

## Searching for axion dark matter with the MeerKAT radio telescope

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 (Received 19 December 2021; revised 3 April 2022; accepted 19 September 2022; published 10 October 2022)

Axions provide a natural and well-motivated dark matter candidate, with the capability to convert directly to photons in the presence of an electromagnetic field. A particularly compelling observational target is the conversion of dark matter axions into photons in the magnetospheres of highly magnetised neutron stars, which is expected to produce a narrow spectral peak centered at the frequency of the axion mass. We point the MeerKAT radio telescope towards the isolated neutron star J0806.4–4123 for 10-hours of observation and obtain the radio spectra in the frequency range 769–1051 MHz. By modeling the conversion process of infalling axion dark matter (DM), we then compare these spectra to theoretical expectations for a given choice of axion parameters. Whilst finding no signal above  $5\sigma$  in the data, we provide a unique constraint on the Primakoff coupling of axion DM,  $g_{\text{a}\gamma\gamma} \lesssim 9.3 \times 10^{-12} \text{ GeV}^{-1}$  at the 95% confidence level, in the mass range 3.18–4.35  $\mu\text{eV}$ . This result serves the strongest constraint in the axion mass range 4.20–4.35  $\mu\text{eV}$ .

DOI: [10.1103/PhysRevD.106.083006](https://doi.org/10.1103/PhysRevD.106.083006)

### I. INTRODUCTION

As a minimal extension of the Standard Model, and in particular, the Peccei-Quinn solution of the strong  $CP$  problem [1–4], axions, and axionlike particles occupy a rare point of convergence in theoretical physics, in that they are

also a generic prediction of the exotic physics of string and M theory [5,6]. Despite the profound differences between these contexts, their resultant properties are also largely universal, creating an easily characterizable theoretical target.

As typically light, long-lived pseudoscalar particles, axions are also natural candidates for the mysterious dark matter (DM) comprising much of the mass of our observable Universe [7,8], becoming a topic of intense ongoing investigation [9,10]. In recent years, a particularly compelling observational mechanism has emerged, Primakoff conversion of these DM axions into photons

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in the magnetospheres of highly magnetized neutron stars (NS) [11]. Following Ref. [11], this idea was revisited more thoroughly in Refs. [12,13], leading to a flurry of theoretical activities [14–19].

Based on the observed DM density, the so-called QCD axions which solve the strong  $CP$  problem are favored to have masses in the 10–100  $\mu\text{eV}$  range [20]. This corresponds to a characteristic signature of NS conversion in the MHz–GHz frequency range, exactly covered by mid-frequency radio telescopes (e.g., Square Kilometre Array [21], and FAST [22]). Radio antennae can therefore provide a unique and complementary method to laboratory axion DM haloscope experiments (e.g., ADMX [23]). Several observational studies recently have been performed using the Karl G. Jansky Very Large Array [24–26], and the Green Bank and Effelsberg radio telescopes [27,28] to search for the signature of this conversion mechanism.

In the following, we will also search for this phenomenon by observing the isolated NS J0806.4–4123 with the MeerKAT radio telescope [29]. MeerKAT has a good sensitivity along with a wide UHF band coverage to search for this type of axion signature in the low mass window where there is a gap between ADMX [30] and RBF [31] laboratory experiments.

## II. AXION CONVERSION IN NS MAGNETOSPHERES

Our starting point is the Lagrangian that represents coupling between the electromagnetic field and the axion field  $a$ ,  $\mathcal{L}_{a\gamma} = -(1/4)g_{a\gamma\gamma}aF_{\mu\nu}\tilde{F}^{\mu\nu} = g_{a\gamma\gamma}a\vec{E}\cdot\vec{B}$ , where  $g_{a\gamma\gamma}$  is the coupling constant and  $\vec{E}/\vec{B}$  are the electric/magnetic fields [10,24,25] (we work throughout in natural units where  $\hbar = c = 1$ , such that mass and frequency have the same dimension).

To estimate the resulting radio flux we follow the analysis of Ref. [12], in common with the other observational studies presented in Refs. [24,25,27]. Therein, using a variant of the Goldreich-Julian model [32] assuming the presence of electrons and positrons only, we have a magnetosphere plasma frequency  $\omega_p \simeq (4\pi\alpha n_e/m_e)^{1/2}$ , where  $\alpha$  is the fine structure constant and  $n_e/m_e$  are the electron number density and mass, respectively. Since the axion/photon conversion probability is maximized on resonance, when  $m_a \simeq \omega_p$ , we focus on the critical radius where this condition is satisfied,

$$r_c = 224 \text{ km} |3 \cos \theta \hat{\mathbf{m}} \cdot \hat{\mathbf{r}} - \cos \theta_m|^{1/3} \times \left( \frac{r_{\text{NS}}}{10 \text{ km}} \right) \left[ \frac{B_0}{10^{14} \text{ G}} \frac{1 \text{ sec}}{P} \left( \frac{1 \text{ GHz}}{m_a} \right)^2 \right]^{1/3}. \quad (1)$$

Here,  $\theta$  and  $\theta_m$  are the angles between the NS axis of rotation and respectively, our observational line of sight and the magnetic dipole axis,  $\hat{\mathbf{m}} \cdot \hat{\mathbf{r}} = \cos \theta_m \cos \theta + \sin \theta_m \sin \theta \cos(\omega t)$ .  $P$  is the NS period and  $\omega = (2\pi/P)$ ,

$r_{\text{NS}}$  is the NS radius, and  $B_0$  is the magnetic field strength at the poles.

