

Involute, minimal, outer, and increasingly trapped spheres

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(Received 13 September 2009; revised manuscript received 21 December 2009; published 26 January 2010)

Seven different refinements of trapped surfaces are proposed, each intended as potential stability conditions. This article concerns spherical symmetry, but each condition can be generalized. Involute trapped spheres satisfy a similar condition to minimal trapped spheres, which are strictly minimal with respect to the Kodama vector. There is also a weaker version of involute trapped spheres. Outer trapped spheres have positive surface gravity. Increasingly (future, respectively, past) trapped spheres generate spheres which are more trapped in a (future, respectively, past) causal direction, with three types: in any such causal direction, along the dual Kodama vector, and in some such causal direction. Assuming the null energy condition, the seven conditions form a strict hierarchy, in the above order. In static space-times, they reduce to three inequivalent definitions, namely, minimal, outer, and increasingly trapped spheres. For a widely considered class of so-called nice (or nondirty) black holes, minimal trapped and outer trapped become equivalent. Reissner-Nordström black holes provide examples of this, and that the increasingly trapped differs. Examples where all three refinements differ are provided by a simple family of dirty black holes parametrized by mass and singularity area.

DOI: [10.1103/PhysRevD.81.024037](https://doi.org/10.1103/PhysRevD.81.024037)

PACS numbers: 04.70.Bw, 02.40.Hw

I. INTRODUCTION

Trapped surfaces, as originally defined by Penrose [1], play an important role in gravitational physics, both for black holes and in cosmology. Such surfaces were crucial in the singularity theorems of Penrose and Hawking [1–3]. More recent years have seen the development of a local, dynamical theory of black holes in terms of marginal surfaces, a limit of trapped surfaces, including laws of black hole dynamics involving physical quantities such as mass and surface gravity [4–8].

However, a problem has become clear recently, as follows. Suppose one has a spherically symmetric space-time such as Vaidya, and considers the region covered by all trapped surfaces, not necessarily spherically symmetric. One might expect the boundary of this trapped region to consist of marginal surfaces, i.e., to be a trapping horizon. However, this is not so: trapped surfaces can poke through the trapping horizon [9–11]. The boundary must be spherically symmetric, but can be outside the trapping horizon. On the other hand, this boundary does not have the special physical properties that trapping horizons have, such as a first law involving surface gravity [12] and a local Hawking temperature [13]. There is thus a conflict between the mathematics and physics, which is something of a crisis for anyone hoping to understand black holes in a practical way.

In the author's view, the physics is clear and so the mathematics must yield. That is, trapped surfaces as simply defined need to be refined in some way to forbid the above behavior. Of course, the condition or conditions should be geometrically natural, and preferably with some physical content. The situation appears to be similar to that which used to hold for marginal surfaces, which are

too general to characterize black holes and needed to be refined, such as by the outer condition expressing positive surface gravity [8,12], or other stability conditions [14–17]. Such stability conditions are inequalities of one differential level higher than trapped or marginal itself.

This article proposes seven such refinements, of essentially three types with some variations, each of which holds for any trapped sphere in a Schwarzschild space-time. They are of some interest in themselves and turn out to be related, assuming the Einstein equation, or more exactly just the null energy condition (NEC). A companion article will deal with general cases [18], while this article will be restricted to spherical symmetry, with everything respecting spherical symmetry. One reason for making this special study is that the physical meaning of all terms which will appear is clear, while in general one must use the available geometrical quantities, which are one step removed from familiar physical quantities, such as mass m , surface gravity κ , and energy flux ψ . Thus while the companion article [18] is intended to be of more interest to mathematicians, this article is intended for those interested in more physical issues related to black holes which can be adequately addressed in spherical symmetry.

The article is organized as follows. Section II explains how any sphere is extremal in some normal direction, and defines minimal trapped spheres. Section III defines outer trapped spheres and shows that the condition is implied by minimal trapped spheres. Section IV defines increasingly trapped spheres of three kinds and shows that the conditions are implied by outer trapped spheres. Section V defines involute trapped spheres of two kinds and shows that they are outer trapped. Section VI turns to static cases and shows that the conditions reduce to only three inequivalent ones, namely, minimal, outer, and increasingly

trapped. Section VII studies so-called nice black holes, as opposed to so-called dirty black holes [19], and shows that minimal trapped spheres and outer trapped spheres become equivalent in such cases. Section VIII studies Reissner-Nordström black holes, finding the boundaries of the regions of the refined trapped spheres. Section IX does the same for a class of dirty black holes for which all three refinements differ. Section X concludes with hierarchy diagrams and remarks.

