and the stable loops with cusps described above.

ACKNOWLEDGMENTS

We are grateful to Jose Blanco-Pillado, Ken Olum, Ben Schlaer, and Alex Vilenkin for discussions and Isha Savani for help during the initial algorithm development. This work was supported by the U.S. Department of Energy at Case Western Reserve University. T. V. was also supported by Grant No. DE-FG02-90ER40542 at the Institute for Advanced Study. The numerical simulations were performed on the facilities provided by the Case ITS High Performance Computing Cluster.

APPENDIX A: LOOP INTERCOMMUTATION

The numerical evolution in this work was performed using the “diamond method” for string evolution which is discussed in [13]. Full details of our implementation along with source code and documentation are available online [6]. Here we will discuss one aspect of the evolution: loop intercommutation.

The diamond method gets its name from the σ_{\pm} grid shown in Fig. 19. The top of the figure shows the grid structure for an initial loop. Our initial loop segments are of uniform length $\Delta\sigma$, producing a regular grid. This grid provides a simple geometric picture of the loop. With this grid structure, a starting point $x(0, 0)$, and values for $p(\sigma_-)$ and $q(\sigma_+)$, taken to be piecewise constant functions, we can calculate the coordinate of the string $x(\sigma, t)$ during its evolution if intersections do not occur.

In Fig. 19 we consider the case of an intersection at points a and b , that is, $x(\sigma_a, t_{\text{int}}) = x(\sigma_b, t_{\text{int}})$. In this picture the intercommutation and production of two loops is geometric. The shaded area is excised from the parent

APPENDIX B: LOOP REST FRAME TRANSFORMATION

The time coordinate of a loop, t , was chosen to be the proper time and the same as the background spacetime $x^0 = t$; see Sec. II. To make \mathbf{p} and \mathbf{q} easier to boost we construct four-vectors from them with the time components give by $p^0 = -1$ and $q^0 = 1$. At a fixed time, t , we may then write in the initial loop's rest frame

$$p^\mu = \frac{\partial a^\mu}{\partial \sigma_-} = \begin{pmatrix} -1 \\ \mathbf{p} \end{pmatrix}, \quad (\text{B1})$$

where $|\mathbf{p}| = 1$.

The center of mass velocity of a loop is given by

$$\begin{aligned} \mathbf{v}_{\text{cm}} &= \frac{1}{2} \left(\int \mathbf{q} d\sigma_+ - \int \mathbf{p} d\sigma_- \right) = \int \mathbf{q} d\sigma_+ \\ &= - \int \mathbf{p} d\sigma_-. \end{aligned} \quad (\text{B2})$$

With $\gamma = 1/\sqrt{1 - \mathbf{v}_{\text{cm}}^2}$ boosting to the rest frame of this loop we find

$$P^\mu(\sigma_-) = \begin{pmatrix} -\gamma(1 + \mathbf{v}_{\text{cm}} \cdot \mathbf{p}) \\ \mathbf{p} + (\gamma - 1)(\mathbf{v}_{\text{cm}} \cdot \mathbf{p})\mathbf{v}_{\text{cm}}/\mathbf{v}_{\text{cm}}^2 + \gamma\mathbf{v}_{\text{cm}} \end{pmatrix}. \quad (\text{B3})$$

Notice that the time component of this four-vector is no longer -1 . To correct this we apply a gauge transformation to the coordinate σ_- . Let

$$\tilde{\sigma}_- \equiv \gamma(\sigma_- + \mathbf{v}_{\text{cm}} \cdot \mathbf{a}), \quad (\text{B4})$$

so that

$$\frac{\partial \tilde{\sigma}_-}{\partial \sigma_-} = \gamma(1 + \mathbf{v}_{\text{cm}} \cdot \mathbf{p}). \quad (\text{B5})$$