Since the NS is relatively close to the Earth ( $d \simeq 250 \text{ pc}$  [33]), we assume a standard Maxwell-Boltzmann DM velocity distribution with local density  $\rho_{\text{DM}}^\infty = 0.45 \text{ GeV cm}^{-3}$  [34,35] asymptotically far from the NS surface, which leads to a radio signal with intrinsic linewidth  $\Delta f/f \simeq v_0^2$  [12], where  $v_0 \sim 10^{-3}$  is the DM velocity dispersion.

There are however additional effects which can further broaden our expected signal. It has, for example, been argued in Refs. [15,18] that misalignment of the NS rotation axis and the critical surface of axion/photon conversion can result in significant broadening above the intrinsic linewidth, although the conclusions of Ref. [27] differ on this point. More recently, this issue has been revisited in more comprehensive detail in [17], where it was found that various effects can broaden the signal linewidth by more than an order of magnitude, albeit primarily in the large  $|\vec{\omega}|$  regime. However, in the  $|\vec{\omega}| \sim 0.1 \text{ s}^{-1}$  regime, we currently operate for NS J0806.4–4123, the additional broadening was found to be minimal, for the fiducial models considered therein.

Therefore, we adopt  $\Delta f/f \simeq 5 \times 10^{-6}$ , in line with Ref. [27] and the original analysis in Ref. [12]. For the frequency range of our interest the resulting linewidth is then comparable with the 16 kHz MeerKAT channel width, such that the expected signal can be completely contained in three adjacent channels. The average flux density in channel  $i$  is then given by

$$\bar{S}_{\nu_i} = \frac{F}{\Delta\nu} = 3.8 \times 10^{-6} \text{ Jy} \left( \frac{100 \text{ pc}}{d} \right)^2 \left( \frac{16 \text{ kHz}}{\Delta\nu} \right) \times \left( \frac{d\mathcal{P}/d\Omega}{5.7 \times 10^9 \text{ W}} \right) \int_{\nu_{i,\text{min}}}^{\nu_{i,\text{max}}} \frac{d\nu}{\sqrt{2\pi}\sigma_0} e^{-\frac{(\nu-m_a)^2}{2\sigma_0^2}}, \quad (2)$$

where we assume a Gaussian shape of the spectrum with intrinsic width  $\sigma_0 = 5 \times 10^{-6} m_a$ .  $\nu_{i,\text{min}}$  and  $\nu_{i,\text{max}}$  are the start and end frequencies of the channel, and  $\Delta\nu$  is the channel width. The radiated power is given by

$$\frac{d\mathcal{P}}{d\Omega} \simeq 5.7 \times 10^9 \text{ W} \left( \frac{g_{a\gamma\gamma}}{10^{-12} \text{ GeV}^{-1}} \right)^2 \left( \frac{r_{\text{NS}}}{10 \text{ km}} \right)^{5/2} \left( \frac{m_a}{\text{GHz}} \right)^{4/3} \times \left( \frac{B_0}{10^{14} \text{ G}} \right)^{5/6} \left( \frac{P}{\text{sec}} \right)^{7/6} \left( \frac{\rho_{\text{DM}}^\infty}{0.45 \text{ GeV cm}^{-3}} \right) \left( \frac{M_{\text{NS}}}{M_\odot} \right)^{1/2} \times \left( \frac{200 \text{ km s}^{-1}}{v_0} \right) \frac{3(\hat{\mathbf{m}} \cdot \hat{\mathbf{r}})^2 + 1}{|3 \cos \theta \hat{\mathbf{m}} \cdot \hat{\mathbf{r}} - \cos \theta_m|^{7/6}}, \quad (3)$$

where, following Ref. [17] we have reintroduced the missing factor of  $1/v(r_c)$  identified in Ref. [18] as being absent from the original expression in Ref. [12]. Time dependence enters in “ $\hat{\mathbf{m}} \cdot \hat{\mathbf{r}}$ ” terms, but lacking sufficient time resolution we average over the NS period to derive our resulting constraint.

It is important to emphasize that this theoretical approach relies on various simplifying assumptions and so has some limitations. For example, following Refs. [16–18], incorporating ray tracing of the emitted photons can diminish the estimated total emitted power relative to the analysis of [12], corresponding to an  $\mathcal{O}(1)$  increase in the minimum detectable  $g_{\text{ax}\gamma}$ . In this sense, our estimated limit can be viewed as less conservative than those arrived at from a full ray-tracing procedure.