II. MINIMAL TRAPPED SPHERES

In terms of the area radius r , the area of the spheres of symmetry is

$$A = 4\pi r^2. \quad (1)$$

The dual Kodama vector is

$$k_* = g^{-1}(dr) \quad (2)$$

where g is the space-time metric. A sphere is said to be *untrapped*, *marginal* or *trapped* if k_* is, respectively, spatial, null or temporal [12,20,21]. If the space-time is time orientable and k_* is future (respectively, past) causal, then the sphere is said to be *future* (respectively, *past*) trapped or marginal.

There is a duality operation on normal vectors η , corresponding to the Hodge dual on 1-forms, defined by

$$g(\eta_*, \eta) = 0, \quad g(\eta_*, \eta_*) = -g(\eta, \eta). \quad (3)$$

Then k is the Kodama vector [12,20,22], defined up to sign, which can be locally fixed so that k is future causal in untrapped regions. Then a sphere is trapped, marginal, or untrapped if k is, respectively, spatial, null or temporal.

Now

$$k \cdot dr = 0 \quad (4)$$

which expresses that any sphere is extremal in the k direction. In the special case $k = 0$, the sphere is extremal in any normal direction, but otherwise k gives the unique such direction. Thus a trapped sphere is equivalently defined as a sphere which is extremal in a unique spatial normal direction. Then it is natural to ask whether the sphere is not merely extremal but minimal, as is the case for any trapped sphere in a Schwarzschild space-time.

Definition 1. A (strictly) *minimal* trapped sphere is a trapped sphere for which

$$k^a k^b \nabla_a \nabla_b r > 0 \quad (5)$$

where ∇ is the covariant derivative operator of g . Note that minimality itself requires only a nonstrict inequality, but the strict sign will turn out to be convenient. This definition has effectively been given by Maeda *et al.* in the context of cosmological wormholes [23].

III. OUTER TRAPPED SPHERES

Surface gravity was defined as [12,13,21]

$$\kappa = \frac{1}{2} * d * dr \quad (6)$$

where d is the exterior derivative and $*$ the Hodge dual in the normal space, i.e., $*d*dr$ is the normal wave operator, fixing the sign convention. Then a trapping horizon was said to be *outer*, *degenerate*, or *inner* if, respectively, $\kappa > 0$, $\kappa = 0$, or $\kappa < 0$. Examples of all types are provided by Reissner-Nordström solutions, where they correctly label the types of Killing horizon. This suggests extending the terminology to trapped spheres.

Definition 2. An *outer* trapped sphere is a trapped sphere for which

$$\kappa > 0. \quad (7)$$

Therefore spheres sufficiently close to an outer trapping horizon will be outer trapped, while those sufficiently close to an inner trapping horizon will not, for instance in a nondegenerate Reissner-Nordström solution. A trapped sphere with $\kappa < 0$ may similarly be called *inner* trapped.

Lemma 1. Assuming the Einstein equation with units $G = 1$,

$$k^a k^b \nabla_a \nabla_b r = \left(\frac{2m}{r} - 1 \right) \kappa - 4\pi r k_* \cdot \psi \quad (8)$$

where m is the mass [12,20,21,24],

$$\frac{2m}{r} - 1 = g(k, k) \quad (9)$$

and ψ is an energy flux defined by [12,21]

$$\psi = T \cdot k_* + w dr \quad (10)$$

where T is the energy tensor and w is an energy density:

$$w = -\frac{1}{2} \text{tr} T \quad (11)$$

where the trace is in the normal space.

