

Gravity-mediated decoherence

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A small quantum system within the gravitational field of a massive body will be entangled with the quantum degrees of freedom of the latter. Hence, the massive body acts as an environment, and it induces nonunitary dynamics, noise, and decoherence to the quantum system. It is impossible to shield systems on Earth from this gravity-mediated decoherence, which could severely affect all experiments with macroscopic quantum systems. We undertake a first-principles analysis of this effect, by deriving the corresponding open system dynamics. We find that near-future quantum experiments are not affected, but there is a strong decoherence effect at the human scale. The decoherence time for a superposition of two localized states of a human with a one meter separation is of the order of one second.

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I. INTRODUCTION

Contrasting the existence of many hypotheses or theories about a quantum theory of gravity, the experimental information about the interplay of gravity and quantum theory is surprisingly sparse, even in the nonrelativistic, weak-field regime. The famous Colella-Overhauser-Werner (COW) experiment [1] established that the effects of a background gravitational field on nonrelativistic particles are accounted by the addition of a potential term in the Hamiltonian operator. Later experiments on neutrons bouncing off a horizontal mirror [2] demonstrated the existence of bound states due to the gravitational field.

There is as yet no experimental test of the gravitational interaction between two different quantum matter distributions. In classical physics, gravity is universal; that is, it affects all bodies. It is always attractive, so it is impossible to shield any body from its effects. Furthermore, in the nonrelativistic weak-field limit, gravity is nondynamical. It is described solely by the gravitational potential, which is completely slaved to the mass density through Poisson's equation.

Taking the gravitational potential to be slaved to the mass density also for quantum systems is perhaps the most conservative assumption about the relation of gravity and quantum theory in the weak-field regime. Nonetheless, it has profound implications. It implies the possibility of gravitational Schrödinger-cat states [3], that is, of measurable superpositions of the gravitational force. It also implies that the gravitational interaction may induce quantum correlations, such as entanglement [4–7], that are experimentally accessible.

This means that a small quantum particle within the gravitational field of a massive body—for example, Earth—is entangled with the quantum degrees of freedom of the latter. Hence, when considering the reduced dynamics of the small particle, the massive body plays the role of an environment, and it leads to nonunitary dynamics, noise, and decoherence.

In this paper, we undertake a first-principles analysis of open system dynamics and decoherence for a particle in the gravitational field of a heavy body. We call this type of decoherence “gravity mediated,” rather than “gravitational,” because it originates from the quantum fluctuations of matter. Gravity plays the role of the transmission channel for those fluctuations, unlike gravitational decoherence models, in which the source of decoherence is the gravitational field itself [8,9].

The main motivation for our analysis is that gravity-mediated decoherence affects any system inside the gravitational field of the Earth. It sets an upper limit to the size of any macroscopic superpositions that can be created in a terrestrial laboratory. The existence of such a limit is unavoidable, but its value cannot be estimated with simple arguments: a detailed analysis and modeling is necessary. There is a good *a priori* possibility that this limit would affect many proposed experiments that involve macroscopic superpositions. This includes tests of dynamical reduction theories [10,11], tests of fundamental/gravitational decoherence [8,9], and generation of gravity-induced effects, such as entanglement [4–7]. It is therefore essential to have a quantitative estimate of the strength of gravity-mediated decoherence.

We undertake a first-principles analysis of the effect, and we evaluate the resulting decoherence rate. For Earth, this rate is small; it will not affect currently proposed experiments on macroscopic quantum systems. However, it is not

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negligible. Gravity-mediated decoherence effects become significant when quantum superpositions reach the scale of humans (or of cats). The decoherence time for a superposition of two localized states with mean separation of one meter for a human is of the order of one second. Macroscopic superpositions of significantly heavier bodies are not possible on the surface of Earth. However, the decoherence rate drops as d^{-3} , where d is a quantum particle's distance from the surface of Earth. The limits set by gravity-mediated decoherence are not fundamental but practical.

To obtain these results, we derived the open system dynamics for a particle inside the gravitational field generated by a large body. This dynamics is well approximated by a quantum Brownian motion (QBM) model [12–15], which is exactly solvable.

Gravity-mediated decoherence affects quantum systems in the vicinity of all massive bodies, except for black holes. The latter are vacuum solutions to Einstein's equation, and they do not contain any quantum matter to generate decoherence. The model constructed here applies to any gravitating system for which the Newtonian description of gravity is a good approximation. For example, it applies to quantum particles in the vicinity of compact stars, such as white dwarfs.