An integration gives $\tilde{\sigma}_- \in [0, 1/\gamma]$ and a coordinate re-scaling can be used to bring the interval back to $[0, 1]$. With this we now have

$$\begin{aligned} P^0(\tilde{\sigma}_-) &= \frac{\partial \tilde{a}^0}{\partial \tilde{\sigma}_-} = \left(\frac{\partial \tilde{a}^0}{\partial \sigma_-} \right) \left(\frac{\partial \tilde{\sigma}_-}{\partial \sigma_-} \right)^{-1} = \frac{P^0}{\gamma(1 + \mathbf{v}_{\text{cm}} \cdot \mathbf{p})} \\ &= -1, \end{aligned} \quad (\text{B6})$$

where we have used (B3) for P^0 . Applying the same transformation to the spatial piece of (B3) we find

$$\mathbf{P}(\tilde{\sigma}_-) = \frac{\mathbf{p} + (\gamma - 1)(\mathbf{v}_{\text{cm}} \cdot \mathbf{p})\mathbf{v}_{\text{cm}}/\mathbf{v}_{\text{cm}}^2 + \gamma\mathbf{v}_{\text{cm}}}{\gamma(1 + \mathbf{v}_{\text{cm}} \cdot \mathbf{p})}. \quad (\text{B7})$$

Again this looks messy but we can verify that $|\mathbf{P}| = 1$, as required. We can further verify that

$$\int P(\tilde{\sigma}_-) d\tilde{\sigma}_- = \int P(\tilde{\sigma}_-) \left(\frac{\partial \tilde{\sigma}_-}{\partial \sigma_-} \right) d\sigma_- = 0. \quad (\text{B8})$$

For \mathbf{q} we proceed in the same way. In this case the gauge transformation is $\tilde{\sigma}_+ = \gamma(\sigma_+ - \mathbf{v}_{\text{cm}} \cdot \mathbf{b})$ and we find

$$\mathbf{Q}(\tilde{\sigma}_+) = \frac{\mathbf{q} + (\gamma - 1)(\mathbf{v}_{\text{cm}} \cdot \mathbf{q})\mathbf{v}_{\text{cm}}/\mathbf{v}_{\text{cm}}^2 - \gamma\mathbf{v}_{\text{cm}}}{\gamma(1 - \mathbf{v}_{\text{cm}} \cdot \mathbf{q})}. \quad (\text{B9})$$

Once again we can show that $|\mathbf{Q}| = 1$.