Proof. In dual-null coordinates x^\pm , the metric has the form

$$ds^2 = r^2 d\Omega^2 - 2e^{2\varphi} dx^+ dx^- \quad (12)$$

where $d\Omega^2$ refers to the unit sphere and (r, φ) are functions of (x^+, x^-) . In these coordinates, trapped spheres are equivalently defined by $\partial_+ r \partial_- r > 0$, with $\partial_\pm r < 0$ (respectively, > 0) for future (respectively, past) trapped spheres, assuming that ∂_\pm are future pointing. Then one finds the explicit expressions

$$dr = \partial_+ r dx^+ + \partial_- r dx^- \quad (13)$$

$$k_* = -e^{-2\varphi} (\partial_+ r \partial_- + \partial_- r \partial_+) \quad (14)$$

$$k = e^{-2\varphi} (\partial_+ r \partial_- - \partial_- r \partial_+) \quad (15)$$

$$2m/r - 1 = 2e^{-2\varphi} \partial_+ r \partial_- r \quad (16)$$

$$\kappa = -e^{-2\varphi} \partial_+ \partial_- r \quad (17)$$

$$w = e^{-2\varphi} T_{+-} \quad (18)$$

$$\psi = -e^{-2\varphi} (T_{++} \partial_- r dx^+ + T_{--} \partial_+ r dx^-) \quad (19)$$

which implies

$$k_* \cdot \psi = e^{-4\varphi} (T_{--} (\partial_+ r)^2 + T_{++} (\partial_- r)^2). \quad (20)$$

The only relevant nonzero connection coefficients are $\Gamma_{\pm\pm}^{\pm} = 2\partial_{\pm}\varphi$. Then

$$\begin{aligned} k^a k^b \nabla_a \nabla_b r &= k^a k^b (\partial_a \partial_b r - \Gamma_{ab}^c \partial_c r) \\ &= e^{-4\varphi} ((\partial_- r)^2 (\partial_+ \partial_+ r - 2\partial_+ \varphi \partial_+ r) \\ &\quad + (\partial_+ r)^2 (\partial_- \partial_- r - 2\partial_- \varphi \partial_- r) \\ &\quad - 2\partial_+ r \partial_- r \partial_+ \partial_- r). \end{aligned} \quad (21)$$

The null-null components of the Einstein equation are

$$\partial_{\pm} \partial_{\pm} r - 2\partial_{\pm} \varphi \partial_{\pm} r = -4\pi r T_{\pm\pm}. \quad (22)$$

Then straightforward calculation using (16), (17), and (20).

Proposition 1. *NEC and minimal trapped implies outer trapped.*

Proof. $NEC \Rightarrow k_* \cdot \psi \geq 0$, as is most easily seen from $T_{\pm\pm} \geq 0$ and (20). For a trapped sphere, $r < 2m$, then inspect signs in (5), (7), and (8).

IV. INCREASINGLY TRAPPED SPHERES

Noting that $g(k, k) = 2m/r - 1$ (9) vanishes for marginal spheres and is positive for trapped spheres, it can be taken as a measure of how trapped a sphere is. The idea then is to ask whether it is increasing to the future (respectively, past) for a future (respectively, past) trapped sphere. Consideration of general cases [18] suggests instead the measure $(2m/r - 1)/r^2$, which yields stricter conditions. Three different definitions can be given, as follows.

Definition 3. An *increasingly* trapped sphere is a trapped sphere for which

$$k_* \cdot d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) > 0. \quad (23)$$

Definition 4. An *anyhow increasingly* trapped sphere is a future (respectively, past) trapped sphere for which, for all future (respectively, past) causal normal vectors ζ ,

$$\zeta \cdot d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) > 0. \quad (24)$$

Definition 5. A *somehow increasingly* trapped sphere is a future (respectively, past) trapped sphere for which, for some future (respectively, past) causal normal vector ζ ,

$$\zeta \cdot d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) > 0. \quad (25)$$

Clearly anyhow increasingly trapped implies increasingly trapped, which implies somehow increasingly trapped.

Lemma 2. Assuming the Einstein equation,

$$\begin{aligned} r^2 \zeta \cdot d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) &= 8\pi r \zeta \cdot \psi - 2\left(\kappa + \frac{1}{r} \left(\frac{2m}{r} - 1\right)\right) \\ &\quad \times \zeta \cdot dr. \end{aligned} \quad (26)$$

Proof. The last term above comes from the $1/r^2$ term, so it suffices to calculate

$$\begin{aligned} \zeta \cdot d(m/r) &= (\zeta^+ \partial_+ + \zeta^- \partial_-) (e^{-2\varphi} \partial_+ r \partial_- r) \\ &= e^{-2\varphi} (\zeta^+ \partial_- r (\partial_+ \partial_+ r - 2\partial_+ \varphi \partial_+ r) \\ &\quad \times \zeta^- \partial_+ r (\partial_- \partial_- r - 2\partial_- \varphi \partial_- r) \\ &\quad + (\zeta^+ \partial_+ r + \zeta^- \partial_- r) \partial_+ \partial_- r) \\ &= 4\pi r \zeta \cdot \psi - \kappa \zeta \cdot dr \end{aligned} \quad (27)$$

using the Einstein equations (22) as before and expressions (17) and (19).