A more realistic modeling of the NS magnetosphere and axion/photon conversion in this environment is also needed. However, from this perspective, our estimate can be regarded as conservative insofar as the analysis of Ref. [19] suggested that more thorough modeling can provide an enhanced signal strength. Most recently, it has also been suggested in Ref. [36] that the underlying resonance condition could be extended to arbitrarily light axions, further modifying the underlying theoretical picture.

General relativistic corrections are also absent, although as noted in Ref. [16] these typically amount to a percent-level correction overall. Reference [18] furthermore suggested that these effects can in fact be helpful insofar as they can counterbalance refractive effects of the magnetosphere, although for our analysis this point is not salient.

As an overall perspective, it is worth noting that as a relatively novel approach to axion discovery there is at present no analysis in the literature accounting for all these effects. As emphasized in Ref. [17] various aspects of the radio flux estimation, especially the true broadening of the expected signal, are still ongoing topics of debate. As such, any limits derived therefrom are to be understood within this context.

### III. DATA ACQUISITION

MeerKAT was pointed towards NS J0806.4–4123 (RA =  $08^{\text{h}}06^{\text{m}}23.3471^{\text{s}}$ , DEC =  $-41^{\circ}22'30.179''$ , J2000), in six observing epochs of  $\sim 100$  minutes, in 2021 (11 March/64 antennas, 14 March/63 antennas, 24 March/61 antennas, 2 April/62 antennas, 4 April/63 antennas, 11 May/60 antennas functional). The complete UHF band (544–1087.983 MHz including the tapered edges of the band) was covered in shared-risk commissioning observations. The backend was configured in 32k mode (32768 channels), resulting in a frequency resolution of 16.602 kHz. The integration time was 8 s and all (linear) polarization products were observed and recorded, although the cross-polarization data were not used for the generation of final data products.

To avoid systematic effects associated with the  $uv$ -coverage we observed the source over three different hour angle ranges. In each period, a bandpass calibrator (either J0408–6545 or J1939–6342, depending on the LST range) was observed for 8 min at the beginning and end of the observations, and a gain calibrator (J0828–3731) was

observed for 2 min before each of the 3 target observations, each lasting 26 min.

We made use of the data products provided by the MeerKAT science-data-processor (SDP) pipeline, which can be accessed through the MeerKAT archive interface. The SDP pipeline provides raw images (channel maps) from an automated calibration routine, which performs flagging of the data for radio frequency interference (RFI), cross calibration, including bandpass and gain calibrations. It then performs a continuum calibration based on the OBIT data reduction package [37], with two rounds of self-calibration.

Wide-field effects are dealt with using a faceting approach and wideband effects are mitigated by using subbands. The continuum model is then subtracted from the visibilities. The continuum-subtracted data are used to produce CLEANed images with Briggs’s robust weighting 0 [38] at full frequency resolution (16.602 kHz) employing a dedicated imaging software created for the usage on GPU units. We combined the single channel maps from the archive into a data cube and performed a further continuum-subtraction employing a median filter. The original images cover the primary beam size ranging from 104 arcmin to 208 arcmin (HPWB), but we made use of an inner fraction of the images only ( $14.9 \text{ arcmin} \times 14.9 \text{ arcmin}$ ) only, sufficient to measure the flux density in the central pixel. The channel maps were convolved with Gaussians such that the resulting synthesized beam size was identical at all frequencies and in all observations,  $20.9 \text{ arcsec} \times 13.4 \text{ arcsec}$  (half-power-beam-width, HPBW) and regridded to the identical pixel size ( $3.5 \text{ arcsec}$ , fully sampling the synthesized beam). This gave us six data cubes, which were inspected individually to identify potential problems or artifacts in the data, then averaged into a final data cube.

### IV. ANALYSIS AND RESULTS

Channels with frequencies below 769 MHz and in the interval of [935, 961] MHz were discarded because the former showed systematic features possibly due to correlator issues, and the latter due to RFI (Fig. 1). Our final spectrum covers the ranges of 769–935 MHz and 961–1051 MHz, thus occupying 256 MHz total bandwidth (16,000 channels, as shown in Fig. 1). Voxel unit is mJy/beam, a measure of the intensity, also representing the flux density in units of mJy in the case of point sources.