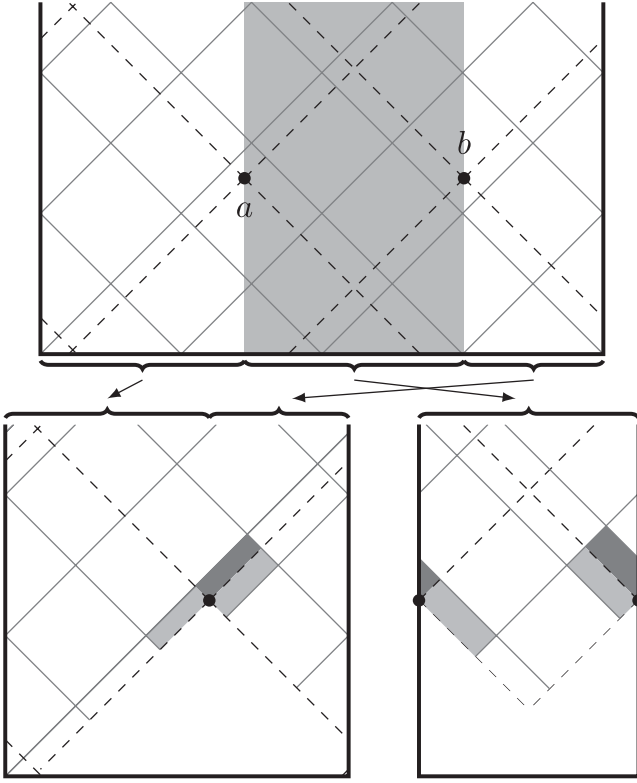


FIG. 19. Geometric splitting of a loop. Shown in this figure are the σ_\pm grids for various loops. The diagonal lines are lines of constant σ_- (up and to the right) and constant σ_+ (up and to the left). Time runs upward and σ rightward. The left and right edges of the grids are identified to describe a loop of string. The initial loop is shown on the top with the intersection occurring at points a and b . The dashed lines show the constant values of σ_\pm at which the intersection points occur and so at which kinks are formed. The daughter loops are formed by cutting the grid and identifying the new edges. In the figure the gray shaded region is cut from the grid and becomes one loop whereas the remaining unshaded pieces are connected. The new loops are shown below the initial loop. The past light cone of the intersection point has been removed since it no longer contains useful information. The two intersecting diamonds are replaced by the three shaded diamonds in the figure.

loop to become one daughter loop (bottom right grid) and the remaining portions of the parent loop are reconnected (bottom left grid). The past light cone of the intersection point no longer contains meaningful information and has been removed from the new grids. When an intersection occurs, the grid structure is modified; an originally uniform grid becomes nonuniform but otherwise contains the same information. In the bottom left grid the light gray shaded diamonds are the remnants of the original diamonds that intersected. The dark gray shaded diamond is a new one created by the reconnection of the grid after the middle portion has been removed. These two grids now describe independent loops which may be evolved in the usual way for all times $t > t_{\text{int}}$.

The above equations are not in the best form for numerical evaluation. If $|\mathbf{v}_{\text{cm}}| \approx 1$, then $\gamma \gg 1$ so large numbers will be subtracted from each other in the numerator. To correct this it is better to write the equations in terms of $\gamma^{-1} = \sqrt{1 - \mathbf{v}_{\text{cm}}^2}$ since $0 \leq \gamma^{-1} \leq 1$. This provides the alternative forms

$$\begin{aligned} \mathbf{p} &= \frac{\gamma^{-1} \mathbf{p} + (1 - \gamma^{-1})(\mathbf{v}_{\text{cm}} \cdot \mathbf{p})\mathbf{v}_{\text{cm}}/\mathbf{v}_{\text{cm}}^2 + \mathbf{v}_{\text{cm}}}{1 + \mathbf{v}_{\text{cm}} \cdot \mathbf{p}}, \\ \mathbf{q} &= \frac{\gamma^{-1} \mathbf{q} + (1 - \gamma^{-1})(\mathbf{v}_{\text{cm}} \cdot \mathbf{q})\mathbf{v}_{\text{cm}}/\mathbf{v}_{\text{cm}}^2 - \mathbf{v}_{\text{cm}}}{1 - \mathbf{v}_{\text{cm}} \cdot \mathbf{q}}. \end{aligned} \quad (\text{B10})$$

APPENDIX C: EXPLICIT SOLUTIONS OF THE STRING EQUATIONS AND CONSTRAINTS

Here we describe an explicit analytical solution of the Nambu-Goto equations of motion [Eq. (1)] and the string constraints [Eqs. (2) and (3)]. By choosing the decomposition in terms of left and right movers we can solve the Nambu-Goto equations as in Eq. (1). To solve the constraint in Eq. (2) we start with

$$\mathbf{p}(\sigma_-) = (\sin\theta \cos\phi, \sin\theta \sin\phi, \cos\theta), \quad (\text{C1})$$

where θ and ϕ are functions of σ_- . We similarly choose $\mathbf{q}(\sigma_+)$ with independent angular functions.

We require that \mathbf{p} be periodic under $\sigma_- \rightarrow \sigma_- + 1$. (Throughout this section we consider a loop of length 1.) Therefore we write

$$\theta(\sigma_-) = 2\pi j_- \sigma_- + \sum_{m=0}^{\infty} \theta_m \cos(2\pi m \sigma_- + \alpha_m), \quad (\text{C2})$$

where j_- is an integer and θ_m and α_m are arbitrary constants that can be chosen randomly to generate random loops. Similarly,

$$\phi(\sigma_-) = 2\pi k_- \sigma_- + \sum_{m=0}^{\infty} \phi_m \cos(2\pi m \sigma_- + \beta_m), \quad (\text{C3})$$

where k_- is an integer and ϕ_m and β_m are arbitrary constants.

