

**Recovery of general relativity in massive gravity via the Vainshtein mechanism**E. Babichev,<sup>1,2</sup> C. Deffayet,<sup>3,4</sup> and R. Ziour<sup>3</sup><sup>1</sup>*ASC, Department für Physik, Ludwig-Maximilians-Universität München, Theresienstrasse 37, D-80333, Munich, Germany*<sup>2</sup>*Institute for Nuclear Research, 60th October Anniversary Prospect, 7a, 117312 Moscow, Russia*<sup>3</sup>*APC, UMR 7164 (CNRS-Université Paris 7), 10 rue Alice Domon et Léonie Duquet, 75205 Paris Cedex 13, France*<sup>4</sup>*IAP (Institut d'Astrophysique de Paris), 98bis Boulevard Arago, 75014 Paris, France*

(Received 2 August 2010; published 3 November 2010)

We study in detail static spherically symmetric solutions of nonlinear Pauli-Fierz theory. We obtain a numerical solution with a constant density source. This solution shows a recovery of the corresponding solution of general relativity via the Vainshtein mechanism. This result has already been presented by us in a recent paper, and we give here more detailed information on it as well as on the procedure used to obtain it. We give new analytic insights into this problem, in particular, for what concerns the question of the number of solutions at infinity. We also present a weak-field limit which allows one to capture all the salient features of the numerical solution, including the Vainshtein crossover and the Yukawa decay.

DOI: [10.1103/PhysRevD.82.104008](https://doi.org/10.1103/PhysRevD.82.104008)

PACS numbers: 04.50.-h, 11.27.+d, 98.80.-k

**I. INTRODUCTION**

Motivated by the wish to find consistent large distance modifications of gravity, theories of “massive gravity” have recently attracted some attention (see e.g. [1] for a review). Starting from the unique consistent quadratic theory for a Lorentz invariant massive spin 2, the Pauli-Fierz theory (henceforth PF) [2], it is easy to build a simple nonlinear completion at the price of introducing a non-dynamical external metric. We will consider here such a nonlinear Pauli-Fierz (henceforth NLPF) theory that we will define precisely below. One of the crucial issues to be answered in models of massive gravity is how to recover metrics sufficiently close to the ones obtained in Einstein’s general relativity (henceforth GR) to pass standard tests of the latter theory, while at the same time having significant deviations from GR at large distances. A first major obstacle in this way is related to the so-called van Dam–Veltman–Zakharov (vDVZ) discontinuity [3]. Indeed, the vDVZ discontinuity states that the quadratic Pauli-Fierz theory does not have linearized GR as a limit when the mass of the graviton is sent to zero. In particular, the quadratic Pauli-Fierz theory leads to physical predictions so different from those of GR, such as those for light bending around the Sun, that one can rule out this theory based on the simplest tests of GR in the Solar System. However, soon after the discovery of the vDVZ discontinuity, it was realized that the discontinuity might in fact disappear in nonlinear Pauli-Fierz theories [4]. This is because a careful examination of static spherically symmetric solutions of those theories by Vainshtein showed that the solutions of the linearized nonlinear Pauli-Fierz theories (i.e. those of simple Pauli-Fierz theory) were only valid at distances larger than a distance scale, the Vainshtein radius  $R_V$ , which goes to infinity when the mass of the graviton is sent to zero. On the other hand Vainshtein showed that there existed a well-behaved (as the mass of the graviton is sent to zero) expansion valid at distances

smaller than  $R_V$ , this expansion being defined as an expansion around the Schwarzschild solution of GR. What Vainshtein did not show is the possibility to join together those two expansions as expansions of one single non-singular underlying solution, as underlined, in particular, by Boulware and Deser soon after Vainshtein’s work first appeared [5]. The vDVZ discontinuity and its possible cure *à la* Vainshtein reappeared later in more complicated (and probably more realistic) models. Those include the Dvali-Gabadadze-Porrati (DGP in the following) model [6] and its sequels [7], which have attracted a lot of attention, in particular, as far as application to cosmology is concerned [8–12]; the “degravitation” models [13], the Galileon and k-Mouflage models [14], the recent models of Refs. [15–17], as well as possibly the Hořava-Lifshitz theory [18] (see e.g. [19]).

The search for an everywhere nonsingular asymptotically flat solution featuring the Vainshtein recovery has proven to be a difficult task. For example, Damour *et al.* found, in Ref. [20] by numerical integration of the equations of motion, that singularities always appeared in static spherically symmetric solutions of the kind considered by Vainshtein and concluded that the Vainshtein mechanism was not working. This issue was recently reexamined by us, and in Ref. [21] we reported briefly (disagreeing with the results of Ref. [20]) on the numerical discovery of a static, spherically symmetric, asymptotically flat solution of massive gravity showing for the first time the Vainshtein recovery of GR. The aim of this work is to provide more details on this solution and on the way it has been obtained. Before doing so, let us however underline that we do not call for a realistic use of the kind of NLPF theory presented here (see however [17,22]). Indeed, such theories are believed to suffer from various deadly pathologies, like the “Boulware-Deser ghost” [5] and the related strong coupling. Rather, we use the NLPF as a toy model to investigate the issue of the success or failure of the Vainshtein mechanism.

This paper is organized as follows. In the next, mostly introductory, Sec. II, we first introduce the NLPF we will be looking at (Sec. II A), the ansatz considered (Sec. II B), the Vainshtein mechanism in brief (Sec. II C), and then with more details, together with the so-called decoupling limit (DL) that played a crucial role in the numerical integration of the field equations (Sec. II D). Then (Sec. III), we give some analytic insight on the putative solution, starting first from standard perturbation theory (Sec. III A) yielding, in particular, a large distance expansion, presenting later expansions valid at intermediate and small distances (Secs. III C and III D). In this section we also discuss with some detail the crucial issue of the number of solutions at infinity, arguing that there exist infinitely many such solutions which are not captured by the standard perturbative expansion (Sec. III B). In the next Sec. IV we introduce a new approximation able to capture both the Vainshtein recovery and the large distance Yukawa decay. Section V is devoted to the detailed presentation of our numerical results. We conclude in Sec. VI. Three Appendixes contain more technicalities.

## II. GENERALITIES ABOUT STATIC SPHERICALLY SYMMETRIC SOLUTIONS

### A. Action and covariant equations

We will consider the theory defined by the action

$$S = \int d^4x \sqrt{-g} \left( \frac{M_P^2}{2} R_g + L_g \right) + S_{\text{int}}[f, g]. \quad (1)$$

In the above action,  $R_g$  is the Ricci scalar computed with the metric  $g_{\mu\nu}$ ,  $f_{\mu\nu}$  is a nondynamical metric,  $M_P$  is a mass scale, and  $L_g$  denotes a generic matter Lagrangian with a minimal coupling to the metric  $g$  (and not to the metric  $f$ ), and  $S_{\text{int}}[f, g]$  is an interaction term with nonderivative couplings between the two metrics. Following the notations of Damour *et al.* [23], these interaction terms are in the form

$$\begin{aligned} S_{\text{int}}^{(a)} &= -\frac{1}{8} m^2 M_P^2 \int d^4x \mathcal{V}^{(a)}(g, f) \\ &\equiv -\frac{1}{8} m^2 M_P^2 \int d^4x \sqrt{-g} V^{(a)}(\mathbf{g}^{-1} \mathbf{f}) \end{aligned} \quad (2)$$

with  $\mathcal{V}^{(a)}(g, f) \equiv \sqrt{-g} V^{(a)}(\mathbf{g}^{-1} \mathbf{f})$  a suitable ‘‘potential’’ density. There is much freedom in the choice of these potentials. In this paper we will mostly concentrate on the following term (henceforth called AGS), considered, in particular, by Arkani-Hamed *et al.* in Ref. [24],

$$S_{\text{int}} = -\frac{1}{8} m^2 M_P^2 \int d^4x \sqrt{-g} H_{\mu\nu} H_{\sigma\tau} (g^{\mu\sigma} g^{\nu\tau} - g^{\mu\nu} g^{\sigma\tau}), \quad (3)$$

where  $g^{\mu\nu}$  denotes the inverse of the metric  $g_{\mu\nu}$ , and  $H_{\mu\nu}$  is defined by

$$H_{\mu\nu} = g_{\mu\nu} - f_{\mu\nu}.$$

We will also comment on other forms for the potential. In the equation above,  $m$  is the graviton mass and  $M_P$  the reduced Planck mass, given by

$$M_P^{-2} = 8\pi G_N,$$

in terms of the Newton constant  $G_N$ . The theory described by the action (1)–(3) is invariant under common diffeomorphisms which transform the metric as

$$\begin{aligned} g_{\mu\nu}(x) &= \partial_\mu x'^\sigma(x) \partial_\nu x'^\tau(x) g'_{\sigma\tau}(x'(x)), \\ f_{\mu\nu}(x) &= \partial_\mu x'^\sigma(x) \partial_\nu x'^\tau(x) f'_{\sigma\tau}(x'(x)). \end{aligned}$$

A crucial property of the potential (3) is that when  $g$  is expanded to second order around the canonical Minkowski metric  $\eta_{\mu\nu}$  as  $g_{\mu\nu} = \eta_{\mu\nu} + h_{\mu\nu}$  and  $f$  has the canonical Minkowski form  $\eta_{\mu\nu}$ , the potential at quadratic order for  $h_{\mu\nu}$  takes the Pauli-Fierz form. The equations of motion, derived from action (1), read

$$M_P^2 G_{\mu\nu} = (T_{\mu\nu} + T_{\mu\nu}^g), \quad (4)$$

where  $G_{\mu\nu}$  denotes the Einstein tensor computed with the metric  $g$ ,  $T_{\mu\nu}$  is the energy momentum tensor of matter fields, and  $T_{\mu\nu}^g$  is the effective energy momentum tensor coming from the variation with respect to the metric  $g$  of the interaction term  $S_{\text{int}}$ .  $T_{\mu\nu}^g$  depends nonderivatively on both metrics  $f$  and  $g$  and is defined as usual as

$$T_{\mu\nu}^g(x) = -\frac{2}{\sqrt{-g}} \frac{\delta}{\delta g^{\mu\nu}(x)} S_{\text{int}}(f, g).$$

A simple, but nontrivial, consequence of Eq. (4) is obtained by taking a  $g$ -covariant derivative  $\nabla$  of both sides of the equation; one gets, using the Bianchi identities and the conservation of the matter energy momentum tensor,

$$\nabla^\mu T_{\mu\nu} = 0, \quad (5)$$

the constraint

$$\nabla^\mu T_{\mu\nu}^g = 0 \quad (6)$$

which the effective energy momentum tensor should obey. This equation will turn out to play an important role in the following.

### B. Static spherically symmetric ansatz and boundary conditions

In this paper we study static spherically symmetric solutions of the theory (1)–(3) and we take the nondynamical metric  $f$  to be flat, i.e. to parametrize (possibly only part of) a Minkowski space-time. Note that it is possible to consider  $f$  to be dynamical as well (see e.g. [20,22,23,25]); for simplicity we will not consider this possibility here, and consider the flatness of  $f$  as a prerequisite of our model. Hence, we consider static spherically symmetric configurations using the following ansatz:

$$g_{\mu\nu}dx^\mu dx^\nu = -e^{\nu(R)}dt^2 + e^{\lambda(R)}dR^2 + R^2d\Omega^2,$$

$$f_{\mu\nu}dx^\mu dx^\nu = -dt^2 + \left(1 - \frac{R\mu'(R)}{2}\right)^2 e^{-\mu(R)}dR^2 \quad (7)$$

$$+ e^{-\mu(R)}R^2d\Omega^2,$$

where  $\nu$ ,  $\lambda$ , and  $\mu$  are unknown functions and  $d\Omega^2$  is the metric of a unit 2-sphere

$$d\Omega^2 = d\theta^2 + \sin^2\theta d\varphi^2.$$

This ansatz, where both metrics are diagonal in a common gauge, is not the most general one (see [25]); however, it is the one relevant for the study of the Vainshtein mechanism of massive gravity which is the aim of the present work (see e.g. [20]). It is easy to check that for the ansatz (7), the tensors  $G_{\mu\nu}$  and  $T_{\mu\nu}^g$  are both diagonal (the diagonal character of  $T_{\mu\nu}^g$  follows from the fact that its components are made by taking products of the diagonal matrices  $\mathbf{f}$ ,  $\mathbf{f}^{-1}$ ,  $\mathbf{g}$ , and  $\mathbf{g}^{-1}$ ). The  $tt$  and  $RR$  components of (4) and the nontrivial part of Bianchi identities (6) read

$$e^{\nu-\lambda}\left(\frac{\lambda'}{R} + \frac{1}{R^2}(e^\lambda - 1)\right) = 8\pi G_N(T_{tt}^g + \rho e^\nu), \quad (8a)$$

$$\frac{\nu'}{R} + \frac{1}{R^2}(1 - e^\lambda) = 8\pi G_N(T_{RR}^g + P e^\lambda), \quad (8b)$$

$$-\frac{1}{m^2 M_P^2} \frac{1}{R} \nabla^\mu T_{\mu R}^g = 0, \quad (8c)$$

where the source energy momentum tensor  $T_{\mu\nu}$  is assumed to have the perfect fluid form

$$T_{\mu\nu} = \text{diag}(-\rho, P, P, P),$$

with total mass

$$M \equiv \int_0^{R_\circ} 4\pi R^2 \rho dR.$$

The matter conservation equation reads

$$P' = -\frac{\nu'}{2}(\rho + P). \quad (9)$$

In the integration of the field equations, we will only consider cases where the source has a constant density inside its radius. For the potential (3), we have

$$T_{tt}^g = m^2 M_P^2 f_\nu, \quad T_{RR}^g = m^2 M_P^2 f_R,$$

$$\nabla^\mu T_{\mu R}^g = -m^2 M_P^2 f_g,$$

where we use the following notations [20]:

$$f_t = \frac{e^{-\lambda-2\mu}}{4} \left[ (3e^{\mu+\nu} + e^\mu - 2e^\nu) \left(1 - \frac{R\mu'}{2}\right)^2 + e^\lambda (2e^\mu - e^\nu) - 3e^{\lambda+\mu} (2e^{\mu+\nu} + e^\mu - 2e^\nu) \right], \quad (10a)$$

$$f_R = \frac{e^{-\nu-2\mu}}{4} \left[ (3e^{\mu+\nu} - e^\mu - 2e^\nu) \left(1 - \frac{R\mu'}{2}\right)^2 + e^\lambda (2e^\mu + e^\nu) - 3e^{\lambda+\mu} (-2e^{\mu+\nu} + e^\mu + 2e^\nu) \right], \quad (10b)$$

$$f_g = -\left(1 - \frac{R\mu'}{2}\right) \frac{e^{-\lambda-2\mu-\nu}}{8R} [8(e^\lambda - 1)(3e^{\mu+\nu} - e^\mu - e^\nu) + 2R((3e^{\mu+\nu} - 2e^\nu)(\lambda' + 4\mu' - \nu')$$

$$- e^\mu(\lambda' + 4\mu' + \nu')) - R^2((3e^{\mu+\nu} - 2e^\nu)(\lambda'\mu' - 2\mu'' - \mu'\nu' + (\mu')^2)$$

$$- e^\mu(\lambda'\mu' - 2\mu'' + \mu'\nu' + (\mu')^2) - 2e^\nu(\mu')^2]. \quad (10c)$$

Besides the components (8a) and (8b), the field equation (4) has nontrivial  $\theta\theta$  and  $\varphi\varphi$  components. However, it is easy to check that the set of equations (8) and (9) is equivalent to Eqs. (4) and (5) once the ansatz (7) is chosen.

The system (8) and (9), is a system of ordinary differential equations (ODEs) that requires five boundary conditions [respectively the values of  $\nu$ ,  $\lambda$ ,  $\mu$ ,  $\mu'$  as well as  $P$  at some (finite) boundary] to be integrated along the radial coordinate  $R$ . One is also looking for an asymptotically flat solution such that  $\nu$  and  $\lambda$  must vanish at large  $R$ , and one should also require that the solution is nonsingular in  $R = 0$ . To fulfill this condition and avoid a conical singularity at the origin, we will require that

$$\lambda(R = 0) = 0. \quad (11)$$

The solution found will also be such that

$$\mu'(R = 0) = 0.$$

In the following sections we first recall some properties of the so-called Vainshtein mechanism which, if valid, opens a way to recover solutions of general relativity in massive gravity for small graviton mass. Then we recall some of our previous results on the so-called decoupling limit of the theory considered, which played a crucial role for the integration of the field equations.

### C. Vainshtein mechanism in brief

The (quadratic) Pauli-Fierz theory is known to suffer from the vDVZ discontinuity, i.e. the fact that when one lets the mass  $m$  of the graviton vanish, one does not recover

predictions of general relativity. For example, if one adjusts the parameters (namely the Planck scale) such that the Newton constant agrees with the one measured by some type of Cavendish experiment, then the light bending as predicted by Pauli-Fierz theory (and for a vanishingly small graviton mass) will be 3/4 of the one obtained by linearizing GR [3].<sup>1</sup> One way to see this is to consider solutions of equations of motion (4) which are static and spherically symmetric and which would describe the metric around a spherically symmetric body such as the Sun. To do so, using the ansatz (7) is especially convenient because in this form, the  $g$  metric can be easily compared with the standard Schwarzschild solution. If one tries to find a solution expanding in the Newton constant  $G_N$ , as we recall in the next section, one finds immediately the vDVZ discontinuity appearing in the form of a different ( $m$  independent) absolute value of the coefficients in front of the first nontrivial correction to flat space-time in  $g_{tt}$  and  $g_{RR}$  components (neglecting the Yukawa decay by assuming the Compton wavelength of the graviton is much larger than other distances of interest). Indeed, those corrections are obtained at first order in  $G_N$ , i.e. by linearizing the field equations, in which case the NLPF and PF theories are equivalent by definition. However, as first noticed by Vainshtein [4], the computation of the next order correction in the NLPF shows that the first order approximation ceases to be valid at distances to the source smaller than a composite scale, the Vainshtein radius defined by

$$R_V = (m^{-4}R_S)^{1/5},$$

where  $R_S$  is the Schwarzschild radius of the source. This Vainshtein radius obviously diverges when one lets  $m$  go to zero. It is also much larger than the Solar System size for a massive graviton with a Compton wavelength of the order of the Hubble radius and  $R_S$  taken to be the Schwarzschild radius of the Sun. Hence, one cannot rule out massive gravity based on Solar System observations and on the results of the original works on the vDVZ discontinuity [3]. Vainshtein also showed that an expansion around the standard Schwarzschild solution (defined as an expansion in the graviton mass  $m$ ) can be obtained (as we recall in the next section). This expansion is well behaved when the mass of the graviton is sent to zero, opening the possibility of a recovery of GR solution at small distances  $R$  of the

<sup>1</sup>The fact it is smaller is easy to understand: the essential difference between Pauli-Fierz theory and linearized GR comes from an extra propagating scalar mode present in the massive theory. This mode exerts an extra attraction in the massive case compared to the massless case. Hence, if one wants measurements of the force exerted between nonrelativistic masses to agree, the coupling constant of the massive theory should be smaller than that of the massless theory. But light bending is blind to the scalar sector—because the light energy momentum tensor is traceless. Hence, provided the two theories agree on the force between nonrelativistic probes, the massive theory would predict a smaller light bending than the massless one.

source. Indeed, the domain of validity of this second expansion was shown to be  $R \ll R_V$ . Moreover, the corrections found by Vainshtein to the Schwarzschild solution are nonanalytic in the Newton constant which can explain the failure of the attempt to obtain an everywhere valid solution expanding in the Newton constant.