The structure of this paper is the following. In Sec. II, we develop the dynamics of our model, and we show that it is mathematically equivalent to QBM. In Sec. III, we present the master equation and derive a general formula for the decoherence rate. In Sec. IV, we make some simplifying approximations to perform an explicit calculation of the decoherence rate. In Sec. V, we summarize and discuss our results.

II. SETUP

We consider a particle of mass m under a potential $V(\mathbf{x})$ localized at \mathbf{x} , and interacting gravitationally with a spherically symmetric mass distribution of total mass M (see Fig. 1). We assume that the massive body is composed of a finite collection of uncoupled harmonic oscillators of masses m_i and frequencies ω_i , each located at $\mathbf{r}_i + \delta\mathbf{r}_i$, where the term $\delta\mathbf{r}_i$ describes small displacements from each oscillator's equilibrium position. The Hamiltonian of the combined system reads

$$H_{\text{tot}} = H_m + H_M + H_{\text{int}}, \quad (1)$$

where

$$H_m = \frac{\mathbf{p}^2}{2m} + V(\mathbf{x}), \quad (2)$$

$$H_M = \sum_i \left(\frac{\mathbf{p}_i^2}{2m_i} + \frac{1}{2} m_i \omega_i^2 \delta\mathbf{r}_i^2 \right), \quad (3)$$

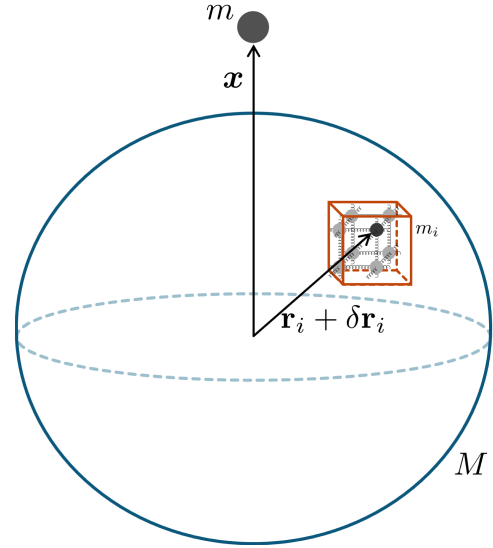


FIG. 1. A particle of mass m lies within the gravitational field of a massive body of a total mass M . The large body can be described as a collection of cubic cells, with each shell localized around a point \mathbf{r}_i and representing a crystal lattice of oscillators.

are, respectively, the free Hamiltonians of the particle and the harmonic oscillators and H_{int} is the interaction term between the small particle of mass m and each of the harmonic oscillators that compose the heavy body. The interaction Hamiltonian is described by the Newtonian gravitational potential

$$H_{\text{int}} = - \sum_i \frac{Gmm_i}{|\mathbf{x} - (\mathbf{r}_i + \delta\mathbf{r}_i)|}, \quad (4)$$

where G is the gravitational constant. Expanding the interaction term for small variations of $\delta\mathbf{r}_i$ around the separation distance $|\mathbf{x} - \mathbf{r}_i|$, we obtain

$$H_{\text{int}} = - \sum_i Gmm_i \left(\frac{1}{|\mathbf{x} - \mathbf{r}_i|} + \frac{(\mathbf{x} - \mathbf{r}_i) \cdot \delta\mathbf{r}_i}{|\mathbf{x} - \mathbf{r}_i|^3} + \dots \right). \quad (5)$$

The first term is of the form $m\phi(\mathbf{x})$, where ϕ is the gravitational potential generated by the massive body. It can therefore be absorbed in H_m . We assume a potential $V(\mathbf{x})$ that compensates for the gravitational acceleration, so that the small particle only performs small oscillations of frequency Ω around an equilibrium point \mathbf{x}_0 . Modulo a constant, the Hamiltonian H_m becomes $H_m = \frac{\mathbf{p}^2}{2m} + \frac{1}{2} m \Omega^2 \delta\mathbf{x}^2$, where $\delta\mathbf{x} = \mathbf{x} - \mathbf{x}_0$.