We now analyze the cleaned cube to extract the flux density and search for potential axion signal. The central pixel coincides with the pulsar’s location and its value reflects the source’s flux density ( $d_{\nu_i}$ ). We adopt an “aperture photometry” method to characterize the background fluctuations [39,40]. We draw two circles that cover the center pixel with radii 2.25 arcmin and 6 arcmin, respectively. The outer circle is sufficiently large so that the ring area can represent the background fluctuations. We then compute the mean ( $\mu_{\nu_i}$ ) and root-mean-square



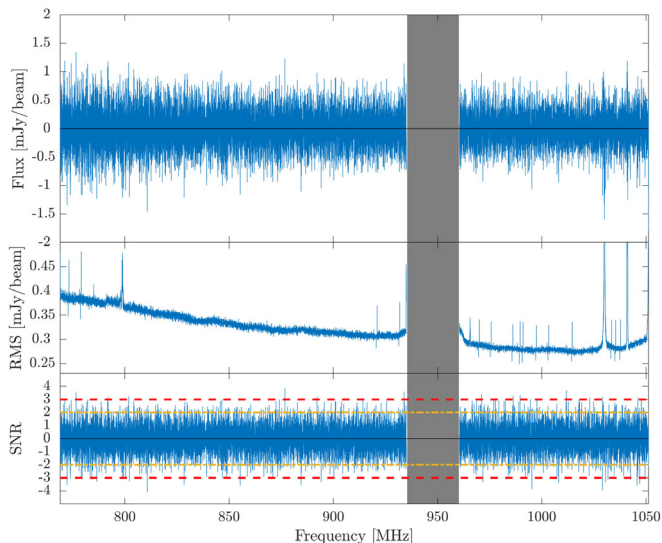


FIG. 1. The measured flux density, i.e., spectra of the NS (top), the root-mean-square (*rms*; middle), and signal-to-noise ratio (SNR; bottom) as a function of frequency for the 10-hour MeerKAT observation. The grey shaded region is the frequency range 935–961 MHz most affected by radio frequency interference (RFI). In the bottom panel, we plot the  $2\sigma$  CL (orange) and  $3\sigma$  CL (red) limits for visual comparison.

(*rms*;  $\sigma_{\nu_i}$ ) values of all pixels contained in the region enclosed by the two circles (roughly 28,500 pixels). We treat ( $\mu_{\nu_i}$ ,  $\sigma_{\nu_i}$ ) as the mean and *rms* of the background continuum, so any possible detection of axion signal in the pulsar location should have an excess compared with the background value [see also Eq. (4)]. We compare the measured and the expected *rms* for five channels using the SEFD provided in the MeerKAT description in the radiometer equation [41], to then multiply with a typical correction factor of 1.45. This “fudge factor” is typical for MeerKAT, accounting for Robust-0 weighting (as compared to Natural weighting, for which the radiometer equation is valid). The measured *rms* matches the theoretical prediction well, confirming the integrity of the data.

The mean and *rms* values for each channel are plotted in the top and middle panels of Fig. 1. One can see that the measured flux density varies for each frequency channel but the mean is about zero across all frequency bands. In addition, the absolute flux value decreases slightly from low to high frequency bands. This trend is exactly captured in the middle panel, in which we plot the *rms* value ( $\sigma_{\nu_i}$ ) for all channels. The *rms* decreases from 0.4 to 0.28 mJy/beam and therefore strengthens the constraint at higher frequencies. The trend can be understood as a consequence of a lower system temperature in the UHF band towards higher frequencies, although the amplitude of the decrease is slightly higher than expected.

In the bottom panel of Fig. 1, we plot the signal-to-noise ratio, defined as the ratio between flux density and the

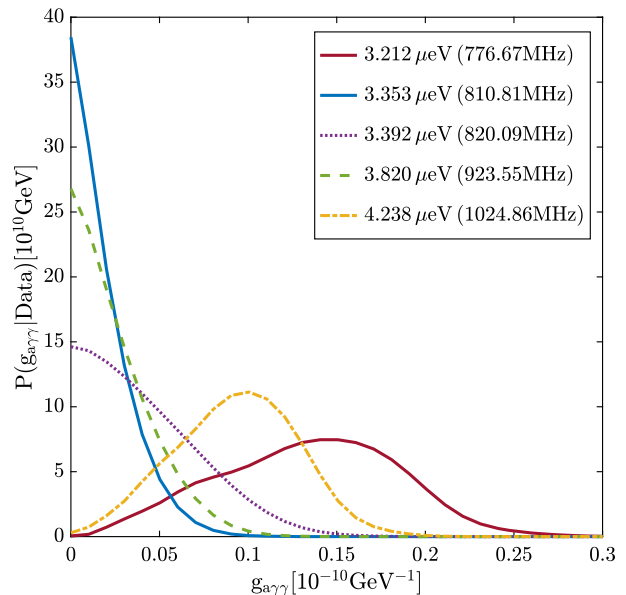


FIG. 2. Posterior distribution of the axion coupling constant  $g_{a\gamma\gamma}$  for five representative axion masses. The 95% CL upper limits from these five probability distributions are ( $g_{a\gamma\gamma} \times 10^{10} \text{ GeV}$ ) < 0.213(3.212  $\mu\text{eV}$ ), 0.054(3.353  $\mu\text{eV}$ ), 0.104(3.392  $\mu\text{eV}$ ), 0.069(3.820  $\mu\text{eV}$ ), and 0.149(4.238  $\mu\text{eV}$ ). Notice that although being significant apparently, the red line does not give more than  $4\sigma$  CL detection.