Proposition 2. *NEC and outer trapped implies anyhow increasingly trapped.*

Proof. $r < 2m$, $\zeta \cdot dr < 0$ for ζ in the given causal quadrant, and $NEC \Rightarrow \zeta \cdot \psi \geq 0$, then inspect signs in (7) and (26).

If one wished to show directly that the sphere was merely increasingly trapped, the corresponding expression would be

$$k_* \cdot d\left(\frac{m}{r}\right) = 4\pi r k_* \cdot \psi + \left(\frac{2m}{r} - 1\right) \kappa. \quad (28)$$

This result will survive in general cases, while the above result will require a stricter definition of outer trapped [18], which in turn suggests a stricter condition than minimal trapped, described below.

V. INVOLUTE TRAPPED SPHERES

The minimality condition (5) can be rewritten as

$$k_*^a k^b \nabla_b k_a < 0 \quad (29)$$

due to orthogonality of k_* and k . This form is useful for general calculations, and suggests further definitions.

Definition 6. An *involute* trapped sphere is a future (respectively, past) trapped sphere for which, for all future (respectively, past) causal normal vectors ζ ,

$$\zeta^a k^b \nabla_b k_a < 0. \quad (30)$$

Definition 7. A *somehow involute* trapped sphere is a future (respectively, past) trapped sphere for which, for some future (respectively, past) causal normal vector ζ ,

$$\zeta^a k^b \nabla_b k_a < 0. \quad (31)$$

Clearly involute trapped implies minimal trapped, which implies somehow involute trapped.

Lemma 3. Assuming the Einstein equation,

$$\zeta^a k^b \nabla_b k_a = \kappa \zeta \cdot dr + 4\pi r \zeta \cdot \psi. \quad (32)$$

Proof.

$$\begin{aligned} \zeta^a k^b \nabla_b k_a &= \zeta^a k^b (\partial_b k_a - \Gamma_{ba}^c k_c) \\ &= e^{-2\varphi} (\zeta^+ \partial_- r (\partial_+ \partial_+ r - 2\partial_+ \varphi \partial_+ r) \\ &\quad + \zeta^- \partial_+ r (\partial_- \partial_- r - 2\partial_- \varphi \partial_- r) \\ &\quad - (\zeta^+ \partial_+ r + \zeta^- \partial_- r) \partial_+ \partial_- r) \\ &= 4\pi r \zeta \cdot \psi + \kappa \zeta \cdot dr \end{aligned} \quad (33)$$

using expressions (17) and (19) and the Einstein equations (22) as before.

Proposition 3. *NEC and somehow involute trapped implies outer trapped.*

Proof. As before, $\zeta \cdot dr < 0$ and $\text{NEC} \Rightarrow \zeta \cdot \psi \geq 0$, then inspect signs in (31) and (32).

VI. STATIC SPACE-TIMES

In static space-times, the static Killing vector ∂_t is proportional to k . Thus $k \cdot df = 0$ for any function f .

Proposition 4. *In static space-times, the three increasingly trapped conditions 3–5 are equivalent.*

Proof. Recall that k and k_* are orthogonal, and span the normal space for a trapped sphere. Any normal vector ζ can then be written as a linear combination of k and k_* , with positive component along k_* for the appropriate causal quadrant. Since

$$k \cdot d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) = 0 \quad (34)$$

the result follows.

Proposition 5. *In static space-times, the minimal trapped condition 1 and the two involute trapped conditions 6–7 are equivalent.*

Proof. As above, since

$$2k^a k^b \nabla_b k_a = k^b \nabla_b (k^a k_a) = 0. \quad (35)$$

Thus the seven (or eight in general) conditions reduce to three, definitions 1–3, which can be seen to be inequivalent in the examples to follow.