Substituting $\mathbf{x} = \mathbf{x}_0 + \delta\mathbf{x}$ in the second term of Eq. (5), we obtain to leading order in $\delta\mathbf{x}$,

$$H'_{\text{int}} = -\sum_i \frac{Gmm_i}{|\mathbf{x}_0 - \mathbf{r}_i|^3} \left((\mathbf{x}_0 - \mathbf{r}_i) \cdot \delta\mathbf{r}_i + \delta\mathbf{x} \cdot \delta\mathbf{r}_i - 3 \frac{(\mathbf{x}_0 - \mathbf{r}_i) \cdot \delta\mathbf{r}_i (\mathbf{x}_0 - \mathbf{r}_i) \cdot \delta\mathbf{x}}{|\mathbf{x}_0 - \mathbf{r}_i|^2} \right). \quad (6)$$

The first term can be absorbed by a shift in the equilibrium positions of the oscillators. The second and third terms give a genuine coupling between $\delta\mathbf{x}$ and $\delta\mathbf{r}_i$. Assuming that the small particle moves along the direction with unit vector \mathbf{n} , we can write $\delta\mathbf{x} = \delta x \mathbf{n}$. Then, the coupling term takes the form

$$H''_{\text{int}} = \delta x \sum_i \mathbf{c}_i \cdot \delta\mathbf{r}_i, \quad (7)$$

where

$$\mathbf{c}_i \equiv \frac{Gmm_i}{|\mathbf{x}_0 - \mathbf{r}_i|^3} \left[\frac{3\mathbf{n} \cdot (\mathbf{x}_0 - \mathbf{r}_i)(\mathbf{x}_0 - \mathbf{r}_i)}{|\mathbf{x}_0 - \mathbf{r}_i|^2} - \mathbf{n} \right] \quad (8)$$

defines coupling constants. In this way, the Hamiltonian of the total system can be placed into the form

$$H_{\text{tot}} = H_{\text{m}} + H_{\text{M}} + \delta x \sum_i \mathbf{c}_i \cdot \delta\mathbf{r}_i, \quad (9)$$

which is equivalent to the QBM Hamiltonian that describes the dynamics of a Brownian particle—modeled as a harmonic oscillator—interacting with a thermal environment comprising a large number of independent harmonic oscillators.

III. MASTER EQUATION

The dynamics of the reduced density matrix $\rho(\tau)$ of the small particle, obtained by tracing out the degrees of freedom of the environment, can be described by the Hu-Paz-Zhang (HPZ) master equation [15]:

$$\frac{d\rho(\tau)}{d\tau} = -\frac{i}{\hbar} [\tilde{H}_m(\tau), \rho(\tau)] - i\gamma(\tau) [\delta x, \{p, \rho(\tau)\}] - D(\tau) [\delta x, [\delta x, \rho(\tau)]] - f(\tau) [\delta x, [p, \rho(\tau)]]. \quad (10)$$

In the above equation, the term $\tilde{H}_m(\tau)$ represents the particle Hamiltonian with a time-dependent frequency shift $\Omega^2 + \delta\Omega^2(\tau)$. The coefficient $\gamma(\tau)$ describes dissipation, while $D(\tau)$ and $f(\tau)$ are diffusion terms. The explicit expressions of the time-dependent coefficients $\delta\Omega^2(\tau)$, $\gamma(\tau)$, $D(\tau)$, and $f(\tau)$, which are rather cumbersome to be reported here, can be found in Refs. [15,16].

We note that the HPZ master equation is exact. It is valid for arbitrary temperatures of the environment and open system-environment coupling strengths. It is derived only under the assumption of a factorized initial state $\rho_{\text{tot}} = \rho \otimes \rho_{\text{env}}$ for the total system, where the

environment is in a thermal equilibrium state ρ_{env} at temperature T .

Straightforward expressions for the coefficients in the master equation (10) can be obtained by employing different approximation schemes during its derivation. Examples include the assumption of a weak coupling between the open system and the environment, and the application of the Markov approximation, which neglects memory effects in the time evolution of the open system [17,18]. Under the latter two approximations, the diffusion coefficient D is given by

$$D = \frac{1}{\hbar^2} \int_0^\infty ds \nu(s) \cos(\Omega s), \quad (11)$$

where

$$\nu(\tau) := \sum_i \frac{\hbar |\mathbf{c}_i|^2}{2m_i \omega_i} \coth\left(\frac{\hbar \omega_i}{2k_B T_i}\right) \cos(\omega_i \tau) \quad (12)$$

is the so-called *noise kernel*. We have assumed a different temperature T_i for each oscillator. Indeed, if the large body is taken to be Earth, the temperature depends on the distance from the center.