corresponding channel’s *rms*, and overlaid the  $2\sigma$  (orange) and  $3\sigma$  (red) confidence level (CL) limits for visual comparison. It is clear that the SNR of each channel varies, but is centered at zero. In addition, we checked that there are 40 channels with SNR exceeding  $\pm 3\sigma$ , and only one channel exceeding  $\pm 4\sigma$  SNR. Since there are 16,000 channels in total, the expected numbers of channels that can exceed  $3\sigma$  and  $4\sigma$  CL due purely to the noise fluctuation are  $16,000 \times 0.27\% = 43.2$ , and  $16,000 \times 0.0063\% = 1.01$ . These numbers are in perfect consistency with the measurements. No peak exceeding  $5\sigma$  has been found in our data. The statistics indicates that the rare peaks we have found are within the allowance of noise fluctuations. We therefore regard the measured flux density as consistent with noise in the radio continuum background, and derive upper limits on the Primakoff coupling constant between axions and photons from the data. The worry might arise whether the signal might have been reduced by the clipping algorithm of the data reduction pipeline. However, since the source is at the phase centre, its potential signal in one channel would be constant over all times and hence either be clipped in the majority of all visibilities (reducing in a missing channel) or just in the minority of all channels (resulting in the preservation of the signal). In the frequency ranges 769–935 MHz and 961–1051 MHz as analyzed in this work no signal was totally clipped, which suggests that the nondetection of the signal is not due to data processing. A possible test by

ingesting an artificial source before data processing was not done because the pipeline was not directly available to the authors.

A Bayesian approach is employed to obtain the constraints on axion parameters. The likelihood function for axions with mass  $m_a$  and coupling  $g_{a\gamma\gamma}$  is given by

$$\mathcal{L}(m_a, g_{a\gamma\gamma}) = \prod_{i=1}^{N_{\text{ch}}} \frac{1}{\sqrt{2\pi}\sigma_{\nu_i}} \exp \left[ -\frac{(d_{\nu_i} - \mu_{\nu_i} - \bar{S}_{\nu_i}(m_a, g_{a\gamma\gamma}))^2}{2\sigma_{\nu_i}^2} \right], \quad (4)$$

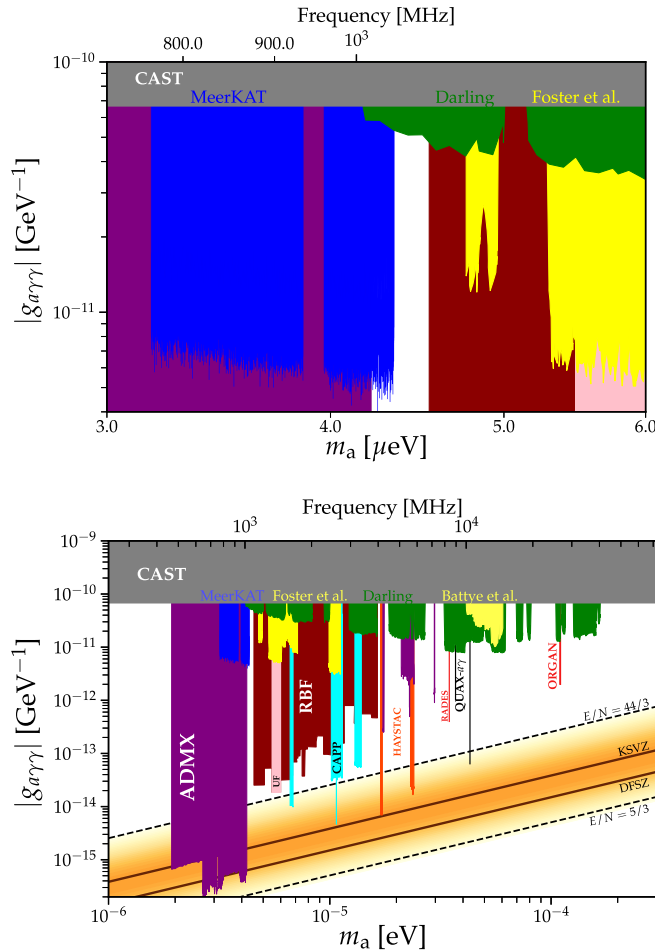


FIG. 3. The 95% CL upper limits on  $g_{a\gamma\gamma}$  derived from 10-hours of MeerKAT observations of NS J0806.4 – 4123 (blue), compared with other results. The top panel is a cutout of the MeerKAT-relevant regime from the bottom panel ( $3 \mu\text{eV} < m_a < 6 \mu\text{eV}$ ). Other constraints including those from the same observational technique for various targets (green, yellow) [24–27], as well as laboratory experiments [30,31,44–52]. The orange shaded region in the lower panel represents the parameter space of typical QCD axion models, such as the well-known DFSZ/KSVZ models, where the mass and couplings are inversely correlated. Figure was produced using the AxionLimits code [53].