VII. NICE BLACK HOLES

Definition 8. A *nice* black hole is a static black hole for which

$$g(\partial_t, \partial_t) = -g^{-1}(dr, dr) \quad (36)$$

for some thereby standard choice of the scaling of the static Killing vector ∂_t . This includes Schwarzschild, Reissner-

Nordström (de Sitter), and many other black holes which have been considered in the literature. A non-nice static black hole is here called a *dirty* black hole, more or less consistently with previous usage [19].

For nice black holes, the Kodama and Killing vectors coincide, $k = \partial_t$, and the metric may be written, except at trapping horizons, as

$$ds^2 = r^2 d\Omega^2 + \left(1 - \frac{2m}{r}\right)^{-1} dr^2 - \left(1 - \frac{2m}{r}\right) dt^2. \quad (37)$$

Note that this is valid both outside the horizon, $r > 2m$, and inside, $r < 2m$.

Proposition 6. *For nice black holes, minimal trapped spheres and outer trapped spheres are equivalent.*

Proof. For a static black hole, one finds the energy density (11)

$$w = \frac{1}{2}(\rho - p) \quad (38)$$

where $\rho = -T_t^t$ is the energy density and $p = T_r^r$ the radial pressure. The energy flux (10) is

$$\psi = \frac{1}{2}(\rho + p)dr. \quad (39)$$

Using (8) one sees that the result will follow if

$$\rho + p = 0. \quad (40)$$

This is a known property of nice black holes, e.g., [25], one derivation going as follows. Any spherically symmetric metric can locally be written in the Regge-Wheeler form

$$ds^2 = r^2 d\Omega^2 + e^{2\varphi} (dr_*^2 - dt^2) \quad (41)$$

where (r, φ) are functions of (t, r_*) and φ reduces to the Newtonian potential in the Newtonian limit, if t reduces to Newtonian time. In static cases, the metric can also be written as

$$ds^2 = r^2 d\Omega^2 + \left(1 - \frac{2m}{r}\right)^{-1} dr^2 - e^{2\varphi} dt^2 \quad (42)$$

for $\varphi(r)$, $m(r)$, where

$$1 - \frac{2m}{r} = (e^{-\varphi} r')^2 \quad (43)$$

and the prime denotes differentiation with respect to r_* . A nice black hole is defined by

$$e^{2\varphi} = 1 - \frac{2m}{r} \quad (44)$$

which gives

$$r' = e^{2\varphi}. \quad (45)$$

One component of the Einstein equations is

$$4\pi r(\rho + p) = -(e^{-2\varphi} r')' \quad (46)$$

which gives the result.

VIII. REISSNER-NORDSTRÖM BLACK HOLES

For a Schwarzschild black hole, all the trapped spheres satisfy all the refinements. The simplest standard example where this is not so is provided by nonextreme Reissner-Nordström black holes, which are nice black holes (37) with

$$m = M - \frac{Q^2}{2r} \quad (47)$$

where M is the mass and Q the charge, being constants satisfying

$$M^2 > Q^2. \quad (48)$$

As is well known, there are both outer and inner trapping horizons, located at

$$r_{\pm} = M \pm \sqrt{M^2 - Q^2} \quad (49)$$

with trapped spheres for

$$r_- < r < r_+. \quad (50)$$

Since there are inner horizons, one knows in advance that trapped spheres close to them will be inner rather than outer trapped spheres. Explicitly, one may calculate surface gravity (6) using the general static formula

$$\kappa = \frac{1}{2\sqrt{|\gamma|}} \partial_r (\sqrt{|\gamma|} \gamma^{rr}) \quad (51)$$

where γ is the normal metric, to find

$$\kappa = \frac{M}{r^2} - \frac{Q^2}{r^3} \quad (52)$$

for which the sign switches at

$$r_o = \frac{Q^2}{M}. \quad (53)$$

Then $\kappa > 0$ for $r > r_o$ to satisfy definition 2, while $\kappa < 0$ for $r < r_o$. Elementary analysis using (48) shows that

$$r_- < r_o < r_+ \quad (54)$$

so that the switch occurs outside the inner horizon as expected.