#### D. The decoupling limit and the Vainshtein mechanism

The Vainshtein mechanism can be enlightened taking the so-called decoupling limit (DL) of the theory considered. Indeed this limit gives the scalings of the different metric coefficients expected to hold (for an interesting range of distances to be specified below), should the Vainshtein mechanism be valid. The DL has been investigated by us in a previous work [26], and those investigations turned out also to be crucial to obtain the solutions of the full system presented here and in Ref. [21]. Hence, it is the purpose of this section to briefly present the salient features of the DL and its application to the study of the Vainshtein mechanism. Instead of the way it was originally introduced [24], we will define here the DL directly in the equations of motion [26]. Indeed, it can be shown that the full system of equations (8) reduces to a much simpler one in the DL defined as follows:

$$M_P \rightarrow \infty, \quad m \rightarrow 0, \quad \Lambda \sim \text{const},$$

$$T_{\mu\nu}/M_P \sim \text{const},$$

where  $\Lambda$  is an energy scale defined by

$$\Lambda = (m^4 M_P)^{1/5}.$$

This scale is associated with the strongest self-interaction of the model (for more details see [24,26]). Taking the DL, Eqs. (8) become

$$\frac{\lambda'}{R} + \frac{\lambda}{R^2} = -\frac{m^2}{2}(3\mu + R\mu') + \frac{\rho}{M_P^2}, \quad (12a)$$

$$\frac{\nu'}{R} - \frac{\lambda}{R^2} = m^2\mu, \quad (12b)$$

$$\frac{\lambda}{R^2} = \frac{\nu'}{2R} + Q(\mu), \quad (12c)$$

where  $Q$  is the only nonlinear part left over in the DL. As seen above, the DL amounts to linearizing  $G_{tt}^g$  and  $G_{RR}^g$  (as well as the sources), keeping the part of  $f_t$  and  $f_R$  linear in  $\mu$ , linearizing  $f_g$  and keeping also the piece of  $f_g$  which is quadratic in  $\mu$ . As appears above, the linearization of  $f_g$  does not contain any pieces depending on  $\mu$ .  $Q$  is the only part of the pieces left over in the DL which depends on the choice of potential  $V^{(a)}$  (provided the theory is in the NLPF category). For AGS potential,  $Q$  is given by

$$Q = -\frac{1}{2}\left(\frac{\mu'^2}{2} + \mu\mu'' + \frac{4\mu\mu'}{R}\right). \quad (13)$$

Note that we have neglected the pressure  $P$  in the above equations (12), since this is a direct consequence of taking



the decoupling limit in the matter conservation equation (9). From the three equations (12) a single equation on  $\mu$  can be extracted, which after integration and matching of the solution to the source, reads

$$2Q(\mu) + \frac{3}{2}m^2\mu = \frac{R_S}{R^3} \quad (\text{outside source}), \quad (14a)$$

$$2Q(\mu) + \frac{3}{2}m^2\mu = \frac{R_S}{R_\circ^3} \quad (\text{inside source}), \quad (14b)$$

where  $R_\circ$  is the radius of the source and where we assumed that the density of the source is constant. Note the right-hand side (rhs) of (14b) is constant, in contrast to that of Eq. (14a).

The essence of the Vainshtein mechanism, as well as the related scalings of the solutions can be understood from Eq. (14a), which allows one to obtain  $\mu$ , as well as from Eqs. (12a) and (12b) which enables one to get  $\lambda$  and  $\nu$  once  $\mu$  is known. Indeed, one can guess the existence of two different regimes for Eq. (14a), namely, the linear and the nonlinear regimes. In the linear regime, the nonlinear part  $Q$  can first be neglected. This regime is valid for  $R \gg R_V$ , and the asymptotic solutions,  $\bar{\mu}_\infty$  and  $\bar{\lambda}_\infty$ , read

$$\begin{aligned} \bar{\mu}_\infty &\simeq \frac{2}{3} \frac{R_S}{m^2 R^3} \left[ 1 + \mathcal{O}\left(\frac{R_V^5}{R^5}\right) \right], \\ \bar{\lambda}_\infty &\simeq -\frac{\bar{\nu}_\infty}{2} \simeq \frac{2}{3} \frac{R_S}{R} \left[ 1 + \mathcal{O}\left(\frac{R_V^5}{R^5}\right) \right]. \end{aligned} \quad (15)$$

From (15), it follows that  $\bar{\nu}_\infty \simeq -2\bar{\lambda}_\infty$ , which is a manifestation of the vDVZ discontinuity, and results, e.g., in the mentioned difference between light bending in GR and in PF theory. In the nonlinear regime, the linear term on the left-hand side (lhs) of (14) is neglected compared to the nonlinear  $Q$  term. In this regime (valid for  $R \ll R_V$ ) two different scalings can be identified.

The first scaling, henceforth called the *Vainshtein scaling*, is found assuming that the lhs cancels the rhs for the leading term in the expansion of  $\mu$ . That is, in vacuum, one assumes that the leading term  $\bar{\mu}_0$  is such that

$$Q(\bar{\mu}_0) \simeq R_S/R^3. \quad (16)$$

For  $Q$  given by (13), the leading behaviors of the metric function read

$$\begin{aligned} \bar{\mu}_V &\simeq \sqrt{\frac{8R_S}{9R}} \left[ 1 + \mathcal{O}\left(\frac{R}{R_V}\right)^{5/2} \right], \\ \bar{\nu}_V &\simeq -\frac{R_S}{R} \left[ 1 + \mathcal{O}\left(\frac{R}{R_V}\right)^{5/2} \right], \\ \bar{\lambda}_V &\simeq \frac{R_S}{R} \left[ 1 + \mathcal{O}\left(\frac{R}{R_V}\right)^{5/2} \right]. \end{aligned} \quad (17)$$

In this regime, one recovers at dominant order the results of linearized general relativity, namely  $\bar{\nu}_V \sim -\bar{\lambda}_V$ . Inside the source, following a similar logic, one can look for a solution to Eq. (14b)  $\bar{\mu}_{\circ,V}$ , such that the leading behavior gives  $Q(\bar{\mu}_{\circ,V}) \sim R_S/R_\circ^3$ . We find

$$\bar{\mu}_{\circ,V} = \sqrt{\frac{R_S}{R_\circ}} \left( \tilde{B}_V - \frac{1}{10\tilde{B}_V} (R/R_\circ)^2 \right) + \frac{3}{20} (mR)^2 + \dots, \quad (18a)$$

$$\bar{\nu}_{\circ,V} = \bar{\nu}_0 + \frac{R_S R^2}{2R_\circ^3} + \frac{\tilde{B}_V}{4} \sqrt{\frac{R_S}{R_\circ}} (mR)^2 + \dots, \quad (18b)$$

$$\bar{\lambda}_{\circ,V} = \frac{R_S R^2}{R_\circ^3} - \frac{\tilde{B}_V}{2} \sqrt{\frac{R_S}{R_\circ}} (mR)^2 + \dots, \quad (18c)$$

where  $\tilde{B}_V$  and  $\bar{\nu}_0$  are constants of integration, to be fixed by matching the solutions outside the source, Eq. (17), and inside the source.

A second scaling is possible, leading to the restoring of general relativity in the nonlinear regime [26]. This is what we called the *Q scaling*, whose leading behavior corresponds to a zero mode of the nonlinear operator  $Q$ . This zero mode is defined such that  $Q(\bar{\mu}) \sim 0$ , in which case the rhs term  $R_S/R^3$  compensating the nonlinear piece of (14a) is subdominant. The *Q*-scaling solution of (14a) reads in vacuum,

$$\bar{\mu}_Q = (mR_V)^2 \left\{ A_Q \left( \frac{R_V}{R} \right)^2 + \left[ B_Q - \frac{1}{3A_Q} \ln\left( \frac{R}{R_V} \right) \right] \frac{R}{R_V} \right\}, \quad (19a)$$

$$\bar{\nu}_Q = -\frac{R_S}{R} + \frac{1}{2} A_Q (R_V m)^4 \ln(R/R_V), \quad (19b)$$

$$\bar{\lambda}_Q = \frac{R_S}{R} - \frac{A_Q}{2} (R_V m)^4. \quad (19c)$$

From (19) it is clearly seen that general relativity is restored for small  $R$ . Note that the *Q* solution (19) is a richer family compared to the Vainshtein scaling (17), since it contains two constants,  $A_Q$  and  $B_Q$ . Inside a source, it is not hard to guess that the leading behavior of the *Q* scaling is also given by the *Q* scaling (19) outside the source

$$\bar{\mu}_{\circ,Q} = (mR_V)^2 \tilde{A}_Q \left( \frac{R_V}{R} \right)^2, \quad \bar{\nu}_{\circ,Q} = -\frac{R_S}{R}, \quad \bar{\lambda}_{\circ,Q} = \frac{R_S}{R}, \quad (20)$$

since the leading term of *Q* scaling does not touch the rhs of Eqs. (14). Note, however, that the subdominant terms are different for the *Q* scalings inside and outside the source. Note also that this scaling leads to a curvature singularity at  $R = 0$  and does not obey the boundary condition (11).

In Ref. [26], we integrated numerically the DL system (12). We found that solutions interpolating between the above described asymptotic regimes indeed exist. We were able to identify the following two families of solutions. For a given choice of parameters  $R_S$ ,  $m$ , and  $R_\circ$ , there is one solution  $\{\bar{\mu}_{N,V}, \bar{\nu}_{N,V}, \bar{\lambda}_{N,V}\}$  which interpolates between the ‘‘Vainshtein’’ asymptotic scalings given respectively by (15), (17), and (18), see Table I. Note that the Vainshtein solution is unique for fixed  $R_S$ ,  $m$ , and  $R_\circ$ , since the free constant  $\tilde{B}_V$  in (18) is fixed by the requirement of the asymptotic flatness (15) [26]. Apart from this Vainshtein solution, we have found a one-parameter family of solutions  $\{\bar{\mu}_{N,Q}, \bar{\nu}_{N,Q}, \bar{\lambda}_{N,Q}\}$  which interpolates between the

TABLE I. Solution containing the Vainshtein scaling.

$R$	$R \ll R_\odot$	$R_\odot \ll R \ll R_V$	$R \gg R_V$
$\bar{\mu}_{N,V}$	const	$\sqrt{8R_S/(9R)}$	$2R_S/(3m^2R^3)$
$\bar{\nu}_{N,V}$	const	$-R_S/R$	$-4R_S/(3R)$
$\bar{\lambda}_{N,V}$	$\text{const} \times R^2$	$R_S/R$	$2R_S/(3R)$

solutions given by (15), and the  $Q$  scalings (19) and (20), see Table II. In contrast to the Vainshtein solution, this solution is not uniquely fixed by the asymptotic flatness. There is a freedom in the choice of  $A_Q$  in (19), so that the family of solutions corresponds to a curve in the  $(A_Q, B_Q)$  plane of solutions (19).

Although the solutions we found in DL are solutions of a simplified version of the full system, they should be valid in some range of distances  $R$  for the full system. Indeed, let us estimate the validity of DL solutions (see also [26]). To do this we compare the terms left over in the DL system of equations (12) with the terms we dropped from the original system (8). First of all, keeping only linear terms in  $\lambda$  and  $\nu$  on the rhs of (8) requires  $R \gg R_S$  outside the source, as in GR. The second condition comes from comparison of the linear terms containing  $\mu$  with those containing  $\lambda$  and  $\nu$  in the expansion of the rhs of (8a) and (8b). In the DL, the linear terms containing  $\mu$  are neglected, and one can see that those terms can only be neglected at distances  $R \ll m^{-1}$  for the full system. That is, those first two conditions state that the DL can at best recover linearized GR and will not capture the expected Yukawa decay of the solution. The third condition arises from the neglecting of terms in the nonlinear regime. Indeed, the quadratic terms in  $\mu$ ,  $\nu$ , and  $\lambda$  on the rhs of the expansion of (8a) and (8b) must be subdominant for DL to be valid for the full system. Similarly, the cubic terms on the rhs of (8c) should be small compared to the quadratic ones, left over in DL. This gives nothing new for the Vainshtein scaling, namely,  $R \gg R_S$ . However, this leads to a more rigid constraint for the  $Q$  scaling

$$R \gg R_V(mR_S)^{1/5}. \quad (21)$$

To summarize, the DL solution, having the Vainshtein asymptotic at small  $R$  and the linear asymptotic solution for large  $R$ , is expected to be a good approximation for the solution of the full system in the range,

$$R_S \ll R \ll m^{-1},$$

while the DL solution, approximating the  $Q$  scaling at small  $R$  and the linear asymptotic solution for large  $R$ ,

TABLE II. Solution containing the  $Q$  scaling.

$R$	$R \ll R_\odot$	$R_\odot \ll R \ll R_V$	$R \gg R_V$
$\bar{\mu}_{N,Q}$	$m^2 R_V^4 \bar{A}_Q \times R^{-2}$	$m^2 R_V^4 A_Q \times R^{-2}$	$2R_S/(3m^2R^3)$
$\bar{\nu}_{N,Q}$	$-R_S/R$	$-R_S/R$	$-4R_S/(3R)$
$\bar{\lambda}_{N,Q}$	$R_S/R$	$R_S/R$	$2R_S/(3R)$

should be a good approximation for the solution of the full system in the range,

$$R_V(mR_S)^{1/5} \ll R \ll m^{-1}.$$

If a not too compact source of radius  $R_\odot \gg R_S$  is included into the picture, one can check using the dominant behavior (18) that the Vainshtein solution even remains a good approximation of a full solution until the origin. This will later be confirmed by our numerical integration. On the other hand, even when a large source is included, the DL  $Q$ -scaling solution stops being relevant for the full system at the distance  $R = R_V(mR_S)^{1/5}$ .

In the following, we will look for a solution of the full system (8), and we will use our understanding of the DL as a guideline. As we will report, our numerical attempts to find a solution of the full system with the  $Q$  scaling asymptotic only lead to singular solutions, and accordingly, we will mostly focus on the extension of the Vainshtein scaling to the full system, rather than on the  $Q$ -scaling solution.

### III. INTEGRATION OF THE FIELD EQUATIONS: ANALYTIC INSIGHTS

Before turning to present the numerical integration of the full system of Eqs. (8) and (9), it is convenient to present below some results one can obtain analytically. Those will mostly be given in the form of series expansions. We have no proof of the convergence of those series, and we in fact believe (and will give some arguments below) that those series are divergent, being then asymptotic expansions. However, we found those series useful for three main reasons: First they allowed us to set boundary or ‘‘initial’’ conditions for the numerical integration with the required precision; second, they were also useful to check that the numerical solutions we found were not singular at the origin of the radial coordinate system; and third, because they help investigating the uniqueness of the solutions at infinity.<sup>2</sup>

Let us first outline the organization of this section as well as its main results, allowing a reader not interested in technical details to skip all but this introduction. The first Sec. III A explains with some detail the solution of standard perturbation theory. This perturbation theory is defined as an expansion in a small parameter  $\epsilon$  where at each order, all the metric functions  $\lambda$ ,  $\nu$ , and  $\mu$  are assumed to be proportional to  $\epsilon$  raised to the same power. The

<sup>2</sup>Similar series expansions were used in Ref. [20] to argue that the system we investigate here has a unique asymptotically flat solution. It was then argued on the basis of this, that it was *a priori* unlikely that one could match this unique solution to a solution obtained *à la* Vainshtein, expanding around the Schwarzschild solution, as will be explained in Sec. III C. As we will see our results do not support this conclusion on the uniqueness of the asymptotically flat solution, hence removing the argument given in Ref. [20] to explain the singularities found by the numerical integration done in this last reference.

parameter  $\epsilon$  is expected to be proportional to the mass of the source (or to its density). Solutions of the perturbation theory out of the source are obtained at each order in  $\epsilon$  in the form of asymptotic series both for small distances and for large distances. Those series are expressed in terms of the dimensionless quantity  $z = Rm$ . Hence, we obtain a double series expansion (into  $\epsilon$  and  $z$ ). The “small distance” expansion, which is an expansion valid at small  $z$  (at a fixed order  $\epsilon$ ), is similar to the regime (15) of the DL, and is found (by comparing successive orders in  $\epsilon$ ) to be no longer valid at distances smaller than the Vainshtein radius in agreement with Vainshtein’s computation. This small distance series expansion depends on two unknown integration constants which have to be fixed by matching to the source. On the other hand, the “large distance” (large  $z$ ) expansion is uniquely fixed and has no free parameter. However, we argue in Sec. III B that this is not enough to define a unique solution at infinity, namely, because we show that a nonperturbative correction can always be added to the unique solution of perturbation theory we identified. This is argued to be related to the fact the series presented are not convergent series, in analogy with what we have shown for the decoupling limit [26]. Then, we present expansions valid inside the Vainshtein radius first outside the source (Sec. III C), then inside it (Sec. III D). The expansion inside the Vainshtein radius (Sec. III C), but outside of the source, is obtained expanding the functions around GR, in a small parameter  $\epsilon'$  which is by definition proportional to  $m^2$ . At each order in  $\epsilon'$ , one can then expand the functions in powers of  $R/R_S$  ( $R_S$  being the Schwarzschild radius of the source). One thus obtains again a double expansion. This expansion allows one to compute the first nontrivial corrections to the GR solution. Those corrections will then later be compared to the ones found by the numerics. One key property of the Vainshtein mechanism is related to the existence of those double expansions, to the fact that the associated series appear not to be convergent allowing the addition of “nonperturbative corrections” such as the ones mentioned above, and to the fact, as discovered by Vainshtein, that the solution in the “below  $R_V$ ” regime features corrections to GR which are nonanalytic in the  $\epsilon$  parameter of the standard perturbation theory. Last (Sec. III D), we obtain a new expansion valid inside the source that will be later compared to the numerical solution, and also used by us to study the regularity of the solution at the origin.

### A. Standard perturbation theory in vacuum

Looking for a solution of the system (8) and (9) far from a source, the first obvious way is to look for a scheme where the functions  $\lambda$ ,  $\nu$ , and  $\mu$  can then be expanded as

$$\begin{aligned}\lambda &= \lambda_0 + \lambda_1 + \dots, \\ \nu &= \nu_0 + \nu_1 + \dots, \\ \mu &= \mu_0 + \mu_1 + \dots,\end{aligned}\tag{22}$$

where  $\lambda_i$ ,  $\nu_i$ , and  $\mu_i$  are assumed to be proportional to  $\epsilon^{i+1}$ ,  $\epsilon$  being a small parameter, and such that  $\lambda_{i+1} \ll \lambda_i$ ,  $\nu_{i+1} \ll \nu_i$ , and  $\mu_{i+1} \ll \mu_i$ , at least for some range of distances. This is what we will call here standard perturbation theory. Before going further, let us first introduce a new dimensionless radial coordinate  $z$  defined by  $z = Rm$ . In this unit,  $R = m^{-1}$  corresponds to  $z = 1$ , and  $R = R_V$  translates to  $z = a$ , where  $a$  is defined as

$$a \equiv R_V m.\tag{23}$$

Using this variable, the lowest order  $\lambda_0$ ,  $\nu_0$ , and  $\mu_0$  are obtained solving the system (8a)–(8c) linearized which reads

$$\begin{aligned}\frac{\lambda'}{z} + \frac{\lambda}{z^2} + \frac{1}{2}(\lambda + 3\mu + z\mu') &= 0, \\ \frac{\nu'}{z} - \frac{\lambda}{z^2} - \frac{1}{2}(\nu + 2\mu) &= 0, \quad \nu' - \frac{2\lambda}{z} = 0.\end{aligned}$$

The asymptotically flat solution of the above linear system can be found analytically; it is given by [20,27]

$$\nu_0 = -C \frac{4b}{3z} e^{-z},\tag{24a}$$

$$\lambda_0 = C \frac{2b}{3} \left(1 + \frac{1}{z}\right) e^{-z},\tag{24b}$$

$$\mu_0 = C \frac{2b}{3z} \left(1 + \frac{1}{z} + \frac{1}{z^2}\right) e^{-z},\tag{24c}$$

where

$$b \equiv a^5 = (R_V m)^5 = R_S m,$$

and  $C$  is a dimensionless integration constant. The value of the constant of integration  $C$  can in principle be determined by integrating the system inward and matching to the source, as can be done in a similar case in GR. There, indeed, an integration constant with the same status arises, when one integrates the linearized Einstein equations in the vacuum around a static spherically symmetric source. This constant of integration is then seen to give the Schwarzschild mass, once matched to the source. Here, an alternative way to proceed is to use our knowledge of the decoupling limit. Indeed, requiring the solution to be very close to the decoupling limit solution (15) in the range  $R_V \ll R \ll m^{-1}$ , i.e.,  $a \ll z \ll 1$ , fixes the constant to be

$$C = 1.\tag{25}$$

In Sec. VB, we will show that this choice of constant of integration is also in agreement with the numerical integration. Accordingly, from here on and until the end of this article, we will assume (except at some scattered places)  $C = 1$ . Note however that when dealing with the equations in vacuum, the value of this constant is irrelevant. Having fixed  $C$  one notices that the solution (24) features the expected behavior for the gravitational potential of a massive graviton: it has a Yukawa decay and the “Newtonian” part  $\nu_0$  is 4/3 larger than the GR result (in the range

of distances  $R \ll m^{-1}$ ) while the radial part  $\lambda_0$  is  $2/3$  smaller, a structure related to the vDVZ discontinuity. Our expansion parameter  $\epsilon$  can then be defined as  $\epsilon \equiv \mathcal{C} \times b$ ; note that it is indeed much smaller than 1 for usual astrophysical sources and a graviton of cosmological size Compton length.