where the product runs over all channels. We substitute the theoretical  $\bar{S}_{\nu_i}$  value calculated via Eq. (2) into Eq. (4) by inserting  $r_{\text{NS}} = 10 \text{ km}$ ,  $M_{\text{NS}} = M_{\odot}$ ,  $P = 11.4 \text{ s}$ ,  $d = 250 \text{ pc}$ ,  $B_0 = 2.5 \times 10^{13} \text{ G}$  [33,42] and adopting  $\rho_{\text{DM}}^{\infty} = 0.45 \text{ GeV cm}^{-3}$  [34,35] and  $v_0 = 200 \text{ km s}^{-1}$ . The velocity dispersion is close to the canonical value usually adopted in direct detection experiments (e.g., [43]). For other choices of parameters, the results can be easily scaled according to Eq. (3). We then sample  $g_{a\gamma\gamma}$  in the range  $[0, 2 \times 10^{-10}] \text{ GeV}^{-1}$  and marginalize over  $\theta, \theta_m$  in the ranges  $[0, \pi]$ ,  $[0, \pi/2]$ , respectively, to obtain the posterior distribution of  $g_{a\gamma\gamma}$  with a flat prior in its sampling range.

Figure 2 shows the posterior probability distributions of  $g_{a\gamma\gamma}$  for five selected axion masses. One can see that because of differing SNRs, the distribution functions are centered at different values of  $g_{a\gamma\gamma}$ . However, even the distribution function with the highest significance ( $m_a = 3.212 \mu\text{eV}$ ) does not favour nonzero  $g_{a\gamma\gamma}$  at more than  $4\sigma$ . We therefore regard all of the distributions as providing upper limits of  $g_{a\gamma\gamma}$ . Our final 95% CL limits of  $g_{a\gamma\gamma}$  along with other observational and experimental constraints are shown in Fig. 3.

## V. CONCLUSION

From 10 hours of MeerKAT observations, we have explored the conversion of DM axions into radio-frequency photons in the magnetosphere of the isolated neutron star J0806.4 – 4123, constraining the axion/photon coupling to be  $g_{a\gamma\gamma} \lesssim 9.3 \times 10^{-12} \text{ GeV}^{-1}$  at 95% CL in the mass range of  $3.18 - 4.35 \mu\text{eV}$ . This result provides the strongest known constraints in the mass range of  $4.20 - 4.35 \mu\text{eV}$ , which is not covered by ADMX [30]. Compared with the other constraints derived from the NS magnetospheric axion/photon conversion, we have extended the studied window of axion to lower masses thanks to the UHF receiver of MeerKAT and have reached better performance at overlapping frequencies due to the high sensitivity of MeerKAT in this range.

These constraints can be improved via future observations in several ways. Given the time-dependent nature of the anticipated signal induced by NS rotation, an increased sensitivity can in principle be achieved by exploiting time-series data, as discussed in [12]. More observations of other highly magnetized sources can also lead to tighter constraints. In the future, the analysis procedure, especially the possible confusion of signal with RFI may also be investigated through injecting synthetic signals either in the experiment or the simulation. On the theoretical side, there are also a number of aspects, recently explored in literature [15–19], which can be incorporated to improve the accuracy of our constraints. These include corrections due to gravitational and relativistic effects, ray tracing of the emitted photons and a more sophisticated modeling of magnetospheric plasma effects.

## ACKNOWLEDGMENTS

The MeerKAT telescope is operated by the South African Radio Astronomy Observatory, which is a facility of the National Research Foundation, an agency of the Department of Science and Innovation. Apart from the raw data provided by SARA0, this study relies on the processed data products as provided by the Science Data Processor pipeline of MeerKAT. We also acknowledge use of the NumPy and SciPy scientific computing packages [54,55]. N.H. is supported by the National Natural Science Foundation of China (NSFC) under Grant No. 12150410317. H. C. is supported by the South African Department of Science and Innovation, the National Research Foundation through a SARChI's South African SKA Fellowship within the SARA0 Research Chair held by RC Kraan-Korteweg, and the Key Research Project of Zhejiang Lab (No. 2021 PE0AC03). Y.-Z. M. is supported by the National Research Foundation of South Africa under Grants No. 120385 and No. 120378, NITheCS program “New Insights into Astrophysics and Cosmology with

Theoretical Models confronting Observational Data”, and National Natural Science Foundation of China with Project No. 12047503. Q. Y. is supported by the Key Research Program of Chinese Academy of Sciences (No. XDPB15) and the Program for Innovative Talents and Entrepreneur in Jiangsu. F. D. is supported by the National Natural Science Foundation of China (NSFC) under Grant No. 11873094. Y. C. acknowledges the support from the NSFC under Grant No. 12050410259, and Center for Astronomical Mega-Science, Chinese Academy of Sciences, for the FAST distinguished young researcher fellowship (19-FAST-02), and MOST for the Grant No. QNJ2021061003L. R. D. is supported in part by the National Key R&D Programme of China (2021YFC2203100). We acknowledge the use of China SKA Regional Center prototype system [56] at Shanghai Astronomical Observatory, funded by the National Key R&D Programme of China (No. 2018YFA0404603) and Chinese Academy of Sciences (No. 114231KYSB20170003).