Regarding increasingly trapped spheres, one has

$$d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) = 2 \frac{2Q^2 - 3Mr + r^2}{r^5} dr \quad (55)$$

which switches sign at

$$R_{\pm} = \frac{3}{2} \left(M \pm \sqrt{M^2 - \frac{8}{9} Q^2} \right). \quad (56)$$

Since $k_* \cdot dr = 1 - 2m/r < 0$ for trapped surfaces, the desired sign in definition 3 then occurs for $R_- < r < R_+$. Elementary analysis using (48) shows that

$$R_+ > r_+ \quad (57)$$

so that all trapped spheres sufficiently close to the outer horizon are increasingly trapped, while

$$r_- < R_- < r_o \quad (58)$$

so the switch occurs inside the inner trapped region.

In summary, as one starts at the outer horizon $r = r_+$ and moves inwards, the spheres are both minimal and outer trapped (by proposition 6) down to $r = r_o$ but not beyond, while still being increasingly trapped down to $r = R_-$ but not beyond, which is still outside the inner horizon $r = r_-$.

IX. DIRTY BLACK HOLES

Simple examples where all three refinements differ are provided by

$$ds^2 = r^2 d\Omega^2 + \left(1 - \frac{2M}{r}\right)^{-1} dr^2 - \left(1 - \frac{2M}{r}\right) \left(1 - \frac{L}{r}\right)^2 dt^2 \quad (59)$$

where the constants (M, L) are here assumed to satisfy

$$2M > L > 0. \quad (60)$$

There is an outer trapping horizon at $r = 2M$ and a spatial singularity at $r = L$, with trapped spheres for

$$L < r < 2M. \quad (61)$$

So the global structure is similar to that of a Schwarzschild black hole, except that the singularity has nonzero area. These are evidently dirty black holes with mass M and singularity area $4\pi L^2$. For $M \gg L$, they will look quite like Schwarzschild black holes from outside.

The analysis is similar to the above. Regarding increasingly trapped spheres, one first finds

$$m = M \quad (62)$$

then

$$d\left(\frac{1}{r^2} \left(\frac{2m}{r} - 1\right)\right) = 2 \frac{r - 3M}{r^4} dr \quad (63)$$

which gives the desired sign for $R < 3M$, by definition 3. Thus all the trapped spheres are increasingly trapped.

Regarding outer trapped spheres, one finds using (51) that

$$\kappa = \frac{(2M + L)r - 4ML}{r^2(r - L)} \quad (64)$$

for which the sign switches at

$$r_o = \frac{4ML}{2M + L}. \quad (65)$$

Then $\kappa > 0$ for $r > r_o$ to satisfy definition 2, while $\kappa < 0$ for $r < r_o$, as before. Elementary analysis using (60) shows that

$$L < r_o < 2M. \tag{66}$$

Thus there are inner trapped spheres near the singularity, in distinction to the Schwarzschild case.

Regarding minimal trapped spheres, one can use the Regge-Wheeler form (41) with

$$e^{2\varphi} = \left(1 - \frac{2M}{r}\right)\left(1 - \frac{L}{r}\right)^2 \tag{67}$$

where there is no problem with regarding φ as imaginary inside the horizon, as it will cancel out in the result. In static cases, one generally has

$$k^a k^b \nabla_a \nabla_b r = -k^t k^t \Gamma_{tt}^* \partial_{*r} \tag{68}$$

in terms of the connection coefficient

$$\Gamma_{tt}^* = \partial_{*r} \varphi. \tag{69}$$

In this case one finds

$$\partial_{*r} = \left(1 - \frac{2M}{r}\right)\left(1 - \frac{L}{r}\right) \tag{70}$$

and

$$k = \left(1 - \frac{L}{r}\right)^{-1} \partial_t \tag{71}$$

so

$$k^a k^b \nabla_a \nabla_b r = -\frac{1}{2} \left(1 - \frac{L}{r}\right)^{-2} \left(1 - \frac{2M}{r}\right) \partial_r (e^{2\varphi}) \tag{72}$$

where

$$\partial_r (e^{2\varphi}) = 2 \frac{(M+L)r^2 - (4M+L)Lr + 3ML^2}{r^4} \tag{73}$$

which gives roots

$$r_m = \frac{3ML}{M+L}, \quad L. \tag{74}$$

The desired sign in definition 1 occurs for $r > r_m$. Elementary analysis using (60) shows that

$$r_o < r_m < 2M \tag{75}$$

so the switch occurs in the outer trapped region.

In summary, as one starts at the horizon $r = 2M$ and moves inwards, the spheres are minimal trapped down to $r = r_m$ but not beyond, outer trapped down to $r = r_o$ but not beyond, while still being increasingly trapped all the way to the singularity $r = L$.