In the position representation, $\rho(x, x', \tau) \equiv \langle x | \rho(\tau) | x' \rangle$, the term of the master equation with the coefficient D can be expressed as

$$-D[\hat{x}, [\hat{x}, \rho(\tau)]] \longrightarrow -D(x - x')^2 \rho(x, x', \tau), \quad (13)$$

which indicates that the off-diagonal components ($x \neq x'$) of the reduced density matrix decohere at a rate $D(x - x')^2$ [19–22]. Thus, the quantity D , which has dimensions $[\text{time}]^{-1} \times [\text{length}]^{-2}$, allows the definition of a *decoherence time*

$$\tau_{\text{dec}} = \frac{1}{D \Delta x^2}, \quad (14)$$

as the characteristic timescale on which spatial coherence over a distance $\Delta x = x - x'$ becomes suppressed.

IV. THE NOISE KERNEL

We treat the heavy body as a solid composed by three-dimensional cubic cells, with each shell centered at \mathbf{r}_i and representing a crystal lattice of oscillators (see Fig. 1) at constant temperature $T(\mathbf{r}_i)$. Each of these cubic cells comprises a total number of normal modes

$$N = \frac{1}{3} \int_0^\infty g_r(\omega) d\omega, \quad (15)$$

where $g_r(\omega)$ is the density of modes whose frequencies lie in the infinitesimal range between ω to $d\omega$. The index \mathbf{r} denotes spatial dependency into the characterization of the

density of modes of each particular cubic cell. Hence, the sum over i in the noise kernel (12) becomes a sum over all distances \mathbf{r}_i and over all modes in a given cell, labeled by λ . That is, $\sum_i \rightarrow \sum_{\mathbf{r}_i} \sum_{\lambda}$.

Employing the continuum limit for the spectrum of the environmental frequencies ω_λ in the noise kernel (12), we obtain

$$\nu(\tau) = \sum_{\mathbf{r}} \int d\omega \frac{\hbar G^2 m^2 m_r}{2\omega} \left(|\mathbf{C}_r|^2 g_r(\omega) \times \coth\left(\frac{\hbar\omega}{2k_B T(\mathbf{r})}\right) \cos(\omega\tau) \right), \quad (16)$$

where for later convenience we have written the coupling constants (8) as $\mathbf{c}_i = Gmm_i\mathbf{C}_i$.

We will next consider a frequency density of the form $g_r(\omega) = \alpha_r \omega^{k(r)}$, up to a cutoff frequency $\omega_c(\mathbf{r})$. The spectral densities of many systems of interest (e.g., phonon baths, EM field baths) indeed follow a power law. More generally, decoherence is primarily due to the infrared frequency modes of the environment, so it is usually sufficient to consider the dominant behavior of the spectral density as $\omega \rightarrow 0$; a power law is generic in this limit [13]. Here, $k > 1$ is a constant specific to each cubic cell in the crystal solid. The normalization constant α_r is directly computed through Eq. (15). It follows that

$$g_r(\omega) = \begin{cases} \frac{3N(k(r)+1)}{\omega_c(\mathbf{r})^{k(r)+1}} \omega^{k(r)}, & \text{for } \omega \leq \omega_c(\mathbf{r}) \\ 0, & \text{for } \omega > \omega_c(\mathbf{r}) \end{cases}. \quad (17)$$

Accordingly, the noise kernel is given by

$$\begin{aligned} \nu(\tau) &= \frac{3}{2} \hbar G^2 m^2 \sum_{\mathbf{r}} \frac{k(\mathbf{r})+1}{\omega_c(\mathbf{r})^{k(\mathbf{r})+1}} \int d\omega N m_r |\mathbf{C}_r|^2 \omega^{k(\mathbf{r})-1} \\ &\quad \times \coth\left(\frac{\hbar\omega}{2k_B T(\mathbf{r})}\right) \cos(\omega\tau) \\ &= \frac{3}{2} \hbar G^2 m^2 \int d^3\mathbf{r} \frac{k(\mathbf{r})+1}{\omega_c(\mathbf{r})^{k(\mathbf{r})+1}} \rho(\mathbf{r}) |\mathbf{C}_r|^2 \int_0^{\omega_c} d\omega \omega^{k(\mathbf{r})-1} \\ &\quad \times \coth\left(\frac{\hbar\omega}{2k_B T(\mathbf{r})}\right) \cos(\omega\tau), \end{aligned} \quad (18)$$