To go beyond the lowest order (24), the first step consists of rewriting the system of equations (8), separating explicitly the linear terms from the nonlinear ones:

$$\begin{aligned} \frac{\lambda'}{z} + \frac{\lambda}{z^2} + \frac{1}{2}(\lambda + 3\mu + z\mu') &= f_{t,\geq 2} - G_{tt,\geq 2}, \\ \frac{\nu'}{z} - \frac{\lambda}{z^2} - \frac{1}{2}(\nu + 2\mu) &= f_{R,\geq 2} - G_{RR,\geq 2}, \\ \frac{\lambda}{z^2} - \frac{\nu'}{2z} &= \frac{1}{z} f_{g,\geq 2}. \end{aligned} \quad (26)$$

In the above system, we have used the expression (10) for the functions  $f_i$  and the  $f_{i,\geq 2}$ , and  $G_{ii,\geq 2}$  stand for the nonlinear part of the functions  $f_i$  and of the components of the Einstein's tensor in spherical coordinates. At each order  $n$ , the functions  $\lambda_n$ ,  $\nu_n$ , and  $\mu_n$  (with  $n \geq 1$ ) can be obtained by solving the above system where the left-hand side consists of a linear operator acting on the unknown  $\lambda_n$ ,  $\nu_n$ , and  $\mu_n$ , while the right-hand side is obtained keeping the order  $\epsilon^{n+1}$  and consists of nonlinear expressions depending on the  $\lambda_i$ ,  $\nu_i$ , and  $\mu_i$  (with  $i < n$ ). As will be shown in detail in Appendix A, the functions  $\lambda_n$ ,  $\nu_n$ , and  $\mu_n$  (with  $n \geq 1$ ) can be expressed as asymptotic (and formal) expansions in large  $z$  (large distance) or small  $z$  (small distance) regimes. We now discuss the main properties of those expansions, starting, for reasons we will explain, by the small distance expansion.

### 1. Small distance expansion

An expansion, asymptotic at small  $z$ , can be obtained in the form

$$\begin{aligned} \mu_n &= z^2 \left( \epsilon \frac{e^{-z}}{z^5} \right)^{n+1} \sum_{i=0}^{\infty} \sum_{k=0}^n \mu_{n,i,k} z^i (z^5 \log z)^{n-k}, \\ \lambda_n &= z^4 \left( \epsilon \frac{e^{-z}}{z^5} \right)^{n+1} \sum_{i=0}^{\infty} \sum_{k=0}^n \lambda_{n,i,k} z^i (z^5 \log z)^{n-k}, \\ \nu_n &= z^4 \left( \epsilon \frac{e^{-z}}{z^5} \right)^{n+1} \sum_{i=0}^{\infty} \sum_{k=0}^n \nu_{n,i,k} z^i (z^5 \log z)^{n-k}. \end{aligned} \quad (27)$$

Note that the form of this expansion was sketched in Ref. [20], with however some important differences: the corresponding expansion given there is stopping at a finite value of  $i$  at each order  $n$  and did not contain any logarithms. As shown in Appendix A, it is however not possible to obtain such a simplified expansion. As detailed in the same Appendix, two important properties of the expansion (27) are as follows. First, when solving for the order  $n$ , the coefficients  $\mu_{n,i,k}$ ,  $\lambda_{n,i,k}$ , and  $\nu_{n,i,k}$  are determined only after

choosing two arbitrary constants. This is because the linear operator appearing on the left-hand side of Eq. (26) has two independent zero modes [among which is the zero order solution (24)] which can be expanded formally in the same way as above.<sup>3</sup> One way to fix those constants would be to enforce at small  $z$  the hierarchy  $\lambda_{i+1} \ll \lambda_i$ ,  $\nu_{i+1} \ll \nu_i$ , and  $\mu_{i+1} \ll \mu_i$ . We have in fact checked that it is possible to do so at order  $n = 1$ . However, it is hard to go further because it is far from obvious that the expansions considered are convergent. The question of convergence is actually quite hard to address, since it requires one to know the behavior of the coefficients  $\lambda_{n,i,k}$ ,  $\nu_{n,i,k}$ , and  $\mu_{n,i,k}$ . This can be achieved quite easily for a subset of those coefficients. Indeed, let us, at each order  $n$  retain the dominant term in the expansion (27) at small  $z$  (i.e. for  $R \ll m^{-1}$ ). We obtain, for  $z \ll 1$  the following expansion for  $\mu$ :

$$\mu \sim \sum_{n=0}^{\infty} \epsilon^{n+1} \frac{\mu_{n,0,n}}{z^{3+5n}}. \quad (28)$$

Equivalently, this is what is obtained in the limit  $m \rightarrow 0$ . In fact, and this is the second important property we would like to mention here, as shown in Appendix A, the expansion (28) is exactly the one obtained in the DL [cf. Ref. [26], Eq. (4.34)]. The coefficient  $\mu_{n,0,n}$  can be easily computed in the DL, and it has been shown in [26] that this series diverges for any  $z$ , providing only an asymptotic expansion of the solution. One can even go further: using the exact values of the coefficients  $\mu_{n,0,n}$  obtained in the DL, one can show that the subseries of (27) given by

$$\sum_{n=0}^{\infty} \epsilon^{n+1} e^{-(n+1)z} \frac{\mu_{n,0,n}}{z^{3+5n}}, \quad (29)$$

diverges for any  $z$ , despite the presence of the exponential damping  $e^{-(n+1)z}$ . Of course, the fact that this subseries of (27) diverges does not lead to any conclusion by itself besides the fact that the expansion (27) cannot be absolutely convergent (if it were the case, any subseries would converge and one could rearrange at will the order of summation). In our case, there is no clear prescription for the order of summation, and the divergence of the subseries (29) does not mean that the full series diverges. Still, we see this divergence as a clue that the convergence of the full series is far from guaranteed, and that any conclusion based on these expansions should be taken with great care. For instance, one may question the approach followed in [20], Sec. IV.B, which consists of building the solution of massive gravity far from the source and arguing on its uniqueness, starting from an expansion of the form (27), which is implicitly assumed to be convergent. Moreover, as also happens in the DL, even if the series expansion (27) turns out to be uniquely defined (once the constants we mentioned have

<sup>3</sup>The presence of logarithms in this expansion is also related to those zero modes.



been fixed e.g. by the procedure outlined above), this does not mean that it defines a unique vacuum solution because it could be, if the series is not convergent, that a family of solutions all have the same asymptotic series expansion. We will in fact come back later to this important issue.

Though, let us state that from a numerical point of view expansions (27) and (28) are very useful. Indeed, if one keeps only the relevant first orders, these expansions appear to be excellent approximations of the solutions and can be used to set properly boundary or initial conditions in the numerical integration.

Moreover, by comparing two successive orders in  $\epsilon$ , it is easily seen that for  $R \ll m^{-1}$ , the expansion (27) ceases to be valid when  $R$  becomes smaller than the Vainshtein radius (i.e. whenever  $\epsilon z^{-5}$  becomes of order 1).

## 2. Large distance expansion

In a similar way as above, an asymptotic expansion at large  $z$  can be found at each order  $n$  in the  $\epsilon$  expansion. It is derived in Appendix A, and reads

$$\begin{aligned}\mu_n &= \epsilon^{n+1} e^{-(n+1)z} \sum_{i=-\infty}^{i=0} \mu_{n,i} z^i, \\ \lambda_n &= \epsilon^{n+1} e^{-(n+1)z} \sum_{i=-\infty}^{i=0} \lambda_{n,i} z^i, \\ \nu_n &= \epsilon^{n+1} e^{-(n+1)z} \sum_{i=-\infty}^{i=0} \nu_{n,i} z^i.\end{aligned}\quad (30)$$

One can verify easily that the zeroth order exact solutions (24) can be expressed as above. In contrast to the expansion (27), there is here no need for logarithms nor to fix ‘‘constants of integration.’’ As such, this could indicate in a clearer way than in the previous case that, as stated in Ref. [20], perturbation theory defines a unique solution. This is however again not obvious because this relies crucially on the fact that the above series converges. In fact, we have computed numerically the first coefficients of the first nontrivial order  $n = 1$  and check that the corresponding series seems to have a vanishing radius of convergence.<sup>4</sup> In the next section we try to address the issue of the uniqueness of the solution at infinity in a different way. This analysis again indicates, in agreement with our numerical investigations, that the solution at infinity is not unique.

### B. Investigations of the number of solutions at infinity

As underlined in the previous section, it is hard to conclude anything on the uniqueness of the solution at infinity from the existence of series expansions, which

<sup>4</sup>One finds numerically that  $|\mu_{1,i}/\mu_{1,i+1}| \propto |i|$  indicating a vanishing radius of absolute convergence.

can be divergent. Here we would like to provide arguments showing that there are in fact infinitely many asymptotically flat solutions at infinity which have the same dominant behavior at large distance once a constant, corresponding to the ‘‘mass’’ of the source in the dominant terms has been chosen. This nonuniqueness appears in the form of a nonperturbative correction that can be added at will to the unique solution of the perturbation theory.<sup>5</sup> This is in striking contrast to the general relativity case. Indeed, in GR, once the asymptotic behavior has been chosen through<sup>6</sup>  $\nu = -R_S/R$  and  $\lambda = R_S/R$ , there is a unique solution which follows this asymptotic behavior at infinity: the Schwarzschild solution which reads  $\nu = \nu_{GR} \equiv \ln(1 - R_S/R)$  and  $\lambda = \lambda_{GR} \equiv -\nu_{GR}$ . In NLPF gravity—at least for potentials of the type (3)—we shall see that the asymptotic behavior is not enough to uniquely determine the solution.

To prove the existence of these infinitely many solutions, let us first assume that there exists a solution of the full nonlinear system of equations (8),  $\bar{\sigma} \equiv \{\bar{\lambda}, \bar{\nu}, \bar{\mu}\}$ , such that

$$\bar{\sigma} \rightarrow \sigma_0 \quad \text{for } z \rightarrow \infty,$$

where  $\sigma_0 \equiv \{\lambda_0, \nu_0, \mu_0\}$  as defined in Eq. (24), and we imply we have fixed all the necessary integration constants such that  $\bar{\sigma}$  and  $\sigma_0$  are fully specified. We then want to show there is an infinite number of solutions in vacuum having the same asymptotic at the infinity. We will look for these solutions  $\sigma$  as small perturbations of the known solution  $\bar{\sigma}$ :

$$\sigma = \bar{\sigma} + \delta\sigma,$$

where  $\delta\sigma \equiv \{\delta\lambda, \delta\nu, \delta\mu\}$  and we will further assume that

$$|\delta\sigma| \ll |\bar{\sigma}|, \quad |\delta\sigma'| \ll |\bar{\sigma}'|, \quad |\delta\sigma''| \ll |\bar{\sigma}''|. \quad (31)$$

Taking into account (31), we linearize the system of equations (8), for  $z \rightarrow \infty$ , around the solution  $\bar{\sigma}$  to obtain the following system of equations on  $\delta\lambda$ ,  $\delta\nu$ , and  $\delta\mu$ :

<sup>5</sup>Note that we found a similar result in the DL case [26], as recalled in Sec. IID. Indeed, in the DL case, these infinitely many solutions allowed to accommodate several types of solutions at small distances: the Vainshtein solution (17) and the  $Q$ -scaling solutions (19) which depend on a free constant  $A_Q$ . These solutions also appear as nonperturbative. It is interesting to note that we can arrive at the same conclusion (that there is an infinite number of solutions at the infinity with the same asymptotic behavior) if we consider the weak-field equation approximation (see Sec. IV) as a starting point for this proof. The general structure of the additional solutions remains the same as for NLPF [see (40)], with, however, different exponent and preexponent factors. This difference is due to the fact that when getting an analog of (32) from the weak-field equations (58), some terms will be missing compared to (32).

<sup>6</sup>We recall that in GR, the vacuum solution depends on one integration constant, which can be identified with the Schwarzschild radius  $R_S$ .

$$\frac{\delta\lambda'}{z} + \frac{\delta\lambda}{z^2} + \frac{1}{2}(\delta\lambda + 3\delta\mu + z\delta\mu') = \frac{be^{-z}}{3z}\delta\nu, \quad (32a)$$

$$\frac{\delta\nu'}{z} - \frac{\delta\lambda}{z^2} - \frac{1}{2}(\delta\nu + 2\delta\mu) = 0, \quad (32b)$$

$$\frac{\delta\nu'}{2z} - \frac{\delta\lambda}{z^2} = \frac{be^{-z}}{3z^2}\left(\delta\mu'' - \delta\mu' - 2\delta\mu + \frac{\delta\lambda'}{z} + 2\delta\nu\right). \quad (32c)$$

Note that we substituted  $\sigma_0$  instead of  $\bar{\sigma}$  in (32), since  $\bar{\sigma} \rightarrow \sigma_0$  at  $z \rightarrow \infty$  and we are interested in solutions at infinity.

Until now, we have not made any assumptions about the form of  $\delta\sigma$ , except (31). If in addition one assumes that the rhs of (32) can be neglected, then the solution for  $\delta\sigma$  will be given by (24) provided that the integration constant  $b$  is replaced by  $\delta b$ , such that  $|\delta b| \ll |b|$ . However in this case one should impose  $\delta b = 0$  since we require  $\sigma \rightarrow \sigma_0$  at  $z \rightarrow \infty$ ; thus we do not obtain any new solution (this is, *mutatis mutandis*, what happens in GR).

Instead, let us try to find another solution such that (some) terms containing  $e^{-z}$  are not negligible in (32). In particular, we would like to keep the term containing  $\delta\mu''$ . Introducing a new variable  $\zeta$  as follows:

$$\zeta \equiv e^z,$$

the system of equations (32) can be rewritten as

$$\frac{\zeta\dot{\delta\lambda}}{\log\zeta} + \frac{\delta\lambda}{(\log\zeta)^2} + \frac{1}{2}(\delta\lambda + 3\delta\mu + \zeta\log\zeta\dot{\delta\mu}) = \frac{b\delta\nu}{3\zeta\log\zeta}, \quad (33a)$$

$$\frac{\zeta\dot{\delta\nu}}{\log\zeta} - \frac{\delta\lambda}{(\log\zeta)^2} - \frac{1}{2}(\delta\nu + 2\delta\mu) = 0, \quad (33b)$$

$$\frac{\zeta\dot{\delta\nu}}{2\log\zeta} - \frac{\delta\lambda}{\log^2\zeta} = \frac{b}{3\zeta\log^2\zeta}\left(\zeta^2\ddot{\delta\mu} - 2\delta\mu + \frac{\zeta\dot{\delta\lambda}}{\log\zeta} + 2\delta\nu\right), \quad (33c)$$

where the dot denotes the derivative with respect to  $\zeta$ . In order to simplify the system of equations (33), let us make further assumptions about the functions  $\delta\lambda$ ,  $\delta\nu$ , and  $\delta\mu$  at  $\zeta \rightarrow \infty$ :

$$|\delta\nu| \ll |\delta\mu| \ll \zeta|\dot{\delta\mu}| \ll \zeta^2|\ddot{\delta\mu}| \ll \zeta^3|\ddot{\delta\mu}|, \quad (34)$$

$$\log\zeta|\delta\lambda| \ll \zeta|\dot{\delta\lambda}| \ll \zeta|\delta\lambda|.$$

Equations (33) then take the simpler form

$$\dot{\delta\lambda} + \frac{1}{2}(\log\zeta)^2\dot{\delta\mu} = 0, \quad (35a)$$

$$\dot{\delta\nu} - \frac{\delta\lambda}{\zeta\log\zeta} - \frac{\log\zeta}{\zeta}\delta\mu = 0, \quad (35b)$$

$$\ddot{\delta\mu} + \frac{3}{2b\zeta}(2\delta\lambda - \zeta\log\zeta\dot{\delta\nu}) = 0. \quad (35c)$$

These equations can be combined to give one differential equation on  $\mu$ :

$$\zeta\ddot{\delta\mu} + \ddot{\delta\mu} - \frac{9}{4b}(\log\zeta)^2\dot{\delta\mu} - \frac{3\log\zeta}{b\zeta}\delta\mu = 0,$$

which—applying the assumptions (34)—can be rewritten as

$$\ddot{\delta\mu} - \frac{9}{4b\zeta}(\log\zeta)^2\dot{\delta\mu} = 0. \quad (36)$$

The asymptotic form of the decaying solution of Eq. (36) can be easily guessed:

$$\delta\mu = F_\infty(\zeta)\exp\left(-3\sqrt{\frac{\zeta}{b}}\log\zeta\right), \quad \text{for } \zeta \rightarrow \infty, \quad (37)$$

where  $F_\infty(\zeta)$  is a slowly varying function compared to  $\exp(-\frac{3}{\sqrt{b}}ze^{z/2})$  that cannot be fixed through this leading behavior analysis (cf. Appendix C for a more detailed approach based on series expansions). Note that a companion growing mode can also be considered

$$\delta\mu = F_\infty^{(\text{grow})}(\zeta)\exp\left(3\sqrt{\frac{\zeta}{b}}\log\zeta\right), \quad \text{for } \zeta \rightarrow \infty, \quad (38)$$

which is fixed to zero by the conditions (31); this growing mode can play an important role while integrating the equations of motion *à la* Runge-Kutta, since it can blow up very rapidly (as exponent of exponent) if sourced by numerical errors. This explains the difficulty one encounters while trying to integrate the equations of motion for distances beyond  $R_V$  for the potential (3). In the following, we will discard such a term since it contradicts the boundary conditions at infinity, but one should keep in mind its possible existence when solving the equations numerically.

From (35) one can find the other functions at  $\zeta \rightarrow \infty$ ,

$$\delta\lambda = -F_\infty(\zeta) \frac{(\log\zeta)^2}{2} \exp\left(-3\sqrt{\frac{\zeta}{b}} \log\zeta\right), \quad (39a)$$

$$\delta\nu = -F_\infty(\zeta) \frac{\sqrt{b}}{3} \frac{1}{\sqrt{\zeta}} \exp\left(-3\sqrt{\frac{\zeta}{b}} \log\zeta\right). \quad (39b)$$

Finally, one can check that the assumptions (31) and (34) are satisfied for the solutions (37) and (39) as  $\zeta \rightarrow \infty$  ( $z \rightarrow \infty$ ). In terms of the variable  $z$ , Eqs. (37) and (39) read

$$\delta\mu = F_\infty(z) \exp\left(-\frac{3}{\sqrt{b}} z e^{z/2}\right), \quad (40a)$$

$$\delta\lambda = -F_\infty(z) \frac{z^2}{2} \exp\left(-\frac{3}{\sqrt{b}} z e^{z/2}\right), \quad (40b)$$

$$\delta\nu = -F_\infty(z) \frac{\sqrt{b}}{3} \exp\left(-\frac{z}{2} - \frac{3}{\sqrt{b}} z e^{z/2}\right). \quad (40c)$$

Note that it is possible to complete this leading behavior analysis through a series expansion of the solution; we refer the reader to Appendix C for more details.

Thus we have shown that apart from the solution  $\bar{\lambda}$ ,  $\bar{\nu}$ , and  $\bar{\mu}$ , there is at least one family of asymptotic solutions at  $z \rightarrow \infty$ , parametrized by the function  $F_\infty$ , and given by

$$\lambda = \bar{\lambda} + \delta\lambda, \quad \nu = \bar{\nu} + \delta\nu, \quad \mu = \bar{\mu} + \delta\mu.$$

Note that the difference between solutions decays extremely fast—as a double exponent—as  $z \rightarrow \infty$ .

### C. Expansion inside the Vainshtein radius and outside the source

Let us now discuss what happens inside the Vainshtein radius, but outside the star. This will prove useful for different reasons. We will indeed recall how one can obtain an expansion around the GR solution, and hence compute the corrections to GR first found by Vainshtein. This will later be compared in Sec. VB to the numerical solution. We will also comment on the role of the integration constants appearing in the expansion procedure.

Let us first rely on our investigation of the DL, which provides interesting insights into the different types of solutions in the region  $R_\odot \ll R \ll R_V$ . It has been shown in [26] (as also recalled in Sec. IID) that, in the DL, two types of solutions exist in the region  $R_\odot \ll R \ll R_V$ : the so-called Vainshtein solution of Eq. (17) and the  $Q$ -scaling solutions (19). We are interested in obtaining, in the full theory, a solution which will be close to one of these DL solutions, for a range of distances as broad as possible. It has been recalled in Sec. IID that the  $Q$ -scaling solutions stop being relevant for the full nonlinear theory at the radius  $R_* = R_V(mR_S)^{1/5}$ . Below this distance, the DL is no longer useful for understanding the theory: new nonlinear terms enter into the game and one has to rely on numerics to investigate the existence of solutions. We will come back to numerics in Sec. V, but

we can already announce that our investigation has been negative in this direction: it has not been possible to us to find a regular solution of the full system which would exhibit a  $Q$ -scaling behavior in the range  $R_* = R_V(mR_S)^{1/5} \ll R \ll R_V$ : all our numerical  $Q$ -scaling solutions were developing strong instabilities around the problematic radius  $R_*$ , and integration was not possible for smaller distances. For this reason, we will mostly focus in the rest of this article on the Vainshtein solution, and try to generalize the DL version of this solution to the full system, under the form of series expansions. A first resolution method consists of looking for an expansion whose zeroth order is just the DL. A convenient parameter for such an expansion is the parameter  $a = R_V m$  introduced in Eq. (23). Indeed, one can show [26] that after a rescaling of the functions and equations, the limit  $a \rightarrow 0$  in the system (8) leads to the DL equations (12). The full solution expansion reads

$$\begin{aligned} \mu &= a^2 \sum_{n=0}^{\infty} \mu_n^{(a)}(R/R_V) a^{2n}, \\ \lambda &= a^4 \sum_{n=0}^{\infty} \lambda_n^{(a)}(R/R_V) a^{2n}, \\ \nu &= a^4 \sum_{n=0}^{\infty} \nu_n^{(a)}(R/R_V) a^{2n}, \end{aligned} \quad (41)$$

where the zeroth order functions  $a^2 \mu_0^{(a)}(R/R_V)$ ,  $a^2 \lambda_0^{(a)}(R/R_V)$ , and  $a^2 \nu_0^{(a)}(R/R_V)$  are the solutions of the DL system, whose expansions are given in Eq. (17). Even if the expansion (41) has the advantage to link explicitly the solution of the full system to the DL, it does not allow for an easy comparison with the GR solution. For this purpose, another expansion is possible as was originally proposed by Vainshtein [4].