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- [1] R. D. Peccei and H. R. Quinn, *Phys. Rev. D* **16**, 1791 (1977).  
 [2] S. Weinberg, *Phys. Rev. Lett.* **40**, 223 (1978).  
 [3] F. Wilczek, *Phys. Rev. Lett.* **40**, 279 (1978).  
 [4] C. Vafa and E. Witten, *Phys. Rev. Lett.* **53**, 535 (1984).  
 [5] P. Svrcek and E. Witten, *J. High Energy Phys.* **06** (2006) 051.  
 [6] A. Arvanitaki, S. Dimopoulos, S. Dubovsky, N. Kaloper, and J. March-Russell, *Phys. Rev. D* **81**, 123530 (2010).  
 [7] M. Dine and W. Fischler, *Phys. Lett.* **120B**, 137 (1983).  
 [8] J. Preskill, M. B. Wise, and F. Wilczek, *Phys. Lett.* **120B**, 127 (1983).  
 [9] D. J. E. Marsh, *Phys. Rep.* **643**, 1 (2016).  
 [10] I. G. Irastorza and J. Redondo, *Prog. Part. Nucl. Phys.* **102**, 89 (2018).  
 [11] M. S. Pshirkov and S. B. Popov, *J. Exp. Theor. Phys.* **108**, 384 (2009).  
 [12] A. Hook, Y. Kahn, B. R. Safdi, and Z. Sun, *Phys. Rev. Lett.* **121**, 241102 (2018).  
 [13] F. P. Huang, K. Kadota, T. Sekiguchi, and H. Tashiro, *Phys. Rev. D* **97**, 123001 (2018).  
 [14] B. R. Safdi, Z. Sun, and A. Y. Chen, *Phys. Rev. D* **99**, 123021 (2019).  
 [15] R. A. Battye, B. Garbrecht, J. I. McDonald, F. Pace, and S. Srinivasan, *Phys. Rev. D* **102**, 023504 (2020).  
 [16] M. Leroy, M. Chianese, T. D. P. Edwards, and C. Weniger, *Phys. Rev. D* **101**, 123003 (2020).  
 [17] S. J. Witte, D. Noordhuis, T. D. P. Edwards, and C. Weniger, *Phys. Rev. D* **104**, 103030 (2021).  
 [18] R. A. Battye, B. Garbrecht, J. I. McDonald, and S. Srinivasan, *J. High Energy Phys.* **09** (2021) 105.  
 [19] A. J. Millar, S. Baum, M. Lawson, and M. C. D. Marsh, *J. Cosmol. Astropart. Phys.* **11** (2021) 013.  
 [20] L. F. Abbott and P. Sikivie, *Phys. Lett.* **120B**, 133 (1983).  
 [21] S. Colafrancesco, M. Regis, P. Marchegiani, G. Beck, R. Beck, H. Zechlin, A. Lobanov, and D. Horns, *Proc. Sci. AASKA14* (2015) 100 [arXiv:1502.03738].  
 [22] R. Nan, D. Li, C. Jin, Q. Wang, L. Zhu, W. Zhu, H. Zhang, Y. Yue, and L. Qian, *Int. J. Mod. Phys. D* **20**, 989 (2011).  
 [23] N. Du, N. Force, R. Khatriwada, E. Lentz, R. Ottens, L. J. Rosenberg, G. Rybka, G. Carosi, N. Woollett, D. Bowring *et al.*, *Phys. Rev. Lett.* **120**, 151301 (2018).  
 [24] J. Darling, *Phys. Rev. Lett.* **125**, 121103 (2020).  
 [25] J. Darling, *Astrophys. J. Lett.* **900**, L28 (2020).  
 [26] R. A. Battye, J. Darling, J. McDonald, and S. Srinivasan, *Phys. Rev. D* **105**, L021305 (2022).  
 [27] J. W. Foster, Y. Kahn, O. Macias, Z. Sun, R. P. Eatough, V. I. Kondratiev, W. M. Peters, C. Weniger, and B. R. Safdi, *Phys. Rev. Lett.* **125**, 171301 (2020).  
 [28] J. W. Foster, S. J. Witte, M. Lawson, T. Linden, V. Gajjar, C. Weniger, and B. R. Safdi, arXiv:2202.08274.  
 [29] J. Jonas and MeerKAT Team, *Proc. Sci. MeerKAT2016* (2015) 001, <https://pos.sissa.it/277/001/pdf>.  
 [30] C. Bartram *et al.* (DMX Collaboration), *Phys. Rev. Lett.* **127**, 261803 (2021).  
 [31] S. DePanfilis, A. C. Melissinos, B. E. Moskowitz, J. T. Rogers, Y. K. Semertzidis, W. U. Wuensch, H. J. Halama, A. G. Prodel, W. B. Fowler, and F. A. Nezrick, *Phys. Rev. Lett.* **59**, 839 (1987).  
 [32] P. Goldreich and W. H. Julian, *Astrophys. J.* **157**, 869 (1969).  
 [33] B. Posselt, S. B. Popov, F. Haberl, J. Truemper, R. Turolla, and R. Neuhaeuser, *Astrophys. Space Sci.* **308**, 171 (2007).  
 [34] J. Bovy and S. Tremaine, *Astrophys. J.* **756**, 89 (2012).  
 [35] J. I. Read, *J. Phys. G* **41**, 063101 (2014).