where $Nm_r = m(\mathbf{r})$ is the mass of each cell. In the second equality the continuum-mass limit is taken and $\rho(\mathbf{r})$ is the density of the spherical symmetric mass distribution. Hence, to evaluate the noise kernel, we need explicit forms for the functions $T(\mathbf{r})$, $\rho(\mathbf{r})$, $\omega_c(\mathbf{r})$, $k(\mathbf{r})$ in the interior of the compact body.

Here, we will consider the case where these functions are constant. For Earth, this does not affect the order of magnitude of D and, hence, the decoherence time. The density from Earth's center to the surface changes at most

by a factor of 5; the temperature is significantly higher at the center, but the contribution of internal layers to D is strongly suppressed.

For constant functions $T(\mathbf{r})$, $\rho(\mathbf{r})$, $\omega_c(\mathbf{r})$, $k(\mathbf{r})$, and for a spherical body of radius R , we only need to compute the integral

$$I(x_0, R) = \int_{r \leq R} d^3\mathbf{r} |\mathbf{C}_r|^2. \quad (19)$$

For x_0 along the z-axis, particle motion along $\mathbf{n} = (0, 1, 0)$, and $\mathbf{r} = (r \sin \theta \cos \phi, r \sin \theta \sin \phi, r \cos \theta)$, we obtain

$$\begin{aligned} I(x_0, R) &= \int_{r \leq R} \frac{d^3\mathbf{r}}{|\mathbf{x}_0 - \mathbf{r}|^6} \left(1 + \frac{9r^4 \sin^4 \theta \sin^2 \phi}{|\mathbf{x}_0 - \mathbf{r}|^4} \right. \\ &\quad \left. - \frac{6r^2 \sin^2 \theta \sin^2 \phi}{|\mathbf{x}_0 - \mathbf{r}|^2} + \frac{9r^2 \sin^2 \theta \sin^2 \phi}{|\mathbf{x}_0 - \mathbf{r}|^4} \right) \\ &\quad \times (x_0 - r \cos \theta)^2. \end{aligned} \quad (20)$$

We evaluate this integral in Appendix A. We note that, for a body at distance $d \ll R$ from the surface of Earth, $x_0 = R + d$ and

$$I(x_0, R) \simeq \frac{3\pi}{16d^3}. \quad (21)$$

The noise kernel reads

$$\begin{aligned} \nu(\tau) &= \frac{9(k+1)\hbar G^2 m^2 M}{8\pi\omega_c^{k+1} R^3} I(x_0, R) \\ &\quad \times \int_0^{\omega_c} d\omega \omega^{k-1} \coth\left(\frac{\hbar\omega}{2k_B T}\right) \cos(\omega\tau), \end{aligned} \quad (22)$$

where M is the total mass of the massive body.

The noise kernel diverges for $d \rightarrow 0$. This means that the decoherence rate is very sensitive on the outer layers of the massive body. When using the noise kernel to model a specific experiment, we must split the mass distribution into a spherical part that corresponds to the bulk of the Earth and a part that models the immediate surroundings of the particle in the experiment. In this sense, d is best understood as an effective distance of the particle from the surrounding masses. For particles near the surface of the Earth, the effective d is of the order of the length scales that characterize the laboratory.

A. The decoherence time

We next focus our attention to the case of a function of the density of states that is quadratic to frequency; i.e, we choose $k = 2$ in Eq. (17). This is equivalent to the Debye model for the density of states of solids [23,24]; see also Appendix B. The cutoff frequency ω_c coincides with the Debye frequency ω_D of the solid, and the environmental

TABLE I. The decoherence time in different regions of the parameter space. The distance of 400 km corresponds to the average altitude at which the International Space Station maintains its orbit. The mass m of 10^{-27} kg corresponds to the atomic mass.