Indeed, it is possible to look for a solution as an expansion in powers of the (square of the) graviton mass  $m^2$  dealing with the mass term as a perturbation around massless GR. Let us define this expansion as

$$\begin{aligned} \mu &= \sum_{k=0}^{\infty} \mu_k^{(m)}(R/R_S) m^{2k}, \\ \lambda &= \sum_{k=0}^{\infty} \lambda_k^{(m)}(R/R_S) m^{2k}, \\ \nu &= \sum_{k=0}^{\infty} \nu_k^{(m)}(R/R_S) m^{2k}. \end{aligned} \quad (42)$$

By definition, it is such that the lowest order is equal to the Schwarzschild solution

$$\lambda_0^{(m)} = -\nu_0^{(m)} = -\ln\left(1 - \frac{R_S}{R}\right) = \frac{R_S}{R} + \frac{1}{2}\left(\frac{R_S}{R}\right)^2 + \dots \quad (43)$$

Following Vainshtein, let us now briefly sketch how to obtain the first other nontrivial terms of the expansion. The function  $\mu_0^{(m)}$  can be found using the Bianchi

identity (8c), where  $\lambda$  and  $\nu$  have been replaced by their GR expression  $\lambda_0^{(m)}$ ,  $\nu_0^{(m)}$ :

$$-\frac{1}{R}f_g[\lambda_0^{(m)}, \lambda_0^{(m)'}, \nu_0^{(m)}, \nu_0^{(m)'}, \mu_0^{(m)}, \mu_0^{(m)'}, \mu_0^{(m)''}] = 0. \quad (44)$$

One can look for a solution for  $\mu_0^{(m)}$  as an expansion in powers of  $(R_S/R)$ , valid far from the source, where  $R \gg R_S$ . In this regime, the functions  $\lambda_0^{(m)}$ ,  $\nu_0^{(m)}$  go to zero, meaning that it makes sense to linearize the Bianchi equation in  $\lambda_0^{(m)}$  and  $\nu_0^{(m)}$  in order to find the dominant behavior of  $\mu_0^{(m)}$ . In addition, if  $\mu_0^{(m)}$  is assumed to also tend toward zero, it is consistent to only keep the terms of lowest order in this function. Since there is no linear term in  $\mu_0^{(m)}$  in the Bianchi equation, this lowest order is nothing else than the quadratic term  $Q$  defined in Eq. (13). The Bianchi equation (44) then reads, at dominant order

$$\frac{\lambda_0}{R^2} = \frac{\nu_0'}{2R} + Q(\mu_0).$$

Using the asymptotic behavior far from the source of the Schwarzschild solution (43), i.e. the Newtonian regime, one gets the equation

$$Q(\mu_0) = \frac{R_S}{2R^3},$$

which is identical to the DL equation (16). If one assumes that the function  $\mu_0^{(m)}$  can be expanded as a power series in  $R$ , one finds

$$\mu_0^{(m)} \sim \sqrt{\frac{8R_S}{9R}}. \quad (45)$$

The higher powers in  $R$  can then be found order by order in  $R$  [20]. Let us now compute the first order terms. The functions  $\lambda_1^{(m)}$  and  $\nu_1^{(m)}$  are computed through Einstein's equations (8a) and (8b). In the limit  $R_S \ll R \ll m^{-1}$ , these equations read

$$\begin{aligned} \frac{\lambda_1^{(m)'}}{R} + \frac{\lambda_1^{(m)}}{R^2} &= -\frac{m^2}{2}(3\mu_0^{(m)} + R\mu_0^{(m)'}), \\ \frac{\nu_1^{(m)'}}{R} - \frac{\lambda_1^{(m)}}{R^2} &= m^2\mu_0^{(m)}. \end{aligned}$$

Using the behavior (45), the leading behavior of the first order functions can be found to be

$$\nu_1^{(m)} = (mR)^2 \left[ \frac{2\sqrt{2}}{9} \sqrt{\frac{R_S}{R}} + \mathcal{O}\left(\frac{R_S}{R}\right) \right], \quad (46)$$

$$\lambda_1^{(m)} = (mR)^2 \left[ -\frac{\sqrt{2}}{3} \sqrt{\frac{R_S}{R}} + \mathcal{O}\left(\frac{R_S}{R}\right) \right],$$

while the function  $\mu_1^{(m)}$  is obtained using the Bianchi identity and reads

$$\mu_1^{(m)} = (mR)^2 \left[ \frac{6}{31} + \mathcal{O}\left(\sqrt{\frac{R_S}{R}}\right) \right]. \quad (47)$$

Comparing the zeroth order terms, given by Eqs. (43) and (45), and the first order ones, given by Eqs. (46) and (47), we reach the conclusion that the expansion in powers of

$m^2$  is only valid for  $R \lesssim R_V \equiv (m^{-4}R_S)^{1/5}$ . Hence, the functions (46) give the first nontrivial corrections to the Schwarzschild solution for distances below the Vainshtein radius. In Sec. V, it will be seen that they agree very well with the numerical results. Notice also that those corrections are nonanalytic in the expansion parameter  $\epsilon$  of the perturbation theory of Sec. III A. It is also interesting to notice that the expansion presented above and the small distance perturbative expansion of Sec. III A 1 both yield the same ‘‘Vainshtein radius’’  $R_V$  as a limiting distance of their respective domain of validity.

One may wonder which of the expansions (41) and (42) is the most appropriate for describing the solution of the full NLPF theory. There is no simple answer to this question since, eventually, all the orders are needed to fully recover the solution. Still, a simple but instructive calculation can be done: one can check that for  $R \lesssim (R_V m)^{8/7} R_V$ , the correction to the Newtonian potential  $\nu_N \sim -R_S/R$  due to the GR nonlinearities (i.e. the first correction to  $\nu_N$  in  $\nu_0^{(m)}$ ) dominates over the correction due to the DL nonlinearities (i.e. the first correction to  $\nu_N$  in  $a^4 \nu_0^{(a)}$ ). In mathematical terms, this reads  $(R_S/R)^2 \gtrsim (mR)^2 \sqrt{R_S/R}$ . In particular, this means that for sources of radius  $R_\odot \ll (R_V m)^{8/7} R_V$  (which is usually the case for sources we considered in our numerical investigations), the GR solution is a better approximation inside the source than the DL solution. We were able to confirm this simple analysis numerically (cf. Fig. 3).

Note that here, for the sake of simplicity, we have followed [20] and discarded any integration constant appearing in the process of obtaining the successive terms [beyond the GR part (43)] in the expansion (42). Such constants appear, in particular, at each order via a double integration of the Bianchi identity. If there is however a solution with such an expansion that can extend from  $R = 0$  to  $R = \infty$ , there is no reason for such constants to vanish, and they should be kept explicit until one does the matching with the source.

#### D. Expansion inside the source

Let us now turn to the interior of the source that we assume to be of constant density  $\rho$  and of radius  $R_\odot$ . In order to gain some understanding of the solution near the origin, it is possible to expand the solution as

$$\begin{aligned} \mu &= \sum_{n=0}^{\infty} \mu_n^{(x)} x^{2n}, & \lambda &= \sum_{n=1}^{\infty} \lambda_n^{(x)} x^{2n}, \\ \nu &= \sum_{n=0}^{\infty} \nu_n^{(x)} x^{2n}, & P &= \sum_{n=0}^{\infty} P_n^{(x)} x^{2n}, \end{aligned} \quad (48)$$

where we have defined the rescaled variable

$$x \equiv \frac{R}{R_S}.$$



Note that the expansion of the function  $\lambda$  starts at  $n = 1$  [in agreement with condition (11)]. These series depend on 3 integration constants that can be chosen to be  $\mu_0^{(x)} \equiv \mu(R = 0)$ ,  $\nu_0^{(x)} \equiv \nu(R = 0)$ , and  $P_0^{(x)} \equiv P(R = 0)$ . The other coefficients  $\mu_{n>0}^{(x)}$ ,  $\lambda_{n>0}^{(x)}$ , and  $\nu_{n>0}^{(x)}$  are then functions of these integration constants  $\mu_0^{(x)}$ ,  $\nu_0^{(x)}$ ,  $P_0^{(x)}$  and the mass  $m$  of the graviton. These coefficients can be found solving order by order in  $x^2$ ; a few of them are given for the AGS potential in Appendix B.

#### IV. WEAK-FIELD APPROXIMATION

As we have just seen, different expansion schemes can be used in order to get an understanding of the putative solution. However, those schemes all fail to describe correctly all the key features of the solution. More specifically, the DL scheme (Sec. IID) describes well the Vainshtein crossover, but misses the Yukawa decay. The perturbation theory approach (Sec. III A) catches the Yukawa decay, but misses the Vainshtein crossover, while the expansion around GR misses both the Yukawa decay and the Vainshtein crossover (Sec. III C). The purpose of this section is to introduce an expansion scheme, called here weak-field approximation, that will retain both the Yukawa decay and the Vainshtein crossover.

We expand the full system of Einstein equations together with Bianchi identity in powers of  $\lambda$ ,  $\nu$ ,  $\mu$ , and assume the smallness of each function  $\mu$ ,  $\nu$ , and  $\lambda$  as well as of their derivatives:

$$\{\lambda, \nu, \mu\} \ll 1, \quad \{R\lambda', R\nu', R\mu'\} \ll 1, \quad R^2\mu'' \ll 1. \quad (49)$$

Expanding (8) and using our assumptions (49) we obtain

$$\frac{\lambda'}{R} + \frac{\lambda}{R^2} = -\frac{m^2}{2}(\lambda + 3\mu + R\mu' + \mathcal{O}(\nu^2)) + \frac{\rho}{M_P^2}, \quad (50a)$$

$$\frac{\nu'}{R} - \frac{\lambda}{R^2} = \frac{m^2}{2}(\nu + 2\mu + \mathcal{O}(\lambda^2) + \mathcal{O}(R^2\mu'^2)), \quad (50b)$$

$$\begin{aligned} \lambda - \frac{R\nu'}{2} &= -R^2\left(\frac{\mu'^2}{4} + \frac{2\mu\mu'}{R} + \frac{\mu\mu''}{2}\right) + \mathcal{O}(\nu^2) \\ &\quad + \mathcal{O}(\lambda^2 R^2) + \mathcal{O}(\nu\lambda'R) \\ &\quad + \mathcal{O}(\{\nu, \lambda'\} \times \{\mu, \mu'R, \mu''R^2\}). \end{aligned} \quad (50c)$$

Note that  $P$  on the rhs of (50b) disappears as a consequence of the conservation equation (9) and the assumption of the weak-field regime, similar to the DL case. One can recognize, on the rhs of Eq. (50c), the nonlinear DL part responsible for the Vainshtein mechanism. Note that so far we did not make any assumptions about the relation between  $\lambda$ ,  $\nu$ , and  $\mu$  or between functions and their derivatives, and this is the reason why we kept the quadratic terms in  $\mu$  on the rhs of (50c).

To further simplify the system of equations (50) we assume that the quadratic terms  $\sim \nu^2$  are small in comparison to the linear terms on the rhs of (50a), i.e.,

$$\nu^2 \ll \max\{\lambda, \mu\},$$

and similarly that the term  $\sim \lambda^2$  on the rhs of (50b) is negligible compared with  $\nu$ . We also assume that the quadratic terms containing  $\mu$  and its derivatives are negligible compared to the linear terms containing  $\mu$ . It is important to note that the last assumptions eliminate the quadratic terms in Eqs. (50a) and (50b), while we have to keep the quadratic terms in  $\mu$  in (50c), since the linear term in  $\mu$  is absent in this equation. Note that if we assumed at this step that  $\mu \sim \nu \sim \lambda$ , then the quadratic terms in  $\{\mu, \mu', \mu''\}$  on the rhs of (50c) would be dropped. It is important, however, that we do not assume any relations between  $\{\mu, \mu', \mu''\}$  and  $\{\nu, \lambda, \nu', \lambda'\}$ , and thus we keep these terms. This will be the key feature of the expansion scheme introduced here. As a result we obtain the system

$$\frac{\lambda'}{R} + \frac{\lambda}{R^2} = -\frac{m^2}{2}(\lambda + 3\mu + R\mu') + 8\pi G_N \rho, \quad (51a)$$

$$\frac{\nu'}{R} - \frac{\lambda}{R^2} = \frac{m^2}{2}(\nu + 2\mu), \quad (51b)$$

$$\frac{\nu'}{2R} - \frac{\lambda}{R^2} = -Q(\mu, \mu', \mu''), \quad (51c)$$

with  $Q$  given by (13). It is then convenient to use once again the rescaled radial coordinate  $z = Rm$  introduced in Sec. III A, such that we obtain from (51),

$$\frac{\lambda'}{z} + \frac{\lambda}{z^2} = -\frac{1}{2}(\lambda + 3\mu + z\mu') + \bar{\rho}, \quad (52a)$$

$$\frac{\nu'}{z} - \frac{\lambda}{z^2} = \frac{1}{2}(\nu + 2\mu), \quad (52b)$$

$$\frac{\lambda}{z^2} - \frac{\nu'}{2z} = Q(\mu, \mu', \mu''), \quad (52c)$$

where

$$Q(\mu, \mu', \mu'') = -\left(\frac{\mu'^2}{4} + \frac{\mu\mu''}{2} + \frac{2\mu\mu'}{z}\right), \quad (53)$$

and we introduced  $\bar{\rho} \equiv 8\pi G_N \rho / m^2$ . This system of equation corresponds to the one left over in the limit considered here. As will be seen, it has the required properties.

First, note that it is possible to rewrite this system as a single ODE on  $\mu$  and two algebraic relations between  $\lambda$ ,  $\nu$ , and  $\mu$  and its derivatives. Indeed, expressing  $\lambda$  from (52b) in terms of  $\nu$ ,  $\nu'$ , and  $\mu$  and substituting it to Eqs. (52a) and (52c) we find two equations containing only  $\nu$  and  $\mu$  (and their derivatives):

$$\nu'' + \frac{2\nu'}{z} - \frac{1}{4}(z^2 + 6)\nu - \frac{1}{2}z\mu' - \frac{1}{2}(z^2 + 3)\mu = \bar{\rho}, \quad (54a)$$

$$-\frac{\nu'}{2z} + \frac{\nu}{2} + \mu + Q = 0. \quad (54b)$$

Taking the derivative of (54b) with respect to  $z$  and substituting the resolved  $\nu''$  into (54a) we obtain a system of two equations:

$$\frac{z^2 + 3}{z} \nu' - \frac{1}{4}(z^2 + 6)\nu - \frac{1}{2}(z^2 + 3)\mu + \frac{z}{2}(4Q' + 3\mu') = \bar{\rho}, \quad (55a)$$

$$-\frac{\nu'}{2z} + \frac{1}{2}\nu + \mu + Q = 0. \quad (55b)$$

From the above system of equations (55) we can find  $\nu'$  in terms of  $\mu$  and its derivatives and  $\nu$  in terms of  $\mu$  and its derivatives,

$$\nu(z) = \mathcal{F}, \quad \nu'(z) = z\mathcal{F} + 2z(Q + \mu). \quad (56)$$

where we denoted

$$\mathcal{F} \equiv -\frac{2}{3(z^2 + 2)}[z(4Q' + 3\mu') + (z^2 + 3)(4Q + 3\mu) - 2\bar{\rho}]. \quad (57)$$

Thus we arrive at the following single ODE on  $\mu$ :

$$\frac{d\mathcal{F}}{dz} = z\mathcal{F} + 2z(Q + \mu).$$

The above equation can be rewritten in a closed form,

$$\begin{aligned} (2 + z^2)(4Q'' + 3\mu'') + \frac{2}{z}(4 + z^2)(4Q' + 3\mu') \\ - (z^2 + 4)(4Q + 3\mu) - Q(z^2 + 2)^2 \\ = (2 + z^2)\frac{2\bar{\rho}'}{z} - 2(4 + z^2)\bar{\rho}. \end{aligned} \quad (58)$$

Note that Eq. (58) is of the fourth order, meaning that the number of boundary conditions needed to solve the equation is four. This is in agreement with the number of boundary conditions needed to specify for Eqs. (51). Equation (58) can further be simplified under appropriate assumptions. First of all, let us recover the solution (24) for the linearized equations. To do so, we drop in (58) all nonlinear terms, i.e. all the terms containing  $Q$  and its derivatives, as well as the source terms (containing  $\bar{\rho}$  and  $\bar{\rho}'$ ) to get

$$(2 + z^2)\mu'' + \frac{2}{z}(4 + z^2)\mu' - (z^2 + 4)\mu = 0.$$

The nongrowing at the infinity solution of the above equation reads

$$\mu = \frac{Ce^{-z}(z^2 + z + 1)}{z^3},$$

which coincides with Eq. (24) and catches the Yukawa decay.

Another limit can be obtained assuming in (58) as follows. We assume  $Qz \ll Q'$  and  $\mu z \ll \mu'$  together with  $z \ll 1$ . Neglecting subleading terms in (58), one obtains

$$(4Q'' + 3\mu'') + \frac{4}{z}(4Q' + 3\mu') = \frac{2\bar{\rho}'}{z} - 4\bar{\rho}.$$

The above equation can be integrated twice to give

$$2Q + \frac{3}{2}\mu = \frac{1}{z^3}\left(C_0 + \int_0^z z^2 \bar{\rho} dz\right) + C_1, \quad (59)$$

where in the last expression we also assumed

$$\int_0^z z^2 \left(\int_0^z \bar{\rho} \tilde{z} d\tilde{z}\right) dz \ll \int_0^z z^2 \bar{\rho} dz.$$

The constant of integration  $C_1$  in (59) must be set to 0 to get correct behavior of  $\mu$  at infinity (otherwise the solution is nonflat at large distances). Similarly,  $C_0 = 0$ , by matching  $\mu$  inside the Vainshtein radius and outside the source. Restoring physical units,  $R = z/m$ , one can see that we arrived at the previously found DL equation (14) catching the Vainshtein crossover. Furthermore, the analytic solution of the weak-field system (52) will be seen to reproduce very well the numerical integration (in the expected range) in the following section.

## V. INTEGRATION OF THE FULL SYSTEM: NUMERICAL RESULTS

We have integrated the equations of motion (8) and (9) using two different approaches. The first one is based on a resolution by relaxation, which turns out to be very well adapted to the problem. This method and the related results are presented, respectively, in the Secs. VA and VB. The second approach, which is described in Sec. VC, consists of a shooting method *à la* Runge-Kutta. Both methods perfectly agree, demonstrating the robustness of our numerical investigations. A reader mainly interested in the numerical results can read only Sec. VB which is self-contained.

### A. Relaxation approach: Method and boundary conditions

The basic idea of the relaxation method consists of choosing an initial configuration for the solution, and then deforming it until the equation we want to solve is satisfied with a good enough precision. The initial guess is of prime importance, since it will determine whether the iteration converges or not. In the case of the NLPF equations of motion (8) and (9), a clever guess is needed: indeed, starting from a naive guess (e.g. all the functions to zero) does not lead to a converging iteration. Here is where the DL plays a crucial role in finding a solution of the full system: we can use the DL solution as a guess for the full solution. Note that the DL being simpler than the full system, the DL solution itself is not too hard to find (no specific guess is needed).

The advantage of the relaxation approach is that boundary conditions can be easily enforced. In particular, this method allows one to fix boundary conditions at different points, for example, at the two ends of the interval on which we want to solve the equations. This flexibility in the boundary conditions fixing makes this method very well fitted for static solution studies in which we want to impose conditions both in the center and very far from the source.