- [36] K. Bondarenko, A. Boyarsky, J. Pradler, and A. Sokolenko, [arXiv:2203.08663](https://arxiv.org/abs/2203.08663).
- [37] W. D. Cotton, *Publ. Astron. Soc. Pac.* **120**, 439 (2008).
- [38] D. S. Briggs, Ph.D. thesis, New Mexico Institute of Mining and Technology, United States (1995).
- [39] Planck Collaboration, P. A. R. Ade, N. Aghanim, M. Arnaud, M. Ashdown, E. Aubourg, J. Aumont, C. Baccigalupi, A. J. Banday, R. B. Barreiro *et al.*, *Astron. Astrophys.* **586**, A140 (2016).
- [40] Y.-C. Li, Y.-Z. Ma, M. Remazeilles, and K. Moodley, *Phys. Rev. D* **97**, 023514 (2018).
- [41] J. M. Wrobel and R. C. Walker, in *Synthesis Imaging in Radio Astronomy II*, edited by G. B. Taylor, C. L. Carilli, and R. A. Perley, Astronomical Society of the Pacific Conference Series Vol. 180 (San Francisco, 1999), p. 171.
- [42] D. L. Kaplan and M. H. van Kerkwijk, *Astrophys. J.* **705**, 798 (2009).
- [43] E. Aprile *et al.* (XENON Collaboration), *Phys. Rev. Lett.* **119**, 181301 (2017).
- [44] V. Anastassopoulos *et al.* (CAST Collaboration), *Nat. Phys.* **13**, 584 (2017).
- [45] T. Braine, R. Cervantes, N. Crisosto, N. Du, S. Kimes, L. J. Rosenberg, G. Rybka, J. Yang, D. Bowring, A. S. Chou *et al.*, *Phys. Rev. Lett.* **124**, 101303 (2020).
- [46] C. Bartram *et al.*, [arXiv:2110.10262](https://arxiv.org/abs/2110.10262).
- [47] C. Hagmann, P. Sikivie, N. S. Sullivan, and D. B. Tanner, *Phys. Rev. D* **42**, 1297 (1990).
- [48] O. Kwon, D. Lee, W. Chung, D. Ahn, H. Byun, F. Caspers, H. Choi, J. Choi, Y. Chung, H. Jeong *et al.*, *Phys. Rev. Lett.* **126**, 191802 (2021).
- [49] K. M. Backes *et al.* (HAYSTAC Collaboration), *Nature (London)* **590**, 238 (2021).
- [50] A. A. Melcón *et al.* (CAST Collaboration), *J. High Energy Phys.* **10** (2020) 075.
- [51] D. Alesini, C. Braggio, G. Carugno, N. Crescini, D. D'Agostino, D. Di Gioacchino, R. Di Vora, P. Falferi, U. Gambardella, C. Gatti *et al.*, *Phys. Rev. D* **103**, 102004 (2021).
- [52] B. T. McAllister, G. Flower, E. N. Ivanov, M. Goryachev, J. Bourhill, and M. E. Tobar, *Phys. Dark Universe* **18**, 67 (2017).
- [53] C. O'Hare (2020), [10.5281/zenodo.3932430](https://zenodo.org/record/3932430).
- [54] C. R. Harris, K. J. Millman, S. J. van der Walt, R. Gommers, P. Virtanen, D. Cournapeau, E. Wieser, J. Taylor, S. Berg, N. J. Smith *et al.*, *Nature (London)* **585**, 357 (2020).
- [55] P. Virtanen, R. Gommers, T. E. Oliphant, M. Haberland, T. Reddy, D. Cournapeau, E. Burovski, P. Peterson, W. Weckesser, J. Bright *et al.*, *Nat. Methods* **17**, 261 (2020).
- [56] T. An, X. Wu, B. Lao, S. Guo, Z. Xu, W. Lv, Y. Zhang, and Z. Zhang, *Sci. China Phys. Mech. Astron.* **65**, 129501 (2020), <https://www.sciengine.com/SCPMA/doi/10.1007/s11433-022-1981-8;JSESSIONID=add73d1a-822a-4bec-9eab-4a0ac42e69fa>.