For completeness, one may note that the energy density vanishes,

$$\rho = 0 \tag{76}$$

which is a general property if m is constant, and that the radial pressure is

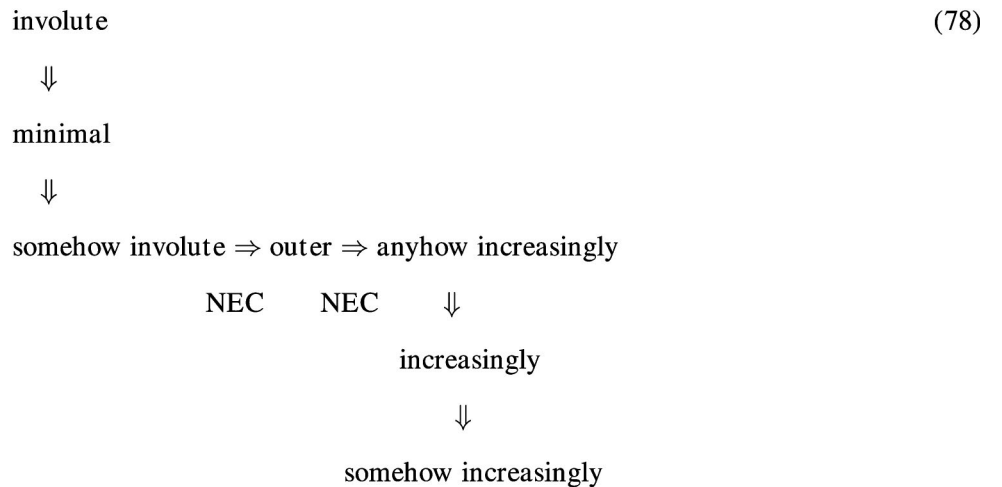
$$p = \frac{(M+L)(r-r_m)}{2\pi r^3(r-L)}. \tag{77}$$

The dominant energy condition requires $\rho \geq p$, which is violated for $r > r_m$. On the other hand, it is satisfied for $r \leq r_m$ and so the hierarchy of trapped surfaces is the standard one according to propositions 1–2.

Since p is rapidly vanishing as $r \rightarrow \infty$, $O(r^{-3})$, but infinite as $r \rightarrow L$, one can interpret it as coming from some modification of Einstein gravity at small length scales. The scale is given by L , which might be the Planck length or some other theoretical scale.

X. SUMMARY

The hierarchy of trapped spheres is illustrated as follows:



where the vertical implications are straightforward, while the horizontal implications require NEC. In general cases, the vertical hierarchy will remain, but the horizontal hierarchy will change and be interwoven [18].

In static cases, the vertical hierarchy collapses to leave just

$$\begin{array}{ccc} \text{minimal} & \Rightarrow & \text{outer} & \Rightarrow & \text{increasingly} & & (79) \\ & & \text{NEC} & & \text{NEC} & & \end{array}$$

and the first two become equivalent for the widely considered subclass of nice (nondirty) black holes. Examples have also been given of dirty black holes where all three refinements differ. The examples show that the refinements can still hold sufficiently close to the outer horizon of a black hole, even if it is dirty and violates energy conditions. On the other hand, the refinements need not hold near an inner horizon or a singularity. This provides some support for the idea to restrict trapped surfaces to those which are regular enough near an outer trapping horizon that the horizon is indeed the boundary of the region of such refined trapped surfaces.

Finally, while geometrical ideas play a large part in the above definitions and results, recall that the original motivation was a physical one, namely, to have a practical definition of black holes and a comprehensive theory. One might therefore argue that perhaps the most physically well motivated refinement is outer trapped, since it expresses positivity of surface gravity κ , which appears in a first law for trapping horizons [12], determines a local Hawking temperature $\kappa/2\pi$ [13], and has other physically relevant properties [12,21]. In any case, the interplay of geometry and physics remains a fascinating aspect of black holes, 95 years after the Schwarzschild solution was found.

ACKNOWLEDGMENTS

Research supported by the National Natural Science Foundation of China under Grants No. 10375081, No. 10473007, and No. 10771140, by the Shanghai Municipal Education Commission under Grant No. 06DZ111, and by the Shanghai Normal University under Grant No. PL609.

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