In our case, we worked with the dimensionless variable

$$\xi \equiv R/R_V,$$

as well as with the rescaled functions  $w$ ,  $u$ , and  $v$  defined by

$$w = a^{-2}\mu, \quad v = a^{-4}\nu, \quad u = a^{-4}\lambda.$$

We fixed the boundary condition at  $\xi \equiv R/R_V = 0$  and at some distance  $\xi_\infty$  which is chosen to be much larger than the Vainshtein radius (typically  $\xi_\infty \sim 100$ ). The boundary conditions in the DL case are given by the DL linear regime (15) and read

$$\begin{aligned} \dot{w}(\xi = 0) &= 0, \\ u(\xi = 0) &= 0, \\ \dot{w}(\xi = \xi_\infty) &= -\frac{2}{\xi_\infty^3}, \\ v(\xi = \xi_\infty) &= -\frac{4}{3\xi_\infty}, \\ P(\xi = \xi_\infty) &= 0. \end{aligned}$$

One may notice that the condition on  $w$  at infinity is imposed on  $\dot{w}$  rather than on  $w$ . This is not necessary, but we found it easier to proceed that way in practice. Once the DL solution has been found, we can turn to the fully nonlinear system. In that case, the boundary conditions are given by the linear solution (24) and read

$$\begin{aligned} \dot{w}(\xi = 0) &= 0, \\ u(\xi = 0) &= 0, \\ w(\xi = \xi_\infty) &= C \times \frac{2a^2}{3\xi_\infty} \left( 1 + \frac{1}{a\xi_\infty} + \frac{1}{(a\xi_\infty)^2} \right) e^{-a\xi_\infty}, \\ v(\xi = \xi_\infty) &= -C \times \frac{4}{3} \frac{e^{-a\xi_\infty}}{\xi_\infty}, \\ P(\xi = \xi_\infty) &= 0. \end{aligned}$$

We have explicitly kept in these boundaries conditions the integration constant  $C$  defined in Eq. (24) [while  $a$  is defined as in Eq. (23)]. This procedure allows one to check numerically that the choice we made  $C = 1$  in Eq. (25) is in agreement with numerics.

In our study of NLPF, we have implemented the relaxation method described in [28], using the C++ language. We typically used a grid of  $N \sim 5 \times 10^4$  points and aimed at a precision of  $10^{-6}$ . We always checked that our numerical

solutions were stable under the change of the step size of the grid and the precision.

The density of the star is assumed to be given by  $\rho = M/(4\pi R_\odot^3)$  inside the star, and null outside.

## B. Numerical results

To summarize here shortly our main results, we were able to show that one can obtain a nonsingular solution showing the recovery of GR *à la* Vainshtein. The relaxation method allowed us to solve the equations of motion (8) and (9) on a wide range of the parameter  $a = R_V m$  ( $a \in [10^{-3}, 0.6]$ ) and for very large intervals (typically for  $R$  between 0 and  $\sim 100R_V$ ). In the following we first discuss the behavior of the functions  $\lambda$  and  $\nu$  which appear in the “physical” metric  $g_{\mu\nu}$  (Sec. VB 1), comparing their behavior in NLPF and GR, and showing, in particular, how our numerical integration agrees with the corrections to GR as predicted by Vainshtein. Then (Sec. VB 2) we turn to discuss the obtained behavior of the pressure inside the star, comparing again the numerical solution with the one obtained in GR (recall that the star’s density is assumed to be constant). Last we discuss our numerical result for the

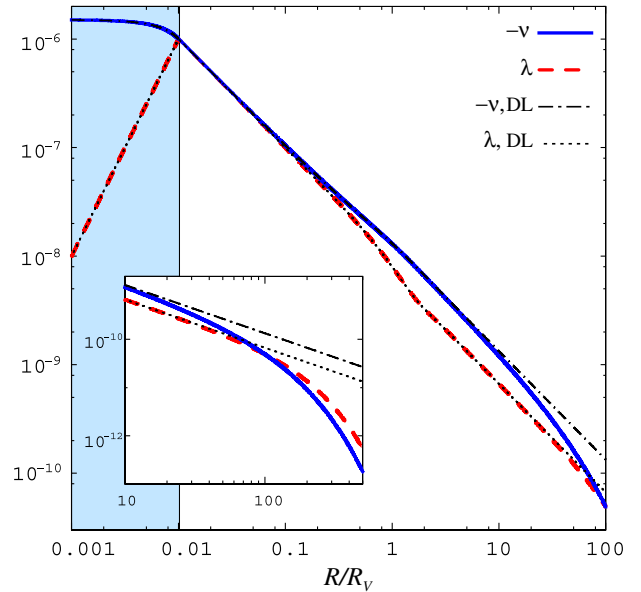


FIG. 1 (color online). Plot of the metric functions  $-\nu$  and  $\lambda$  vs  $R/R_V$ , in the full nonlinear system and the decoupling limit (DL), with a star of radius  $R_\odot = 10^{-2}R_V$  and  $m \times R_V = 10^{-2}$ . For  $R \ll R_V$ , the numerical solution is close to the GR solution (where in particular  $\nu \sim -\lambda \sim -R_S/R$  for  $R > R_\odot$ ). For  $R \gg R_V$ , the solution enters a linear regime. Between  $R_V$  and  $m^{-1}$ , where the DL is still a good approximation, one has  $\nu \sim -2\lambda \sim -4/3 \times R_S/R$ . At distances larger than  $m^{-1}$  the metric functions decay *à la* Yukawa as appearing more clearly in the inset. The latter shows the same solution but for larger values of  $R/R_V$ , and in the range of distance plotted there, the numerical solutions are indistinguishable from the analytic solutions of the linearized field equations (24). This plot was already presented as Fig. 1 of Ref. [26].

function  $\mu$  appearing in the “nonphysical” metric with our gauge choice (9).

### 1. Behavior of the functions $\lambda$ and $\nu$

Figure 1 summarizes our main results and illustrates the following points:

- (i) at short distances ( $R \ll R_V$ ), the solution is very close to the GR solution, and in particular  $\nu \sim -\lambda \sim -R_S/R$  in the neighborhood of the source,
  - at large distances ( $R \gg R_V$ ), the solution is in the linear regime (24). Between  $R_V$  and  $m^{-1}$ , the solution is very close to the DL solution (Table I), and in particular  $\nu \sim -2\lambda \sim -4/3 \times R_S/R$ . Beyond  $m^{-1}$ , the solution is exponentially decreasing,
- (ii) the behavior at infinity allows one to check that the choice of integration constant we made in Eq. (25) is consistent with numerics. More precisely, we tried to solve numerically the equations of motion with an integration constant  $\mathcal{C}$  slightly different from 1, and we noted that the relaxation method was not converging, unless  $|\mathcal{C} - 1| \ll 1$ .

We have seen in Sec. III C that the leading behavior of the difference between the functions  $\lambda$  and  $\nu$  and their GR

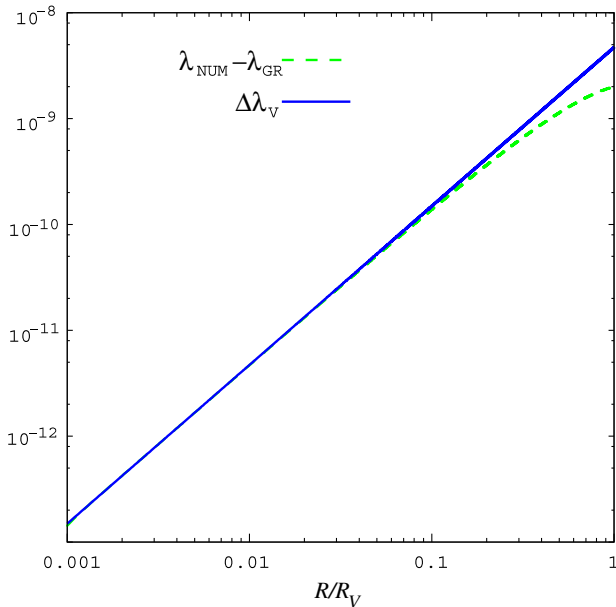


FIG. 2 (color online). Plot of the difference between the numerical solution for  $\lambda$ ,  $\lambda_{\text{NUM}}$  and the one of GR,  $\lambda_{\text{GR}}$ . This is compared to the computation of the same quantity using the first term in the expansion in  $m^2$  proposed by Vainshtein,  $|\Delta\lambda_V| = (\sqrt{2}/3)(mR)^2\sqrt{R_S/R}$ . Both functions are plotted vs  $R/R_V$ , for a source of radius  $R_\odot = 10^{-3}R_V$  and a choice of parameters such that  $a \equiv m \times R_V = 10^{-2}$ . It can be seen that the numerical solution agrees very well with Vainshtein’s expression given by Eq. (46) for  $R \ll R_V$ . For  $R \sim R_V$ , the expansion in  $m^2$  is, as expected, no longer a good approximation of the solution.

equivalent,  $\lambda_{\text{GR}}$  and  $\nu_{\text{GR}}$  is given by the expressions (46). It is possible to check that this theoretical prediction is very well satisfied by our numerical solution, as can be seen in Fig. 2. In this figure, we plotted the difference between the function  $\lambda$  obtained from our numerical integration and the one of GR, and compared this with the first correction to GR as computed using the Vainshtein mechanism. As expected, the two agree very well for distances smaller than the Vainshtein radius.

Besides, we have shown at the end of Sec. III C that deep inside the Vainshtein radius, the NLPF solution should be closer to the GR solution than to the DL solution, illustrating the fact that the GR nonlinearities are starting to play an important role. This is seen in Fig. 3, where we compared the function  $\nu$  (after an appropriately chosen rescaling) for various values of the parameters. It appears on this figure that the numerical solutions well agree with the GR behavior, while slightly differing from the DL expression for large  $a$ . This illustrates the fact that inside the source, the nonlinearities coming from the GR terms of Eq. (8), and which are not taken into account in the DL, are important.

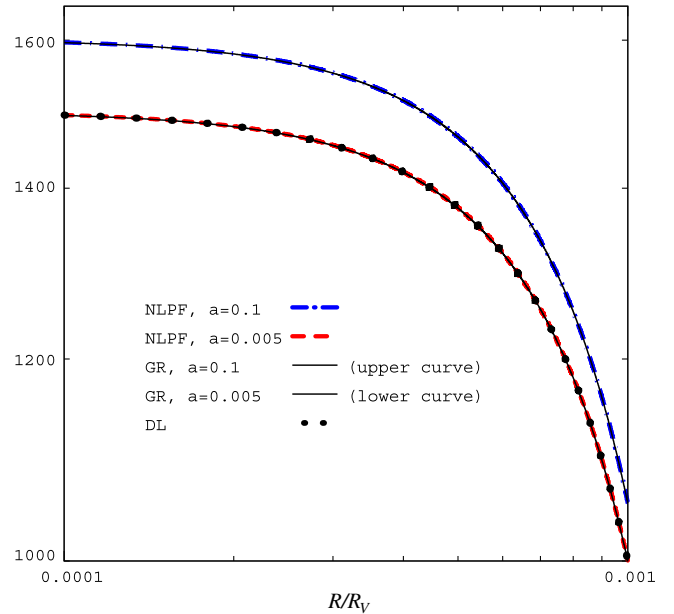


FIG. 3 (color online). Plot of  $-\nu \times a^{-4}$  vs  $R/R_V$  (the  $a^{-4}$  factor is included for convenience such that, in the DL, all plotted theories would exactly coincide) inside the source of radius  $R_\odot = 10^{-3}R_V$ , for three different values of  $a \equiv m \times R_V$ . Along with the numerical solution, the GR analytical solution for  $\nu$  is plotted for each value of  $a$ . It can be seen that the numerical solutions perfectly agree with the GR behavior, while slightly differing from the DL expression for large  $a$ . This illustrates the fact that inside the source, the nonlinearities coming from the GR terms of Eq. (8), and which are not taken into account in the DL, are important. For small  $a$ , both the GR and the numerical solutions agree with the DL solution, which encodes for most of the physics, in agreement with the fact that the DL corresponds to the limit  $a \rightarrow 0$ .



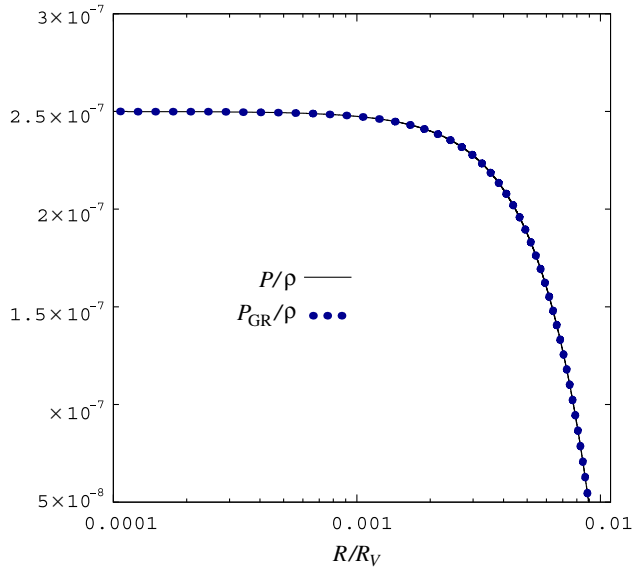


FIG. 4 (color online). Plot of the functions  $P/\rho$  and  $P_{GR}/\rho$  vs  $R/R_V$ , for a source of radius  $R_\circ = 10^{-2}R_V$ , of pressure  $P$ , and of constant density  $\rho$ , and for a parameter choice such that  $a \equiv m \times R_V = 10^{-2}$ . The function  $P/\rho$  is found numerically, while  $P_{GR}/\rho$  is the pressure computed in GR and is given by the analytical expression (60). The two curves are indistinguishable from each other, illustrating the fact that Vainshtein's conjecture is valid inside the star.

## 2. Behavior of the pressure $P$

In GR, the pressure inside a static spherically symmetric source of constant density can be computed analytically to be

$$P_{GR} = \rho \frac{\sqrt{1 - \frac{R_s}{R_\circ^3} R^2} - \sqrt{1 - \frac{R_s}{R_\circ}}}{3\sqrt{1 - \frac{R_s}{R_\circ}} - \sqrt{1 - \frac{R_s}{R_\circ^3} R^2}}. \quad (60)$$

This is compared to Fig. 4 with our numerical solution, confirming once again the validity of Vainshtein's conjecture.

## 3. The function $\mu$

For the range of parameters investigated here, the numerical behavior of the function  $\mu$  between  $R = 0$  and  $R = m^{-1}$  is very close to the Vainshtein solution (Table I) of the DL. Beyond  $m^{-1}$ , the Yukawa decay seen in the linear solution (24) is found to occur. These various regimes are shown in Fig. 5 which plots the (rescaled) function  $\mu$  for three different values of the parameters.

Figure 6 shows the same, zoomed inside the star. We notice there that the solution of the full theory goes toward the DL solution as the parameter  $a \equiv R_V \times m$  tends toward 0 and we recover the fact that the corrections to the DL due to the nonlinearities which are not taken into account in the DL are of order  $\mathcal{O}(a^2)$  and goes to zero as  $a \rightarrow 0$ . This plot also shows that our numerical integration reproduces nonlinearities not present in the DL.

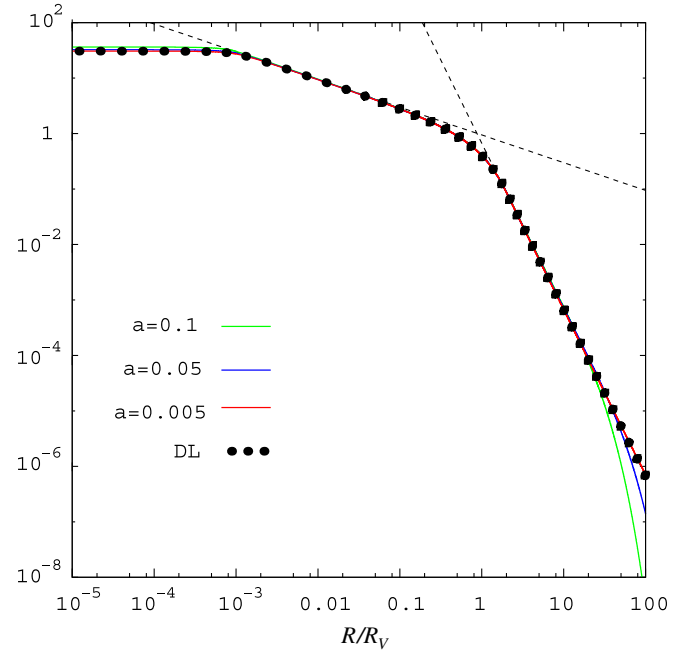


FIG. 5 (color online). Plot of  $a^{-2}\mu$  vs  $R/R_V$  (the  $a^{-2}$  factor is included for convenience such that, in the DL, all plotted theories would exactly coincide) for a source of radius  $R_\circ = 10^{-3}R_V$  for three different values of  $a \equiv m \times R_V = 0.005, 0.05$ , and  $0.1$  as well as in the DL case (which corresponds to  $a \rightarrow 0$ ). The three DL regimes of Table I are clearly distinguishable: for  $R < R_\circ$ ,  $\mu \sim \text{const}$ ; for  $R_\circ < R \ll R_V$ ,  $\mu \sim \sqrt{(8R_V)/(9R)}$ ; for  $R_V \ll R \ll m^{-1}$ ,  $\mu \sim 2R_V^3/(3R^3)$ . For  $R \gg m^{-1}$ , we observe the Yukawa decay of the linear solution (24). At the resolution of the picture, the three numerical solutions are hard to distinguish for distances  $R$  smaller than  $m^{-1}$ . The next figure shows a zoom of this figure at small distances (namely inside the source) where the differences between the three solutions are easily seen.

## C. Shooting method

In addition to the relaxation approach that we presented in the previous sections, we developed a shooting approach to solve the system of equations (8) and (9). This shooting approach, which is based on a direct Runge-Kutta integration, appears to be quite difficult to implement for the theory (3). Indeed, there appear numerical instabilities for distances beyond the Vainshtein radius and probably related to the growing mode (38). As a consequence of these instabilities, the initial conditions need to be extremely finely tuned.<sup>7</sup> Nevertheless, we have been able to integrate the system of equations outward from a point

<sup>7</sup>Note however that the shooting integration appears more stable for some NLPF theories which have a different interaction term from (3). This is the case, for example, for the BD theory (using the terminology of Ref. [26]). In this theory, there is no admissible Vainshtein scaling, but only  $Q$  scalings in the DL. However, in the DL and in the full nonlinear case, the shooting method appears much more stable than in the model considered here, even though the  $Q$ -scaling solution does not seem to continue into a nonsingular solution in the beyond DL regime.

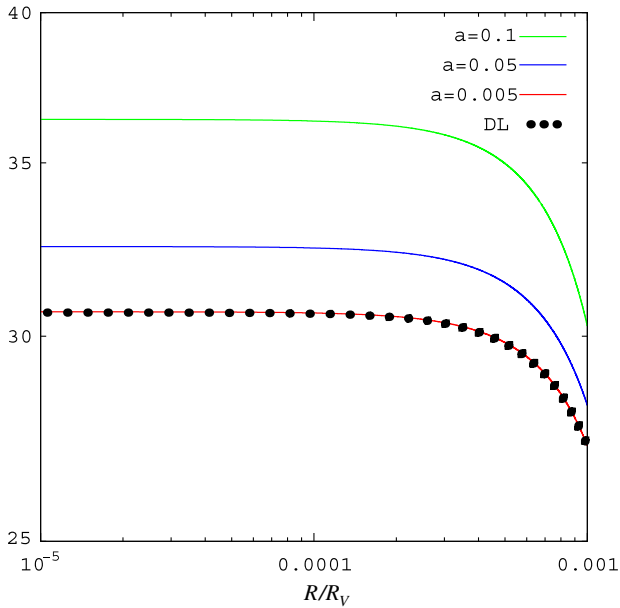


FIG. 6 (color online). Plot of  $a^{-2}\mu$  vs  $R/R_V$  (the  $a^{-2}$  factor is included for convenience such that, in the DL, all plotted theories would exactly coincide) for a source of radius  $R_\odot = 10^{-3}R_V$  for three different values of  $a \equiv m \times R_V = 0.005, 0.05,$  and  $0.1$  as well as in the DL case. The picture is a zoom of Fig. 5, inside the star. One can check that even if the three curves for  $a = 0.005, 0.05,$  and  $0.1$  are close to each other and to the DL solution, they differ inside the star due to nonlinearities not captured by the DL. One notes that the full solution goes toward the DL solution when  $a \equiv R_V \times m$  tends toward 0.

deep inside the star up to a point farther than the Vainshtein radius (typically  $R \sim 3R_V$ ). We have also been able to integrate inward starting from a point far from the source (e.g.  $R \sim 3R_V$ ) up to the core of the source. In all the integrations, we have made an extensive use of series expansions of the solution in order to fix the initial conditions as well as possible.