Massive body	Distance from surface	Mass m	Decoherence time
Earth	10 cm	100 kg	$\tau_{\text{dec}} \sim 1$ s
Earth	10 cm	10^5 kg	$\tau_{\text{dec}} \sim 10^{-6}$ s
Earth	400 km	100 kg	$\tau_{\text{dec}} \sim 10^{20}$ s
Earth	10 cm	10^{-27} kg	$\tau_{\text{dec}} \sim 10^{58}$ s
White dwarf	10 cm	100 kg	$\tau_{\text{dec}} \sim 10^{-6}$ s
White dwarf	10 m	100 kg	$\tau_{\text{dec}} \sim 1$ s

degrees of freedom are characterized by an Ohmic spectral density (in an Ohmic environment, the damping force experienced by the Brownian particle is linear to its velocity [15,25]).

The decoherence rate reads

$$D = \frac{27G^2m^2M}{8\pi\hbar\pi\omega_D^3R^3} I(x_0, R) \times \int_0^\infty d\tau \cos(\Omega\tau) \int_0^{\omega_D} d\omega \omega \coth\left(\frac{\hbar\omega}{2k_B T}\right) \cos(\omega\tau). \quad (23)$$

At the limit $\omega_D \rightarrow \infty$, we obtain

$$D = \frac{27G^2m^2M}{8\pi\hbar\omega_D^3R^3} I(x_0, R) \left(\frac{\pi\Omega}{2} \coth\left(\frac{\hbar\Omega}{2k_B T}\right)\right). \quad (24)$$

In the physically relevant regime, $k_B T \gg \hbar\Omega$, we approximate $\coth(\hbar\Omega/2k_B T) \approx 2k_B T/\hbar\Omega$. Then,

$$D = \frac{27G^2m^2M}{8R^3} \frac{k_B T}{\hbar^2\omega_D^3} I(x_0, R). \quad (25)$$

In Table I, we present the decoherence time for a superposition of two localized states with mean separation of $\Delta\mathbf{x} = 1$ m in various cases, calculated by means of Eq. (25). For a human at a distance 10 cm from the surface of Earth ($R_\oplus = 6.4 \times 10^6$ m, $M_\oplus = 5.9 \times 10^{24}$ kg), we find a decoherence time of order of 1 s. Earth is described as a Debye solid with a Debye frequency of order $\omega_D \sim 10^{13}$ Hz. We also give the decoherence rate for a white dwarf. During the crystallization phase of its ion lattice, a white dwarf ($M \sim 1M_\odot = 2 \times 10^{30}$ kg, $R \sim R_\oplus$) can also be described as a Debye solid [26]. In this case, the Debye frequency is of the order $\omega_D \sim 10^{18}$ Hz.

V. CONCLUSIONS

We showed that any quantum particle within the gravitational field of a massive body is subject not only to the

gravitational pull but also to nonunitary dynamics that originate from the intrinsic fluctuations of the matter in the massive body and are mediated by the gravitational field.

The nonunitary contribution is of second-order to the gravitational constant, and for this reason, it is implicitly assumed to be negligible and it is usually ignored. However, the strength of a nonunitary term depends on the strength of the fluctuations and may become quite strong, for example, at high temperatures or in an environment with a strong concentration of modes in the deep infrared. In any case, the gravitational field will mediate a decoherence process, from which no experiment in the gravitational field of the massive body can be shielded.

Our analysis demonstrated that this effect does not affect the regime of interest for near-future experiments with macroscopic quantum systems. It is expected to appear for superpositions at the human scale.

We note that our treatment is completely general: it can be applied to describe quantum phenomena within the gravitational field of any compact body that is compatible with the Newtonian description of gravity. It allows for the incorporation of fine details, such as density and temperature gradients, or the gravitational influence of the immediate surroundings of the quantum system.

Our analysis in this paper is restricted to positional decoherence. However, this is not the only way that the environment can destroy quantum coherences. Since the environmental degrees of freedom are entangled with those of the quantum particles, it will also affect the generation of entanglement in multipartite systems due to the gravitational force. It is therefore important to analyze the degree to which proposed experiments for gravity-induced entanglement are affected by the presence of gravity-mediated decoherence.

Finally, we note that our analysis straightforwardly applies to the decoherence induced by a large charged body on microscopic charged particles. In this case, the Coulomb-mediated decoherence effects are much stronger and, hence, experimentally accessible.