Next we come to the closure condition in Eq. (3). The integral of \mathbf{p} will, in general, not vanish. To correct this, define

$$\mathbf{v}_- \equiv - \int \mathbf{p}(\sigma_-) d\sigma_- \quad (\text{C4})$$

and boost \mathbf{p} to velocity \mathbf{v}_- followed by a gauge transformation as described in Appendix B. This gives us the final solution

$$\mathbf{P}(\tilde{\sigma}_-) = \frac{\mathbf{p} + (\gamma_- - 1)(\mathbf{v}_- \cdot \mathbf{p})\mathbf{v}_-/\mathbf{v}_-^2 + \gamma_- \mathbf{v}_-}{\gamma_-(1 + \mathbf{v}_- \cdot \mathbf{p})}, \quad (\text{C5})$$

where $\gamma_- = (1 - \mathbf{v}_-^2)^{-1/2}$ and

$$d\tilde{\sigma}_- = \gamma_-(1 + \mathbf{v}_- \cdot \mathbf{p})d\sigma_-. \quad (\text{C6})$$

As in Appendix B, an integration shows that $\tilde{\sigma}_- \in [0, 1/\gamma_-]$ and a coordinate rescaling can be used to bring the interval back to $[0, 1]$.

It is straightforward to check directly that $|\mathbf{P}| = 1$ and also

$$\int \mathbf{P}(\tilde{\sigma}_-) d\tilde{\sigma}_- = 0. \quad (\text{C7})$$

In a similar way we construct $\mathbf{Q}(\tilde{\sigma}_+)$.

This scheme has the advantage that it solves the constraints exactly and explicitly though it still requires integrating $\mathbf{p}(\sigma_-)$ to find \mathbf{v}_- and $\tilde{\sigma}_-$. We have not employed this scheme in the work reported here but have remained with the SP algorithm to facilitate comparison.

-
- [1] A. Vilenkin and E.P.S. Shellard, *Cosmic Strings and Other Topological Defects*, Cambridge Monographs on Mathematical Physics (Cambridge University Press, Cambridge, England, 1994).
 - [2] R.J. Scherrer and W.H. Press, *Phys. Rev. D* **39**, 371 (1989).
 - [3] R.J. Scherrer, J.M. Quashnock, D.N. Spergel, and W.H. Press, *Phys. Rev. D* **42**, 1908 (1990).
 - [4] P. Casper and B. Allen, *Phys. Rev. D* **52**, 4337 (1995).
 - [5] D. Garinkle and T. Vachaspati, *Phys. Rev. D* **36**, 2229 (1987).
 - [6] C.J. Copi, Cosmic String Information, <http://www.phys.cwru.edu/projects/strings/> (2010), source code and documentation are available.
 - [7] D.P. Bennett and F.R. Bouchet, *Phys. Rev. Lett.* **60**, 257 (1988).
 - [8] E.J. Copeland and T.W.B. Kibble, *Phys. Rev. D* **80**, 123523 (2009).
 - [9] E.J. Copeland and N. Turok, *Phys. Lett. B* **173**, 129 (1986).
 - [10] A.L. Chen, D.A. DiCarlo, and S.A. Hotes, *Phys. Rev. D* **37**, 863 (1988).
 - [11] D. DeLaney, K. Engle, and X. Scheick, *Phys. Rev. D* **41**, 1775 (1990).
 - [12] X.A. Siemens and T.W.B. Kibble, *Nucl. Phys. B* **438**, 307 (1995).
 - [13] V. Vanchurin, K. Olum, and A. Vilenkin, *Phys. Rev. D* **72**, 063514 (2005).