While integrating inward from large distances (i.e.  $R \sim 3R_V$ ), the solution adopts almost systematically a  $Q$ -scaling behavior, similar to the one of Eq. (19), except if the initial conditions are extremely fine-tuned to get the Vainshtein solution. In the cases in which the  $Q$  scaling is found, our numerical solutions always encountered a singularity around the radius we identified in Eq. (21) and which corresponds to the scale at which nonlinearities nontaken into account by the DL start to play a role. Of course, this does not mean that a  $Q$ -scaling solution cannot be found but we were not able to find any and it convinced us to focus on the Vainshtein solution.

To obtain the latter, we have found it particularly convenient to integrate from a point located inside the source (typically  $R_{\min} = 10^{-3}R_\odot$ ). To do so, we needed initial conditions at  $R_{\min}$ , i.e.  $\mu(R_{\min}), \mu'(R_{\min}), \lambda(R_{\min}), \nu(R_{\min}),$  and  $P(R_{\min})$ . These initial conditions are computed using the expansion (48) where the three integration constants  $\mu_0^{(x)} \equiv \mu(R=0), \nu_0^{(x)} \equiv \nu(R=0),$  and

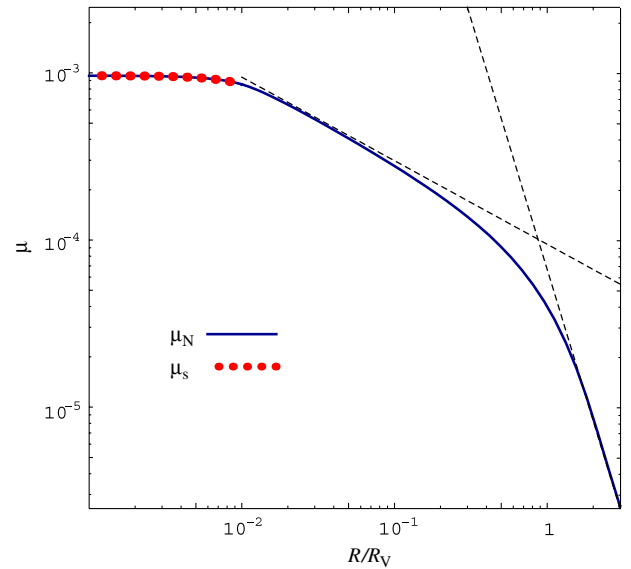


FIG. 7 (color online). Plot of the function  $\mu$  ( $\mu_N$  as given by the numerical integration) vs  $R/R_V$ , for a source of radius  $R_\odot = 10^{-2}R_V$  and of constant density  $\rho$ , and for a choice of parameters such that  $a \equiv m \times R_V = 10^{-2}$ . The plot has been obtained using an direct integration *à la* Runge-Kutta. We also added the plot of series  $\mu_S = \mu_0^{(x)} + \mu_1^{(x)}x^2 + \mu_2^{(x)}x^4$ , which was introduced in the main text as an expansion inside the star: the two curves are indistinguishable from each other.

$P_0^{(x)} \equiv P(R=0)$  have been fixed using the values obtained through the relaxation method. In order to be able to obtain a wide enough range of integration, the required relative precision on the three integration constants  $\mu_0^{(x)}, \nu_0^{(x)},$  and  $P_0^{(x)}$  is very stringent. For example, if one wants to be able to integrate from  $R_{\min} = 10^{-3}R_\odot$  to  $R_{\max} = 3R_V$ , in the case  $a = R_V m = 10^{-2}$  and  $R_\odot = 10^{-2}R_V$ , the required relative precision on the integration constant  $\mu_0^{(x)}$  is smaller than  $10^{-8}$ , in order to ensure that the numerical solution at  $R_{\max} = 3R_V$  differs from the linear solution (24) (which is supposed to be an excellent approximation of the solution at that distance) by less than 5%. The result of such an integration is given in Fig. 7. This figure also shows part of the expansion (48) inside the star. It agrees very well with the numerical solution.

#### D. Weak-field approximation

We have also been able to check numerically the weak-field approximation introduced in Sec. IV. Namely, as appears on Fig. 8, this approximation reproduces very well the numerical solution on a range of distances much larger than the decoupling limit, since this range extends at distances larger than the graviton Compton length  $m^{-1}$ . At small distances, however, as also seen on Fig. 8, the weak-field approximation deviates from the numerical solution. This is expected since the weak-field approximation reduces there to the DL and does not capture some nonlinearities.

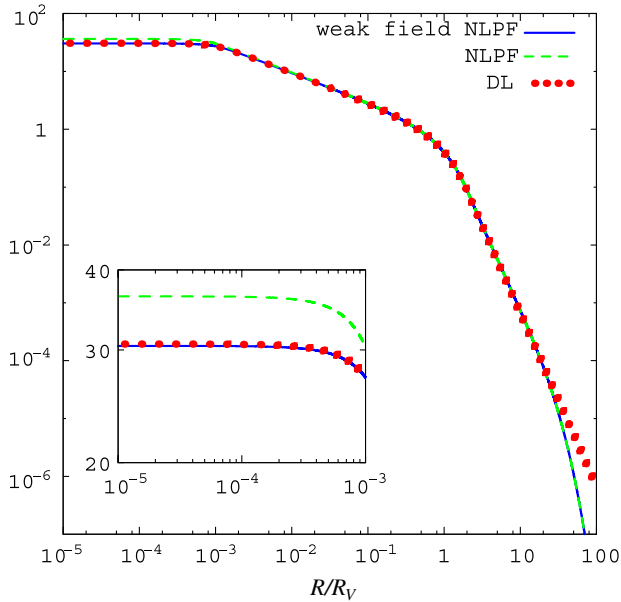


FIG. 8 (color online). Plot of the function  $a^{-2}\mu$  (the  $a^{-2}$  is added to make the comparison between  $\mu$  and its DL behavior easier), vs  $R/R_V$ , for a source of radius  $R_\circ = 10^{-3}R_V$ , of constant density  $\rho$  and a choice of parameters such that  $a \equiv m \times R_V = 10^{-1}$ . The three curves correspond, respectively, to the weak-field solution of Eq. (58), to the solution of the fully nonlinear theory and to the DL solution. One can see that the weak-field model allows one to recover the general behavior of the full solution, except inside the star (as seen in the panel showing a zoom of the plots for small  $R$ ) where it is very close to the DL, and thus misses some nonlinearities.

## VI. CONCLUSIONS

We have studied in detail the Vainshtein mechanism for nonlinear Pauli-Fierz theory as defined in Sec. II. More specifically, we have been able to show in the case of a specific NLPF model [the one defined by the interaction term (3)] that the recovery of the Schwarzschild solution works as forecasted by Vainshtein. This proof was provided via numerical integration of the field equation using two different methods: a relaxation scheme and a shooting method. It also heavily used series expansions (to fix boundary/initial conditions) which were first presented in this work. Our result, already presented in brief in [21] disagrees with a previous work on the same issue [20], which concluded that the Vainshtein mechanism was not working, because it could not find numerically any nonsingular asymptotically flat solution of the NLPF field equation. We also presented arguments, based on various expansions, showing that the solution at infinity is not uniquely defined by standard perturbation theory, but has nonperturbative hairs allowing one to match sources at smaller distances. This, as well as the more sophisticated integration methods we used, could explain the differences between our results and the ones of Ref. [20]. Besides the model defined by the interaction term (3), we also investigated other models, in

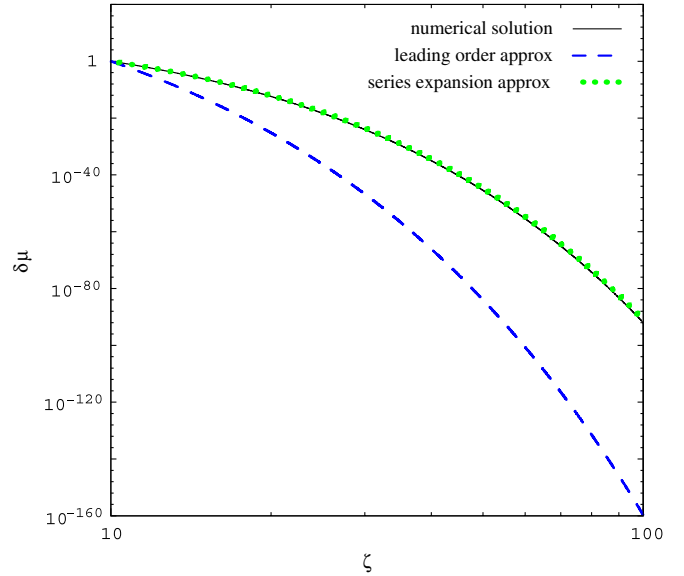


FIG. 9 (color online). Plot of the numerical solution for  $\delta\mu(\zeta)$  for the choice of parameters  $b = 0.1$  and  $\delta\mu(\zeta = 10) = 1$ . We also plotted the leading behavior  $\delta\mu(\zeta) = \exp(-3\sqrt{\frac{\zeta}{b}} \log \zeta)$  of Eq. (37), and the first series expansion correction obtained from Eqs. (C5) and (C13). In terms of the  $\zeta$  variable, this first order series expansion correction reads  $\delta\mu(\zeta) = \frac{1}{(\log \zeta)^{5/4} (1 + \log \zeta)^{1/4}} \exp(\frac{\log \zeta}{4} - \frac{1}{\sqrt{b}} (3 \log \zeta - \frac{15}{2} + \frac{9}{8 \log \zeta}) \sqrt{\zeta})$ . One can see that while the leading behavior already encodes the main behavior [exponentially decreasing  $\delta\mu(\zeta)$ ], the first correction coming from the series expansion improves greatly the agreement with the numerical solution, confirming our analytical analysis of the system (C2).

particular, those which were shown in Ref. [26] to possess only a non-Vainshtein-like scaling (dubbed  $Q$  scaling in Ref. [26]) in the DL. According to our numerical investigations, the DL solutions with  $Q$  scaling do not seem to continue into nonsingular solutions of the full field equations. We have also introduced a new limit, the weak-field approximation, which captures all the salient features of the solution, including the Yukawa decay and the Vainshtein crossover. Of course, the theory considered is believed to suffer from various pathologies, like a Boulware-Deser type of instabilities [5,29,30]. However, our investigations show for the first time that the Vainshtein mechanism can actually work in theories of massive gravity, some of which could be safe from the pathologies of NLPF theory. Several questions are however left unanswered. First, we have not been able to solve numerically the equation for dense objects (or even realistic stars). Indeed, when one increases the density of the object, the numerics becomes unstable and singularities are found to appear. It remains to understand if those singularities are physical or just numerical artifacts. Along the same line, it is not known if standard black holes can be recovered via the Vainshtein mechanism. Second, it is not clear if the solution we found is stable or not. The Boulware-Deser type of instability could show up if one would try to let the

solution be time dependent or to be reached by a dynamical process such as spherical collapse. These and other issues are left for future investigations.

### ACKNOWLEDGMENTS

We thank T. Damour for interesting discussions and suggestions. The work of E.B. was supported by the TRR 33 ‘‘The Dark Universe.’’

### APPENDIX A: DETAILS ON PERTURBATION THEORY

Here, we explain with some detail how one can obtain the series expansions (27) and (30), and how the former is related to the DL expansion (28). Our starting point is the system (8) rewritten as in (26).

Looking for an expansion of the form (22), and assuming such an expansion has been found up to order  $n - 1$  (with  $n \geq 1$ ), the  $n$ th order system is of the form

$$\begin{aligned} \frac{\lambda'_n}{z} + \frac{\lambda_n}{z^2} + \frac{1}{2}(\lambda_n + 3\mu_n + z\mu'_n) &= f_{t,n} - G_{tt,n}, \\ \frac{v'_n}{z} - \frac{\lambda_n}{z^2} - \frac{1}{2}(v_n + 2\mu_n) &= f_{R,n} - G_{RR,n}, \\ \frac{\lambda_n}{z^2} - \frac{v'_n}{2z} &= \frac{1}{z}f_{g,n}, \end{aligned} \quad (\text{A1})$$

where the functions  $f_{t,n}$ ,  $f_{R,n}$ ,  $f_{g,n}$ ,  $G_{tt,n}$ , and  $G_{RR,n}$  correspond, respectively, to the terms of order  $\epsilon^{n+1}$  of  $f_t$ ,  $f_R$ ,  $f_g$ ,  $G_{tt}$ , and  $G_{RR}$ . It is possible to reorganize the system (A1) to obtain the equivalent system

$$v'_n = \frac{2\lambda_n}{z} - 2f_{g,n}, \quad (\text{A2})$$

$$\begin{aligned} \lambda_n = -\frac{z^2\mu_n + z^3\mu'_n}{2+z^2} - \frac{1}{3(2+z^2)}[-2z^2f_{t,n} + 2z^3f'_{R,n} \\ + 2z^2G_{tt,n} - 2z^3G'_{RR,n} - 2z(2+z^2)f_{g,n} + 4z^3f'_{g,n}], \end{aligned} \quad (\text{A3})$$

$$\begin{aligned} L[\mu_n] = -2z(4+z^2)f_{t,n} + 2(2+z^2)f'_{t,n} \\ - (16+6z^2+z^4)f'_{R,n} - (4z+2z^3)f''_{R,n} \\ + 2z(4+z^2)G_{tt,n} - (4+2z^2)G'_{tt,n} \\ + (16+6z^2+z^4)G'_{RR,n} + (4z+2z^3)G''_{RR,n} \\ + \frac{(2+z^2)^2(4+z)}{z^2}f_{g,n} - \frac{16}{z}f'_{g,n} \\ - (8+4z^2)f''_{g,n}, \end{aligned} \quad (\text{A4})$$

where  $L$  stands for the linear differential operator

$$\begin{aligned} L[f] \equiv -3z(4+z^2)f(z) + 6(z^2+4)f'(z) \\ + 3z(z^2+2)f''(z). \end{aligned}$$

Let us first focus on Eq. (A4), which allows one to find  $\mu_n$ . This equation can be written in the schematic form

$$L[\mu_n] = F[\{\mu_{m<n}, \lambda_{m<n}, \nu_{m<n}\}], \quad (\text{A5})$$

where  $F$  gathers the nonlinear terms of the equation of order  $\epsilon^{n+1}$ , which only depend on the terms of order  $m < n$ . Note that the above discussion holds both for the large  $z$  and the small  $z$  expansions, since it is only based on the matching between terms of the same order in the parameter  $\epsilon$  on the left- and right-hand sides of Eqs. (26). To go further, let us consider the action of the linear operator  $L$  on one term of the series (27) for  $\mu_n$  of the form

$$\begin{aligned} M_{n,i,k} &\equiv \left(\frac{e^{-z}}{z^5}\right)^{n+1} z^{i+2}(z^5 \log z)^{n-k}, \\ &= e^{-(n+1)z} z^{i-3-5k} (\log z)^{n-k}, \end{aligned}$$

such that, with those notations, one is looking for an expansion of  $\mu_n$  in the form

$$\mu_n = \epsilon^{n+1} \sum_{i=0}^{\infty} \sum_{k=0}^n \mu_{n,i,k} M_{n,i,k}. \quad (\text{A6})$$

Henceforth, we will denote by  $S_{n,m}$  the set of formal series of the form

$$z^m \epsilon^{n+1} \left(\frac{e^{-z}}{z^5}\right)^{n+1} \sum_{i=0}^{\infty} \sum_{k=0}^n s_{n,i,k} z^i (z^5 \log z)^{n-k},$$

where  $s_{n,i,k}$  denotes real coefficients. With this notation, we see that the expansion for  $\mu_n$  given in Eqs. (27) and (A6) belongs to the set  $S_{n,2}$ , while the expansions for  $\lambda_n$  and  $\nu_n$  belong to the set  $S_{n,4}$ . It is easy to check that those expansions hold for  $n = 0$ , as can be seen from Eq. (24). One finds easily that

$$\begin{aligned} zL[M_{n,i,k}] = p_{n,i,k}M_{n,i,k} + q_{n,i,k}M_{n,i+1,k} + r_{n,i,k}M_{n,i+2,k} \\ + s_{n,i,k}M_{n,i+3,k} + t_{n,i,k}M_{n,i+4,k} + (k-n) \\ \times [a_{n,i,k}M_{n,i+5,k+1} + b_{n,i,k}M_{n,i+6,k+1} \\ + c_{n,i,k}M_{n,i+7,k+1} + d_{n,i,k}M_{n,i+8,k+1}] \\ + 3(k-n)(k-n+1)(2M_{n,i+10,k+2} \\ + M_{n,i+12,k+2}), \end{aligned} \quad (\text{A7})$$

where one has

$$\begin{aligned} p_{n,i,k} &= 6(i-5k)(i-5k-3), \\ a_{n,i,k} &= 6(3-2i+10k), \end{aligned}$$

and  $q_{n,i,k}$ ,  $r_{n,i,k}$ ,  $s_{n,i,k}$ , and  $t_{n,i,k}$  are quadratic polynomials of  $n$ ,  $i$ , and  $k$  with integer coefficients,  $b_{n,i,k}$ ,  $c_{n,i,k}$ , and  $d_{n,i,k}$  are linear polynomials of  $n$ ,  $i$ , and  $k$  with integer coefficients. The exact expressions of these polynomials will not be needed here (except for  $t_{n,i,n}$ , see below). It is then easy to see that the action of the linear operator  $L$  on any series



of  $S_{n,2}$  results in a new series element of  $S_{n,1}$ . At each order  $n$ , a necessary condition to find a solution for  $\mu_n$  in the above form (A6) is then that  $F[\{\mu_{m<n}, \lambda_{m<n}, \nu_{m<n}\}]$  can be written as an element of  $S_{n,1}$ . Let us show that this is indeed the case. To do so, let us first consider the different functions appearing on the right-hand side of Eq. (26). Those functions have different dependencies on  $\lambda$ ,  $\nu$ ,  $\mu$ , and  $z$  that can be read from Eqs. (8a), (8b), and (10a)–(10c) and are given as follows:

- (i) The functions  $f_{t,\geq 2}$  and  $f_{R,\geq 2}$  come from the mass term and, as a consequence, can only depend on  $\lambda$ ,  $\nu$ ,  $\mu$ , and  $z\mu'$ . This means that once expanded in terms of powers of  $\lambda$ ,  $\nu$ ,  $\mu$ , and  $z\mu'$ , the functions  $f_{t,\geq 2}$  and  $f_{R,\geq 2}$  are sums of terms of the form  $\lambda^{n_1} \nu^{n_2} \mu^{n_3} (z\mu')^{n_4}$ , with  $n_1, n_2$ , and  $n_3$  positive integers and  $n_4$  an integer verifying  $0 \leq n_4 \leq 2$ .
- (ii) The functions  $G_{tt,n}$  and  $G_{RR,n}$  come from the Einstein's tensor, in the gauge (7). Their expression is given by the left-hand side of Eqs. (8a) and (8b); one can then easily realize that the functions  $G_{tt,n}$  and  $G_{RR,n}$  can be written as an infinite sum of terms of the form  $\lambda^{n_1} \nu^{n_2} z^{-2}$ ,  $\lambda^{n_1} \nu^{n_2} \nu' z^{-1}$  or  $\lambda^{n_1} \nu^{n_2} \lambda' z^{-1}$ , with  $n_1$  and  $n_2$  positive integers.
- (iii) The function  $f_g$  corresponds to the Bianchi identity (8c), i.e. to the derivative of the mass term; it is an infinite sum of terms of the form  $z^{-1} \lambda^{n_1} \nu^{n_2} \mu^{n_3} (z\mu')^{n_4} (z\nu')^{n_5} (z\lambda')^{n_6} (z^2 \mu'')^{n_7}$ , where  $n_i$  ( $i = 1 \dots 7$ ) are positive integers and  $0 \leq n_4 \leq 3$ , and  $0 \leq n_i \leq 1$  for  $5 \leq i \leq 7$ .