APPENDIX A: EXPLICIT EXPRESSIONS FOR THE DIFFUSION CONSTANT

We evaluate the quantity $I(x_0, R)$ defined by Eq. (19). Assuming that $\mathbf{x} = (0, 0, x_0)$, $\mathbf{n} = (0, 1, 0)$, and $\mathbf{r} = (r \sin \theta \cos \phi, r \sin \theta \sin \phi, r \cos \theta)$,

$$I(x_0, R) = \int_{r \leq R} \frac{d^3\mathbf{r}}{|\mathbf{x}_0 - \mathbf{r}|^6} \left(1 + \frac{9r^4 \sin^4 \theta \sin^2 \phi}{|\mathbf{x}_0 - \mathbf{r}|^4} - \frac{6r^2 \sin^2 \theta \sin^2 \phi}{|\mathbf{x}_0 - \mathbf{r}|^2} + \frac{9r^2 \sin^2 \theta \sin^2 \phi}{|\mathbf{x}_0 - \mathbf{r}|^4} \right) \times (x_0 - r \cos \theta)^2. \quad (A1)$$

For a sphere of radius R , we define $\mathbf{z} = \mathbf{x}_0/R$, $\mathbf{x} = \mathbf{r}/R$ to obtain

$$I(x_0, R) = \frac{1}{R^3} F(z), \quad (\text{A2})$$

where

$$F(z) = 2\pi \int_0^1 dx \int_{-1}^1 d\xi \frac{x^2}{|\mathbf{z} - \mathbf{x}|^6} \left(1 + \frac{9x^4(1 - \xi^2)}{2|\mathbf{z} - \mathbf{x}|^4} - \frac{3x^2(1 - \xi^2)}{|\mathbf{z} - \mathbf{x}|^2} + \frac{9z^2x^2(1 - \xi^2)}{2|\mathbf{z} - \mathbf{x}|^4} - \frac{9zx^3(1 - \xi^2)\xi}{|\mathbf{z} - \mathbf{x}|^4} \right), \quad (\text{A3})$$

and we wrote $\xi = \mathbf{z} \cdot \mathbf{x} / (zx)$.

We carry out the integration over ξ to obtain

$$F(z) = \frac{4\pi}{3} \frac{1}{(z^2 - 1)^3} + 4\pi \int_0^1 dx \frac{x^4}{(x^2 - z^2)^4} = \frac{4\pi}{3} \frac{1}{(z^2 - 1)^3} + \pi \frac{3z + 8z^3 - 3z^5}{12z^3(z^2 - 1)^3} + \frac{\pi}{4z^3} \operatorname{arccoth}(z). \quad (\text{A4})$$

For $z = 1 + x$, with $x \ll 1$, we obtain

$$F(x) \simeq \frac{3\pi}{16x^3}. \quad (\text{A5})$$

APPENDIX B: DEBYE SOLID

A solid can be viewed as an ordered array of atoms, where each atom is fixed to a lattice site and can oscillate about its equilibrium position. The Debye model [23,24] treats these atomic vibrations as sound waves that propagate through the crystal lattice at the speed of sound v_s . The vibrational energy is quantized, and the quanta of

vibrational energy are known as *phonons*, analogous to photons in electromagnetic waves. The frequencies of these phonons are linearly related to their wave vectors \mathbf{k} through the dispersion relation $\omega(\mathbf{k}) = v_s |\mathbf{k}|$.

The Debye model allows for a continuum spectrum of oscillation frequencies, resulting in a total number of normal modes of vibration equal to $3N$, expressed as

$$3N = \int_0^{\omega_D} g(\omega) d\omega, \quad (\text{B1})$$

where N is the number of atoms in the crystal, $g(\omega)$ is the density of normal modes meaning that $g(\omega)d\omega$ represents the number of normal modes with frequencies in the infinitesimal range between ω and $\omega + d\omega$, and ω_D is a maximum allowed phonon frequency, known as the *Debye frequency*. The Debye frequency acts as a cutoff for phonon frequencies and is determined by the minimum wavelength allowed for sound wave propagation in the lattice, constrained by the finite interatomic distance.

The Debye model approximates the density of modes as

$$g(\omega) = \begin{cases} \frac{9N}{\omega_D^3} \omega^2, & \text{for } \omega \leq \omega_D \\ 0, & \text{for } \omega > \omega_D \end{cases}. \quad (\text{B2})$$

The Debye frequency is defined by

$$\omega_D^3 = 18\pi^2 n \left(\frac{1}{v_L^3} + \frac{2}{v_T^3} \right)^{-1}, \quad (\text{B3})$$

where n is the atomic density, v_L is the velocity of the longitudinal sound mode, and v_T is the velocity of the transverse sound modes.

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