It is then clear that, if one assumes that the series expansions hold at order  $\epsilon^k$ , with  $k < n$ , the right-hand side of Eq. (A4) can be written as a series of the form

$$\epsilon^{n+1} \sum_{i=i_{\min}}^{\infty} \sum_{k=k_{\min}}^n m_{n,i,k} M_{n,i,k} = (\epsilon e^z)^{n+1} \times \sum_{i=i_{\min}}^{\infty} \sum_{k=k_{\min}}^n m_{n,i,k} z^{i-3-5k} (\log z)^{n-k}. \quad (\text{A8})$$

It remains to show that  $k_{\min} = 0$  and  $i_{\min} = -1$  for the above series to be in the  $S_{n,1}$  set. It is clear that the first relation holds because the highest power of  $\log z$  that can possibly appear on the right-hand side of (A4) is necessarily strictly smaller than the power of  $\epsilon$ , i.e.,  $n + 1$ ; this is because it is so in each term of the series  $S_{n,m}$ . Let us now compute  $i_{\min}$ . To do so, we need to look at the terms on the right-hand side of (A4) which are the most singular power of  $z$  when  $z$  goes to zero. Keeping in mind that the series  $S_{n,2}$  are more singular when  $z \rightarrow 0$  than the series  $S_{n,4}$ , looking at the above described decompositions of the functions  $f_{t,\geq 2}$ ,  $f_{R,\geq 2}$ ,  $G_{tt,n}$ ,  $G_{RR,n}$ , and  $f_g$ , and inspecting Eq. (A4), one sees that the most singular power of  $z$  of the right-hand sides of this equation comes from the terms proportional to  $z^{-2} f_{g,n}$ ,  $z^{-1} f'_{g,n}$ , and  $f''_{g,n}$ ; and more specifically from products of  $\mu^{n_i}$ ,  $(z\mu')^{n_j}$ , and  $(z^2 \mu'')^{n_l}$

appearing there (with  $n_i, n_j$ , and  $n_l$  positive integers strictly smaller than  $n$ ). In those products, the piece containing the most singular power of  $z$  will be given by terms of the form<sup>8</sup>

$$z^{-3} \prod_{j=1}^{j=J} (\epsilon^{n_j+1} M_{n_j,0,k_j})^{N_j}$$

with  $N_j, n_j, k_j$ , and  $J$  positive integers verifying  $k_j \leq n_j < n, J \geq 2$  and

$$\sum_{j=1}^{j=J} (n_j + 1) N_j = n + 1, \quad (\text{A9})$$

where the above equality selects the order  $\epsilon^{n+1}$ . The above expression (A10) can be rewritten as

$$(\epsilon e^z)^{n+1} z^{-3 - \sum_{j=1}^{j=J} (3+5k_j) N_j} (\log z)^{n+1 - \sum_{j=1}^{j=J} (1+k_j) N_j}.$$

Defining then  $k$  and  $i$  as

$$k = \sum_{j=1}^{j=J} (1 + k_j) N_j - 1, \quad i = -5 + 2 \sum_{j=1}^{j=J} N_j,$$

the above expression reads now

$$(\epsilon e^z)^{n+1} z^{i-3-5k} (\log z)^{n-k}. \quad (\text{A10})$$

With the above definition,  $i$  takes its minimal value equal to  $-1$  for  $\sum_{j=1}^{j=J} N_j = 2$ , which is achieved for  $J = 2$  and  $N_1 = N_2 = 1$  [for which there is always a solution to Eq. (A9)]. This proves that the right-hand side of Eq. (A4) is in the  $S_{n,1}$  set, as announced. Moreover, taking now into account the logarithm, it is easy to see what are the most singular terms in the right-hand side of Eq. (A4) (as  $z$  goes to zero). Those are obtained by the terms (A10) with  $i = -1$  and  $k = n$ , which is achieved again under the same conditions as above:  $\sum_{j=1}^{j=J} N_j = 2, J = 2$ , and  $N_1 = N_2 = 1$  (with the additional condition that  $\forall j, k_j = n_j$ ). This shows in fact that those most singular terms are simply given by the series expansion of the decoupling limit. Indeed, taking this limit, the only terms which are left on the right-hand side of Eq. (A4) are those in  $f_g$  which are quadratic in  $\mu$ . The same reasoning also shows that the nonlinear pieces on the right-hand side of Eqs. (A2) and (A3) are, respectively, in the  $S_{n,3}$  and  $S_{n,4}$  families and have their most singular piece at small  $z$  given by the part of  $f_{g,n}$  which is quadratic in  $\mu$ .

This first shows that a (formal) solution of the system (A2)–(A4) can *a priori* be searched in the form of the series (27), but also that, provided such a solution can be found, a subseries of it which corresponds to the most singular

<sup>8</sup>The factor of  $z^{-3}$  in front of the expression below comes from a factor  $z^{-2}$  in front of  $f_{g,n}$  and another factor of  $z^{-1}$  inside the expression of  $f_{g,n}$ .

terms at small  $z$  will be given *exactly* by the series expansion obtained in the decoupling limit [Eq. (15)].

Let us now show explicitly how to obtain the coefficients  $\mu_{n,i,k}$ , and solve Eq. (A4) for  $\mu_n$ . It is first of interest to look at the expression (A7) for  $k = n$  which reads

$$\begin{aligned} zL[M_{n,i,n}] &= 6(i-5n)(i-5n-3)M_{n,i,n} + q_{n,i,n}M_{n,i+1,n} \\ &\quad + r_{n,i,n}M_{n,i+2,n} + s_{n,i,n}M_{n,i+3,n} \\ &\quad + (6n+3n^2)M_{n,i+4,n}. \end{aligned} \quad (\text{A11})$$

The necessity to incorporate logarithms in the series expansion for  $\mu_n$  clearly appears above. Indeed, imagine we look for an expansion without logarithms, i.e. at each order  $n$ , we look for an expansion for  $\mu_n$ ,  $\nu_n$ , and  $\lambda_n$  of the form

$$\begin{aligned} \mu_n &\sim \epsilon^{n+1} \sum_{i=0}^{\infty} m_{n,i} M_{n,i,n} \\ &= \epsilon^{n+1} \sum_{i=0}^{\infty} m_{n,i} e^{-(n+1)z} z^{i-3-5n}, \\ \nu_n &\sim \epsilon^{n+1} z^2 \sum_{i=0}^{\infty} n_{n,i} M_{n,i,n}, \\ \lambda_n &\sim \epsilon^{n+1} z^2 \sum_{i=0}^{\infty} l_{n,i} M_{n,i,n}. \end{aligned} \quad (\text{A12})$$

Note that this form is equivalent to the form (A6) for  $n = 0$ , and it is in agreement with the zero order solution given in Eq. (24), and we have proven above that those series must start from  $i = 0$ . Then, at fixed  $n$ , assuming the above form holds for terms of order  $\epsilon^k$  with  $k \leq n$ , and matching formally the terms proportional to  $z^{-1} \times M_{n,i,n}$  on the left-hand side and right-hand side of Eq. (A5), one can try to solve for the values of the coefficients  $m_{n,i}$ , starting from  $i = 0$  to increasing values of  $i$ . This is possible only until one reaches  $i = 5n$  where the coefficient in front of  $M_{n,i,n}$  in (A11) vanishes. This vanishing means that the coefficient in front of the term  $z^{-1} \times M_{n,5n,n}$  on the left-hand side of (A5) will not depend on  $\mu_{n,5n}$ . On the other hand, this coefficient will only depend on the coefficients  $\mu_{n,i}$  with  $i < 5n$  which are already determined by matching the terms  $M_{n,i,n}$ , with  $i < 5n$ , and one has thus no freedom left to match the term proportional to  $M_{n,5n,n}$  on the two sides of (A5). This means that it is not possible in general to find for  $\mu_n$  a series expansion of the above form (A12). Note that the coefficient in front of  $M_{n,i,n}$  in (A11) also vanishes for  $i = 5n + 3$ , and hence one has the same problem for finding the coefficient  $\mu_{n,5n+3}$ . Including logarithms however solves this problem and allows one to get a formal series solution in the family  $\mathcal{S}_{n,2}$  as we now show.

Indeed, let us fix  $n$  and see how one can compute the coefficients  $\mu_{n,i,k}$ . Let us first assume that all the

coefficients  $\mu_{n,i,k}$ , with  $k < K$  ( $K$  some chosen positive integer) and  $i < 5k$  are known, and see how to obtain the coefficients  $\mu_{n,i,K}$ , with  $i < 5K$ . This is immediate, indeed, the terms proportional to  $z^{-1} \times M_{n,i,K}$  (for some  $i < 5K$ ) on the left-hand side of Eq. (A5) depend only on the coefficients  $\mu_{n,i,k}$ , with  $k \leq K$  and  $i < 5k$ , as can be seen from Eq. (A7). Hence, by matching formally those terms on the left-hand side and right-hand side of Eq. (A5) one obtains a system of  $5K$  equations for the  $5K$  unknown coefficients  $\mu_{n,i,K}$ , with  $i < 5K$ . This system is invertible because, as follows from Eq. (A7) it is represented by a triangular matrix with diagonal elements, given by the  $p_{n,i,K}$ , which are nonvanishing (this is because we assumed  $i < 5K$ ). Hence one can solve for the  $\mu_{n,i,K}$ , and by recursion one obtains all coefficients  $\mu_{n,i,k}$ , with  $i < 5k$ . The coefficients  $\mu_{n,5k,k}$  can then be determined as follows:  $\mu_{n,0,0}$  is determined uniquely by matching the term proportional to  $z^{-1} \times M_{n,5,1}$  on the left-hand side and right-hand side of Eq. (A5). Indeed,  $\mu_{n,5,1}$  does not appear in this term, since  $p_{n,5,1}$  vanishes, but it is replaced there by  $\mu_{n,0,0}$  because  $a_{n,0,0}$  does not vanish. By recursion, one can obtain in a similar way all coefficients  $\mu_{n,5k,k}$ , with  $k < n$ ; indeed, one can check that the coefficients  $a_{n,i=5k,k} = 18$  and hence do not vanish, which means that at fixed  $K$ ,  $\mu_{n,5(K-1),K}$  is uniquely determined by looking at the term  $z^{-1} \times M_{n,5K,K}$ . The last coefficient  $\mu_{n,5n,n}$  is left undetermined and can be chosen at will being related to a zero mode, as we will see below. The coefficients  $\mu_{n,i,k}$ , with  $5k < i < 5k + 3$  are then determined in a similar manner as those with  $i < 5k$ . One then obtains the  $\mu_{n,i=5k+3,k}$  in way similar to the way one obtained the  $\mu_{n,5k,k}$  (checking again that  $a_{n,i=5k+3,k} = -18$  do not vanish). One is left with one arbitrary choice of  $\mu_{n,5n+3,n}$ , and the remaining coefficients  $\mu_{n,i,k}$ , with  $i > 5k + 3$  are obtained inverting an infinite triangular matrix (or a finite one, if one restricts oneself to some maximum value of  $i$ ). The two arbitrary constants to be fixed each correspond to one of the two independent zero modes of the operator  $L$ . One of these zero modes is  $\mu_0$  given in Eq. (24), and another independent zero mode is given by

$$\mu_{0,bis} = \frac{1}{z} \left( 1 + \frac{1}{z} + \frac{1}{z^2} \right) e^{-z} - \frac{1}{z} \left( 1 - \frac{1}{z} + \frac{1}{z^2} \right) e^z.$$

One can check that both functions  $\mu_0$  and  $\mu_{0,bis}$  can be expanded as in (A6) with all coefficients  $\mu_{n,i,k}$  vanishing but those with  $k = n$ , and  $i \geq 5n$  for  $\mu_0$  and those with  $k = n$ , and  $i \geq 5n + 3$  for  $\mu_{0,bis}$ . Note, as discussed in the main text, that the inclusions of such zero modes in the series expansions might be necessary to enforce the hierarchy  $\mu_n \ll \mu_{n-1}$ .

Having shown how one can obtain  $\mu_n$  in the announced form, it is easy to obtain the corresponding expansions for

$\lambda_n$  and  $\nu_n$  from Eqs. (A2) and (A3). There, there is no free constant left to be chosen.

Let us now turn to the asymptotic expansion (30) for large  $z$ . Following a similar reasoning as previously, it is first easy to see that, provided such an expansion holds up to order  $\epsilon^n$ , then the right-hand side of Eq. (A5) will be a series of the form

$$\epsilon^{n+1} e^{-(n+1)z} \sum_{i=-\infty}^{i=3} m_{n,i} z^i.$$

But one also sees from Eq. (A11) that this form is also the one resulting from the action of the linear operator  $L$

on the series for  $\mu_n$  given in (30). The expression (A11) then shows that one can solve uniquely for all the coefficients  $\mu_{n,i}$  starting from the largest values of  $i$ , and no free constants appear in this case because the coefficient  $t_{n,i,n} = 6n + 3n^2$  does not vanish. One can then easily obtain the expressions for  $\nu_n$  and  $\lambda_n$  in the form (30) with no extra free constant.

## APPENDIX B: SERIES INSIDE THE SOURCE

The coefficients of the series expansions (48) can be found solving order by order in  $x^2$ ; the first orders read

$$\begin{aligned} \mu(x) &= \mu_0^{(x)} + \frac{1}{20(2e^{\nu_0^{(x)}} + e^{\mu_0^{(x)}} - 3e^{\nu_0^{(x)} + \mu_0^{(x)}})} \left[ (3 + 4e^{-\nu_0^{(x)}} - 39e^{\nu_0^{(x)}} - 2e^{\nu_0^{(x)} - 2\mu_0^{(x)}} + 18e^{\nu_0^{(x)} - \mu_0^{(x)}} - 2e^{-\mu_0^{(x)}} - e^{\mu_0^{(x)}} - 5e^{\mu_0^{(x)} - \nu_0^{(x)}} \right. \\ &\quad \left. + 24e^{\nu_0^{(x)} + \mu_0^{(x)}}) m^2 + 2 \left( 3e^{\nu_0^{(x)} + \mu_0^{(x)}} \left( P_0^{(x)} - \frac{5}{x_\odot^3} \right) - 2e^{\nu_0^{(x)}} \left( P_0^{(x)} - \frac{3}{x_\odot^3} \right) + e^{\mu_0^{(x)}} \left( P_0^{(x)} + \frac{7}{x_\odot^3} \right) \right) \right] x^2 + \mathcal{O}(x^4), \\ \lambda(x) &= \frac{1}{4} \left( \frac{4}{x_\odot^3} - e^{-\nu_0^{(x)} - 2\mu_0^{(x)}} (-1 + e^{\mu_0^{(x)}}) (-e^{\nu_0^{(x)}} + e^{\mu_0^{(x)}} + 2e^{\nu_0^{(x)} + \mu_0^{(x)}}) m^2 \right) x^2 + \mathcal{O}(x^4), \\ \nu(x) &= \nu_0^{(x)} + \frac{1}{4} e^{-\nu_0^{(x)} - 2\mu_0^{(x)}} \left( -e^{\nu_0^{(x)}} m^2 + e^{\mu_0^{(x)}} m^2 - 2e^{2\mu_0^{(x)}} m^2 + 2e^{\nu_0^{(x)} + 2\mu_0^{(x)}} \left( m^2 + P_0^{(x)} + \frac{1}{x_\odot^3} \right) \right) x^2 + \mathcal{O}(x^4), \\ P(x) &= P_0^{(x)} - \frac{1}{8} \left( P_0^{(x)} + \frac{3}{x_\odot^3} \right) e^{-\nu_0^{(x)} - 2\mu_0^{(x)}} \left( -e^{\nu_0^{(x)}} m^2 + e^{\mu_0^{(x)}} m^2 - 2e^{2\mu_0^{(x)}} m^2 + 2e^{\nu_0^{(x)} + 2\mu_0^{(x)}} \left( m^2 + P_0^{(x)} + \frac{1}{x_\odot^3} \right) \right) x^2 + \mathcal{O}(x^4). \end{aligned}$$

## APPENDIX C: INFINITELY MANY SOLUTIONS AT INFINITY: A SERIES EXPANSION APPROACH

In this Appendix, we would like to complete the argument on the existence, at infinity, of infinitely many solutions of the system (8) having the same asymptotic behavior given by the linear solution (24). Our starting point is Sec. III B, where we assumed that there exists a solution  $\bar{\sigma} \equiv \{\bar{\lambda}, \bar{\nu}, \bar{\mu}\}$ , such that

$$\sigma \rightarrow \sigma_0 \quad \text{for } z \rightarrow \infty,$$

where  $\sigma_0 \equiv \{\lambda_0, \nu_0, \mu_0\}$ ; we looked for a solution  $\sigma$  as small perturbations of the known solution  $\bar{\sigma}$ :

$$\sigma = \bar{\sigma} + \delta\sigma,$$

where  $\delta\sigma \equiv \{\delta\lambda, \delta\nu, \delta\mu\}$ . We gave the leading behavior of these solutions at Eq. (40). In the following, we would like to go beyond this leading behavior analysis and find a consistent way of looking for the solution  $\delta\sigma$  through a series expansion.

### 1. Equations for $\delta\sigma$

The first step is to write down the equations for  $\delta\sigma$  that would generalize the leading order equations (32).

Schematically, taking into account (31), the system of equations (8) can be linearized around the solution  $\bar{\sigma}$ ,

$$\begin{aligned} \frac{\partial F_t}{\partial \sigma} \Big|_{\bar{\sigma}} \delta\sigma + \frac{\partial F_t}{\partial \sigma'} \Big|_{\bar{\sigma}} \delta\sigma' &= 0, \\ \frac{\partial F_R}{\partial \sigma} \Big|_{\bar{\sigma}} \delta\sigma + \frac{\partial F_R}{\partial \sigma'} \Big|_{\bar{\sigma}} \delta\sigma' &= 0, \\ \frac{\partial f_g}{\partial \sigma} \Big|_{\bar{\sigma}} \delta\sigma + \frac{\partial f_g}{\partial \sigma'} \Big|_{\bar{\sigma}} \delta\sigma' + \frac{\partial f_g}{\partial \mu''} \Big|_{\bar{\sigma}} \delta\mu'' &= 0. \end{aligned} \tag{C1}$$

It is important to note that in the linearized Bianchi identity, the last equation of (C1), we do not assume that the terms containing  $\delta\mu''$  are negligible, so we keep them. We are interested in asymptotic solutions at infinity; thus we can substitute  $\sigma_0$  instead of  $\bar{\sigma}$  in (C1), since  $\bar{\sigma} \rightarrow \sigma_0$  at  $z \rightarrow \infty$ .

Further, since we are considering a range of distances where the solution is deep in the linear regime, we can safely keep only the lowest order in powers of  $\sigma_0$ , while making an expansion of coefficients of Eqs. (C1). This leads to the equations,

$$\begin{aligned}
\frac{\delta\lambda'}{z} + \frac{\delta\lambda}{z^2} + \frac{1}{2}(\delta\lambda + 3\delta\mu + z\delta\mu') &= be^{-z} \left[ \left( \frac{1}{3z^3} - \frac{1}{z^2} - \frac{2}{3z} \right) \delta\lambda + \left( \frac{2}{z^2} + \frac{2}{3z} \right) \delta\lambda' + \frac{1}{3z} \delta\nu + \frac{1}{z} \delta\mu \right. \\
&\quad \left. + \left( -\frac{1}{2z^2} - \frac{1}{2z} + \frac{1}{3} + \frac{z}{6} \right) \delta\mu' \right], \\
\frac{\delta\nu'}{z} - \frac{\delta\lambda}{z^2} - \frac{1}{2}(\delta\nu + 2\delta\mu) &= be^{-z} \left[ \left( \frac{1}{z^3} + \frac{1}{z^2} \right) \delta\lambda + \left( \frac{2}{3z^3} + \frac{2}{3z^2} + \frac{4}{3z} + \frac{1}{3} \right) \delta\nu \right. \\
&\quad \left. + \left( -\frac{1}{3z} + \frac{2}{3} \right) \delta\mu + \left( -\frac{1}{3z^2} - \frac{1}{3z} \right) \delta\mu' \right], \\
\frac{\delta\nu'}{2z} - \frac{\delta\lambda}{z^2} &= \frac{be^{-z}}{z^2} \left[ \left( \frac{1}{z^3} + \frac{1}{z^2} - \frac{2}{3z} + \frac{1}{3} \right) \delta\lambda + \left( \frac{1}{3z^2} + \frac{1}{3z} \right) \delta\lambda' + \left( \frac{2}{3z} + \frac{2}{3} \right) \delta\nu + \left( -\frac{1}{z^2} - \frac{1}{z} - \frac{2}{3} \right) \delta\nu' \right. \\
&\quad \left. + \left( -\frac{2}{3z} - \frac{2}{3} \right) \delta\mu + \left( \frac{1}{3z^2} + \frac{1}{3z} - \frac{1}{3} \right) \delta\mu' + \left( \frac{1}{3z} + \frac{1}{3} \right) \delta\mu'' \right]. \tag{C2}
\end{aligned}$$

One can of course check that keeping only the leading behavior in the limit  $z \rightarrow \infty$  leads to the equations (32).

## 2. Ansatz for $\delta\sigma$

The next step is to adopt an appropriate ansatz, which can be chosen of the form

$$\begin{aligned}
\delta\mu &= F_\mu(z) \exp\left(\frac{z}{4} - f(z)e^{z/2}\right), \\
\delta\lambda &= -F_\lambda(z) \frac{z^2}{2} \exp\left(\frac{z}{4} - f(z)e^{z/2}\right), \tag{C3} \\
\delta\nu &= -F_\nu(z) \frac{\sqrt{b}}{3} \sqrt{1 + \frac{1}{z}} \exp\left(-\frac{z}{4} - f(z)e^{z/2}\right),
\end{aligned}$$

where the function  $f$  is assumed to satisfy the equation

$$f'(z) = \frac{3z^{3/2}}{2\sqrt{b(1+z)}} - \frac{f(z)}{2}. \tag{C4}$$

The form of this ansatz will be justified later. The formal solution of Eq. (C4) is given by

$$f(z) = e^{-z/2} \int_{z_0}^z \frac{3t^{3/2}}{2\sqrt{b(1+t)}} e^{t/2} dt.$$

The exact choice of the lower integral bound  $z_0$  does not matter so much, since it will just translate into an overall constant factor in Eqs. (C3). An asymptotic expansion of the solution can be easily found to be

$$f(z) = \frac{1}{\sqrt{b}} \left( 3z - \frac{15}{2} + \frac{9}{8z} + \dots \right). \tag{C5}$$

Note that the leading behavior of the ansatz (C3) matches the asymptotic behavior (40).

## 3. The equations for the $F_i$ 's

Once the ansatz (C3) has been chosen, we can compute the equations for the  $F_i$ 's. They read<sup>9</sup>

$$\begin{aligned}
E_t &\equiv E_t^{(0)} + (\sqrt{b}e^{-z/2})E_t^{(1)} + (be^{-z})E_t^{(2)} + (b^{3/2}e^{-3z/2})E_t^{(3)} + (b^2e^{-2z})E_t^{(4)} = 0, \\
E_R &\equiv E_R^{(0)} + (\sqrt{b}e^{-z/2})E_R^{(1)} + (be^{-z})E_R^{(2)} + (b^{3/2}e^{-3z/2})E_R^{(3)} = 0, \tag{C6} \\
E_B &\equiv E_B^{(0)} + (\sqrt{b}e^{-z/2})E_B^{(1)} + (be^{-z})E_B^{(2)} + (b^{3/2}e^{-3z/2})E_B^{(3)} = 0,
\end{aligned}$$

with

<sup>9</sup>Note that we have multiplied the  $tt$  equation by a factor  $\sqrt{b}e^{-z/2}$  in order to have consistent form among the three equations.



$$\begin{aligned}
E_t^{(0)}[F_i] &\equiv \frac{3z^{5/2}}{4\sqrt{z+1}}[F_\lambda(z) - F_\mu(z)], \\
E_t^{(1)}[F_i] &\equiv -\frac{1}{8}[2(z^2 + z + 12)F_\lambda(z) + 4zF'_\lambda(z) - (12 + z)F_\mu(z) - 4zF'_\mu(z)], \\
E_t^{(2)}[F_i] &\equiv \frac{1}{4\sqrt{z}\sqrt{z+1}}[-2z^2(z+3)F_\lambda(z) + (z^3 + 2z^2 - 3z - 3)F_\mu(z)], \\
E_t^{(3)}[F_i] &\equiv -\frac{1}{24z^2}[(6z^3 - 10z^2 - 52z)F_\lambda(z) - (8z^3 + 24z^2)F'_\lambda(z) + (z^3 + 2z^2 + 21z - 3)F_\mu(z) \\
&\quad + (4z^3 + 8z^2 - 12z - 12)F'_\mu(z)], \\
E_t^{(4)}[F_i] &\equiv \frac{1}{9z^2}\sqrt{z^2 + z}F_\nu(z), \\
E_R^{(0)}[F_i] &\equiv \frac{1}{2}[F_\lambda(z) + F_\nu(z) - 2F_\mu(z)], \\
E_R^{(1)}[F_i] &\equiv \frac{1}{12z^{5/2}\sqrt{z+1}}[(2z^3 + 3z^2 + z + 2)F_\nu(z) - 4(z^2 + z)F'_\nu(z) - 6(z^3 + z^2)F_\mu(z)], \\
E_R^{(2)}[F_i] &\equiv \frac{1}{12z^2}[6z(1+z)F_\lambda(z) + (1+5z-8z^2)F_\mu(z) + 4(1+z)F'_\mu(z)], \\
E_R^{(3)}[F_i] &\equiv \frac{1}{9z^{7/2}\sqrt{z+1}}[(z^4 + 5z^3 + 6z^2 + 4z + 2)F_\nu(z)], \\
E_B^{(0)}[F_i] &\equiv \frac{1}{4}[2F_\lambda(z) + F_\nu(z) - 3F_\mu(z)], \\
E_B^{(1)}[F_i] &\equiv -\frac{1}{24\sqrt{z^5}(z+1)}[-6(z^2 + z^2)F_\lambda(z) + (z^2 + z + 2)F_\nu(z) - 4(z + z^2)F'_\nu(z) + (36z + 30)F_\mu(z) \\
&\quad + 24(z + z^2)F'_\mu(z)], \\
E_B^{(2)}[F_i] &\equiv \frac{1}{48z^4}[2z(4z^3 - 7z^2 + 21z + 20)F_\lambda(z) + (8z^3 + 8z^2)F'_\lambda(z) + 8z(2z^2 + 3z + 3)F_\nu(z) \\
&\quad + (35z^2 + 27z - 4)F_\mu(z) + (8z^2 - 24z - 16)F'_\mu(z) - 16(z^2 - 16z)F''_\mu(z)], \\
E_B^{(3)}[F_i] &\equiv \frac{1}{36z^5\sqrt{z(z+1)}}[(10z^4 + 21z^3 + 18z^2 + 9z + 6)F_\nu(z) - 4z(2z^3 + 5z^2 + 6z + 3)F'_\nu(z)]. \tag{C7}
\end{aligned}$$

Note that we have used Eq. (C4) to eliminate  $f'$  and  $f''$ ; surprisingly enough, Eqs. (C7) do not contain  $f$  either.

#### 4. Expansion in powers of $\sqrt{b}e^{-z/2}$

The next step is to solve for the functions  $F_i$ . To do so, let us assume they are of the form

$$F_i(z) = \sum_{n=0}^{\infty} F_i^{(n)}(z)(\sqrt{b}e^{-z/2})^n \tag{C8}$$

with  $F_i^{(n)}(z)$  at most polynomial, i.e.

$$F_i^{(n)}(z) = \mathcal{O}(z^{P_i^{(n)}}), \tag{C9}$$

and solve Eqs. (C2) order by order in  $\sqrt{b}e^{-z/2}$ .

#### a. 0th order

*A priori*, out of the three equations (C6), one can get the three 0th order equations for  $F_i^{(0)}$

$$E_k^{(0)}[F_i^{(0)}] = 0,$$

where  $k = \{t, R, B\}$  and  $i = \{\lambda, \nu, \mu\}$ . However, only two of them are independent. To see this more algebraically, one can write the 0th order equations in a matrix form

$$\mathbf{M} \cdot \mathbf{F}^{(0)} = 0,$$

where

$$\mathbf{M} = \begin{pmatrix} \frac{3z^{5/2}}{4\sqrt{z+1}} & 0 & -\frac{3z^{5/2}}{4\sqrt{z+1}} \\ \frac{1}{2} & \frac{1}{2} & -1 \\ \frac{1}{2} & \frac{1}{4} & -\frac{3}{4} \end{pmatrix} \quad \text{and} \quad \mathbf{F}^{(0)} = \begin{pmatrix} F_\lambda^{(0)} \\ F_\nu^{(0)} \\ F_\mu^{(0)} \end{pmatrix}.$$

The vector

$$\mathbf{V}_{\text{Ker}} = \begin{pmatrix} \frac{2}{z} \\ \frac{3z^{3/2}}{\sqrt{z+1}} \\ -\frac{6z^{3/2}}{\sqrt{z+1}} \end{pmatrix}$$

belongs to the kernel of  $\mathbf{M}^\top$ , while the vectors

$$\mathbf{V}_\lambda = \begin{pmatrix} \frac{4\sqrt{z+1}}{3z^{5/2}} \\ 0 \\ 0 \end{pmatrix}, \quad \mathbf{V}_\nu = \begin{pmatrix} -\frac{4\sqrt{z+1}}{3z^{5/2}} \\ 2 \\ 0 \end{pmatrix}$$

can be used to get the equations

$$\mathbf{N} = \begin{pmatrix} \frac{\sqrt{b}}{4}z(f(z) + 2f'(z)) & 0 & -\frac{\sqrt{b}}{4}z(f(z) + 2f'(z)) \\ \frac{1}{2} & \frac{\sqrt{b}\sqrt{z(z+1)}(f(z)+2f'(z))}{6z^2} & -1 \\ \frac{1}{2} & \frac{\sqrt{b}\sqrt{z(z+1)}(f(z)+2f'(z))}{12z^2} & -\frac{b(z+1)(f(z)+2f'(z))^2}{12z^3} \end{pmatrix}.$$

In order to have  $\mathbf{N}$  of rank 2 (otherwise  $\mathbf{F}^{(0)}$  is trivially 0), it is necessary to impose that its determinant is 0. If in addition we require  $f$  to be positive, the condition  $\det \mathbf{N} = 0$  translates into the differential equation (C4).

To conclude, we have found that the 0th order equations are not enough to fix the  $F_i^{(0)}$ 's; however, they allow one to understand the origin of the condition (C4) and enforce  $F_\lambda^{(0)}(z) = F_\nu^{(0)}(z) = F_\mu^{(0)}(z)$ . We now have to go to next order to solve for  $F_\mu^{(0)}(z)$ .

### b. First order

The first order equations read

$$E_k^{(0)}[F_i^{(1)}] + E_k^{(1)}[F_i^{(0)}] = 0. \quad (\text{C11})$$

Let us first project this set of equations onto the Kernel vector  $\mathbf{V}_{\text{Ker}}$ , since we know this procedure will cancel the  $E_k^{(0)}[F_i^{(1)}]$  contribution:

$$\mathbf{V}_{\text{Ker}}^\top \cdot \mathbf{E}^{(1)}[F_i^{(0)}] = \frac{3\sqrt{b}}{2z(z+1)} [(6z+5)F_\mu^{(0)}(z) + 4z(z+1)F_\mu^{(0)'}(z)] = 0, \quad (\text{C12})$$

where we have used the relations (C10) to replace  $F_\lambda$ ,  $F_\nu$  by  $F_\mu$ . Equation (C12) can be straightforwardly solved, leading to

$$F_\lambda^{(0)}(z) = F_\nu^{(0)}(z) = F_\mu^{(0)}(z) = \frac{A}{z^{5/4}(1+z)^{1/4}}, \quad (\text{C13})$$

with  $A$  being an arbitrary integration constant. A few remarks can be made at this point. First, the arbitrariness of the constant  $A$  explicitly shows that out of a solution  $\bar{\sigma}$  of the system of equations (8), it is possible to find a family of infinite number of new solutions, with the same asymptotic

$$\begin{aligned} \mathbf{V}_\lambda^\top \cdot \mathbf{M} \cdot \mathbf{F}^{(0)} &= F_\lambda^{(0)}(z) - F_\mu^{(0)}(z) = 0, \\ \mathbf{V}_\nu^\top \cdot \mathbf{M} \cdot \mathbf{F}^{(0)} &= F_\nu^{(0)}(z) - F_\mu^{(0)}(z) = 0. \end{aligned} \quad (\text{C10})$$

The fact that  $\mathbf{M}$  is of rank 2 is a direct consequence of the ansatz (C3), and, in particular, of the differential equation (C4) satisfied by  $f$ . More precisely, if one does not assume anything on the function  $f$ , and then computes the 0th order equations out of the ansatz (C3), they are still of the homogeneous form  $\mathbf{N} \cdot \mathbf{F}^{(0)} = 0$ , but with the matrix  $\mathbf{N}$  given by

behavior at infinity. Second, we note that  $F_i^{(0)} = \mathcal{O}(z^{-3/2})$ , in agreement with the condition (C9).

Let us turn now to the two other first order equations. Projecting Eq. (C11) onto  $\mathbf{V}_\lambda$  and  $\mathbf{V}_\nu$  leads to

$$\begin{aligned} F_\lambda^{(1)}(z) - F_\mu^{(1)}(z) - \frac{1}{3}\sqrt{1 + \frac{1}{z}F_\mu^{(0)}(z)} &= 0, \\ F_\nu^{(1)}(z) - F_\mu^{(1)}(z) + \frac{-2z^3 - z^2 + z + 2}{6\sqrt{z^5(z+1)}}F_\mu^{(0)}(z) \\ - \frac{4z^2 + 4z}{6\sqrt{z^5(z+1)}}F_\mu^{(0)'}(z) &= 0. \end{aligned}$$

These equations allow one to replace  $F_\lambda^{(1)}$  and  $F_\nu^{(1)}$  by  $F_\mu^{(1)}$  in the equations of order 2. These order 2 equations will lead to an equation for  $F_\mu^{(1)}$ , and two other equations relating  $F_\lambda^{(2)}$  and  $F_\nu^{(2)}$  to  $F_\mu^{(2)}$ , and so on and so forth.

### c. nth order

The  $n$ th order equations read

$$\begin{aligned} E_k^{(0)}[F_i^{(n)}] + E_k^{(1)}[F_i^{(n-1)}] + E_k^{(2)}[F_i^{(n-2)}] \\ + E_k^{(3)}[F_i^{(n-3)}] + E_k^{(4)}[F_i^{(n-4)}] &= 0. \end{aligned}$$

By iteration, it is possible to prove that

$$F_\lambda^{(n)}(z) \sim F_\nu^{(n)}(z) \sim F_\mu^{(n)}(z) = \mathcal{O}(z^{-3/2+n}). \quad (\text{C14})$$

We can sketch briefly how the demonstration goes: after having solved the equation of order  $n$ , the functions  $F_i^{(n-1)}$  are assumed to be known, and  $F_\lambda^{(n)}$ ,  $F_\nu^{(n)}$  can be expressed as functions of  $\{F_\mu^{(n)}, F_i^{(n-1)}\}$ . Then, projecting the equations of order  $n+1$  onto  $\mathbf{V}_{\text{Ker}}$  allows one to find  $F_\mu^{(n)}$ , leading to the scaling (C14). Projecting onto  $\mathbf{V}_\lambda$  and  $\mathbf{V}_\nu$

relates  $F_\lambda^{(n+1)}$  and  $F_\nu^{(n+1)}$  to  $F_\mu^{(n+1)}$ , closing the iteration process.

To summarize the findings of this section, we can say that we have found a class of solutions for the system of equations (C2). These solutions are of the form (C3), where the functions  $F_i$  can be expanded into the series of Eq. (C8). There are an infinite number of such solutions. This shows that out of a solution  $\bar{\sigma}$  of the system of equations (8), it is possible to find an infinite family of new solutions, with the same asymptotic behavior at

infinity. This freedom is crucial to match the vacuum solution with the one with source.

## 5. Numerical validation

To conclude this study of the infinitely many solutions satisfying the same boundary conditions at infinity, we checked that our solution (C3) fits well the numerical solution of the system of equations (C2). An example is shown in Fig. 9.

- 
- [1] V. A. Rubakov and P. G. Tinyakov, *Phys. Usp.* **51**, 759 (2008).
- [2] M. Fierz, *Helv. Phys. Acta* **12**, 3 (1939); M. Fierz and W. Pauli, *Proc. R. Soc. A* **173**, 211 (1939).
- [3] H. van Dam and M. J. G. Veltman, *Nucl. Phys.* **B22**, 397 (1970); V. I. Zakharov, *JETP Lett.* **12**, 312 (1970); Y. Iwasaki, *Phys. Rev. D* **2**, 2255 (1970).
- [4] A. I. Vainshtein, *Phys. Lett.* **39B**, 393 (1972).
- [5] D. G. Boulware and S. Deser, *Phys. Rev. D* **6**, 3368 (1972).
- [6] G. R. Dvali, G. Gabadadze, and M. Porrati, *Phys. Lett. B* **485**, 208 (2000).
- [7] C. de Rham, G. Dvali, S. Hofmann, J. Khoury, O. Pujolas, M. Redi, and A. J. Tolley, *Phys. Rev. Lett.* **100**, 251603 (2008); C. de Rham, S. Hofmann, J. Khoury, and A. J. Tolley, *J. Cosmol. Astropart. Phys.* **02** (2008) 011; C. de Rham, J. Khoury, and A. J. Tolley, *Phys. Rev. D* **81**, 124027 (2010).
- [8] C. Deffayet, *Phys. Lett. B* **502**, 199 (2001); C. Deffayet, G. R. Dvali, and G. Gabadadze, *Phys. Rev. D* **65**, 044023 (2002).
- [9] K. Koyama, *Phys. Rev. D* **72**, 123511 (2005); D. Gorbunov, K. Koyama, and S. Sibiryakov, *Phys. Rev. D* **73**, 044016 (2006); C. Charmousis, R. Gregory, N. Kaloper, and A. Padilla, *J. High Energy Phys.* **10** (2006) 066; R. Gregory, N. Kaloper, R. C. Myers, and A. Padilla, *J. High Energy Phys.* **10** (2007) 069.
- [10] G. Dvali, [arXiv:hep-th/0402130](https://arxiv.org/abs/hep-th/0402130); K. Izumi, K. Koyama, O. Pujolas, and T. Tanaka, *Phys. Rev. D* **76**, 104041 (2007); K. Izumi and T. Tanaka, *Prog. Theor. Phys.* **121**, 419 (2009).
- [11] C. Deffayet, G. Gabadadze, and A. Iglesias, *J. Cosmol. Astropart. Phys.* **08** (2006) 012.
- [12] M. Minamitsuji, *Phys. Lett. B* **684**, 92 (2010); C. de Rham, J. Khoury, and A. J. Tolley, *Phys. Rev. Lett.* **103**, 161601 (2009); N. Agarwal, R. Bean, J. Khoury, and M. Trodden, *Phys. Rev. D* **81**, 084020 (2010).
- [13] N. Arkani-Hamed, S. Dimopoulos, G. Dvali, and G. Gabadadze, [arXiv:hep-th/0209227](https://arxiv.org/abs/hep-th/0209227); G. Dvali, S. Hofmann, and J. Khoury, *Phys. Rev. D* **76**, 084006 (2007).
- [14] A. Nicolis, R. Rattazzi, and E. Trincherini, *Phys. Rev. D* **79**, 064036 (2009); C. Deffayet, G. Esposito-Farese, and A. Vikman, *Phys. Rev. D* **79**, 084003 (2009); E. Babichev, C. Deffayet, and R. Ziour, *Int. J. Mod. Phys. D* **18**, 2147 (2009); C. Deffayet, S. Deser, and G. Esposito-Farese, *Phys. Rev. D* **80**, 064015 (2009); N. Chow and J. Khoury, [arXiv:0905.1325](https://arxiv.org/abs/0905.1325).
- [15] G. Gabadadze, *Phys. Lett. B* **681**, 89 (2009); C. de Rham, *Phys. Lett. B* **688**, 137 (2010).
- [16] C. de Rham and G. Gabadadze, [arXiv:1006.4367](https://arxiv.org/abs/1006.4367).
- [17] A. H. Chamseddine and V. Mukhanov, *J. High Energy Phys.* **08** (2010) 011.
- [18] P. Hořava, *Phys. Rev. D* **79**, 084008 (2009).
- [19] S. Mukohyama, *J. Cosmol. Astropart. Phys.* **09** (2009) 005.
- [20] T. Damour, I. I. Kogan, and A. Papazoglou, *Phys. Rev. D* **67**, 064009 (2003).
- [21] E. Babichev, C. Deffayet, and R. Ziour, *Phys. Rev. Lett.* **103**, 201102 (2009).
- [22] T. Damour, I. I. Kogan, and A. Papazoglou, *Phys. Rev. D* **66**, 104025 (2002).
- [23] T. Damour and I. I. Kogan, *Phys. Rev. D* **66**, 104024 (2002).
- [24] N. Arkani-Hamed, H. Georgi, and M. D. Schwartz, *Ann. Phys. (N.Y.)* **305**, 96 (2003).
- [25] A. Salam and J. A. Strathdee, *Phys. Rev. D* **16**, 2668 (1977); C. J. Isham and D. Storey, *Phys. Rev. D* **18**, 1047 (1978); D. Blas, C. Deffayet, and J. Garriga, *Phys. Rev. D* **76**, 104036 (2007); *Classical Quantum Gravity* **23**, 1697 (2006); Z. Berezhiani, D. Comelli, F. Nesti, and L. Pilo, *J. High Energy Phys.* **07** (2008) 130; *Phys. Rev. Lett.* **99**, 131101 (2007).
- [26] E. Babichev, C. Deffayet, and R. Ziour, *J. High Energy Phys.* **05** (2009) 098.
- [27] C. Aragone and J. Chela-Flores, *Nuovo Cimento Soc. Ital. Fis. A* **10**, 818 (1972).
- [28] W. Press, S. Teukolsky, W. Vetterling, and B. Flannery, *Numerical Recipes* (Cambridge University Press, Cambridge, England, 2007), 3rd ed.
- [29] C. Deffayet and J. W. Rombouts, *Phys. Rev. D* **72**, 044003 (2005).
- [30] P. Creminelli, A. Nicolis, M. Papucci, and E. Trincherini, *J. High Energy Phys.* **09** (2005) 003.